

Synergy between asteroseismology and exoplanet science: an outlook

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Abstract

Space-based asteroseismology has been playing an important role in the characterization of exoplanet-host stars and their planetary systems. The future looks even brighter, with space missions such as NASA’s TESS and ESA’s PLATO ready to take on this legacy. In this contribution, we provide an outlook on the synergy between asteroseismology and exoplanet science, namely, on the prospect of conducting a populational study of giant planets around oscillating evolved stars with the TESS mission.

1 Introduction

The asteroseismology revolution initiated by *Kepler* (Borucki *et al.*, 2010) is set to continue over the coming decades with the launches of TESS (Ricker *et al.*, 2015), PLATO (Rauer *et al.*, 2014), as well as WFIRST (Spergel *et al.*, 2013), with these missions expected to raise the number of solar-like oscillators to a few million stars (Huber, 2018). Note that over 90% of all detections are expected to be for evolved stars, with PLATO by far contributing the most detections for dwarfs and subgiants ($\sim 80,000$). If we combine this with dedicated ground-based efforts, such as the SONG network (Grundahl *et al.*, 2017) of 1-meter telescopes, we are then positive that the synergy between asteroseismology and exoplanet science can only continue to grow (Campante *et al.*, 2018).

Synergetic studies of evolved stars are made possible by even moderate photometric cadences, which can be used to simultaneously detect transits and stellar oscillations. A very exciting prospect is that of conducting asteroseismology of red-giant hosts using the 30-minute cadence of TESS full-frame images (FFIs). Based on an all-sky stellar and planetary synthetic population (Sullivan *et al.*, 2015), we predict that solar-like oscillations will be detectable in up to 200 low-luminosity red-giant branch (LLRGB) stars hosting close-in giant planets (Campante *et al.*, 2016).

The population of transiting planets around evolved stars is so far largely unexplored (Huber, 2018). And although radial-velocity surveys are mostly complete for planets near or above $1 M_{\text{Jup}}$ at > 0.2 AU, there is a dearth of planets with orbital periods $P < 80$ d. Nonetheless, *Kepler/K2* have discovered several close-in giant planets around LLRGB stars (e.g., Grunblatt *et al.*, 2016, 2017), hinting at a population of warm sub-Jovian planets around evolved stars that would be accessible to TESS. *Kepler/K2* mainly targeted main-sequence stars, and observed too few LLRGB stars to

detect enough planets for robust statistics. TESS will increase the number of LLRGB stars with space-based photometry by one order of magnitude over *Kepler/K2*, providing an unprecedented opportunity to address a number of key questions in exoplanet science, namely:

- The role of stellar flux on hot-Jupiter inflation;
- Giant-planet occurrence as a function of stellar mass and evolution;
- Correlation between metallicity and giant-planet occurrence around evolved stars.

In this work, we start by characterizing the parent population of LLRGB stars to be searched for transits based on a GALAXIA (Sharma *et al.*, 2011) simulation (Sect. 2). We focus on the southern ecliptic hemisphere, which will be surveyed during year 1 of TESS’s primary mission. We next implement a software tool for planetary-transit search based on a Box-fitting Least Squares (BLS) algorithm (Sect. 3). The tool is tested both for statistical false positive rates and detection sensitivity using artificial TESS light curves. The former involves running the code on a sufficiently large number of light curves containing only instrumental/shot noise and stellar (correlated) signals, namely, granulation and oscillations. The latter involves running the code on the same light curves, although now with injected transits. We conclude with a few considerations regarding the use of TESS photometry alone in candidate vetting (Sect. 4).

2 Parent stellar population

2.1 Synthetic population

We start with an all-sky, magnitude-limited synthetic stellar population generated with GALAXIA. Output absolute

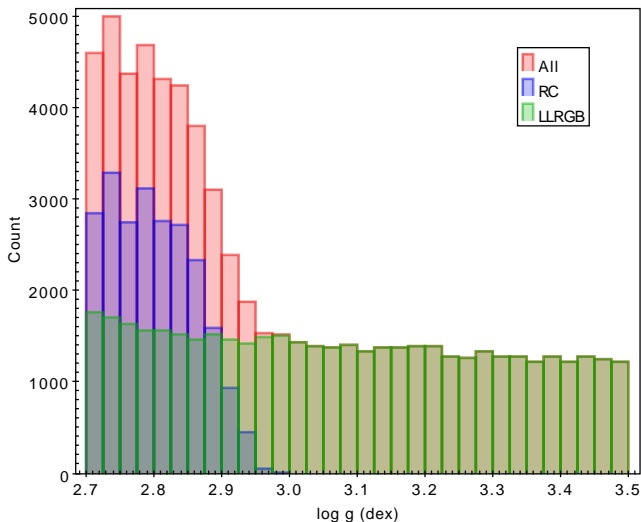


Figure 1: Contamination of RC stars in the initial sample.

magnitudes were converted to apparent magnitudes and extinction applied. We ended up only retaining stars down to magnitude 13 in the Johnson–Cousins I_C band. Although somewhat optimistic, this magnitude cut is used to ensure that all detectable oscillating LLRGB stars are captured (Campante *et al.*, 2016). Note that the simulation is undersampled by a factor of 10 to ease up on the data handling.

We made an initial selection of *putative* LLRGB stars from this synthetic population by applying the following cuts on $\log g$ and T_{eff} : $2.7 < \log g < 3.5$ and $T_{\text{eff}} < 5500$ K. The upper $\log g$ cut ensures that stars are evolved enough to oscillate with frequencies detectable with 30-minute-cadence data (Chaplin *et al.*, 2014). The lower $\log g$ cut is purely empirical (Hekker *et al.*, 2011) and leads to contamination of the sample by red clump (RC) stars (on which more below). This step was followed by selecting only those stars located in the southern ecliptic hemisphere. Finally, we determined the median number of sectors (by considering a number of different initial pointings) over which each star would be observed with TESS using the PYTHON package TVGUIDE¹ (each sector corresponds to a 27.4-day coverage), having discarded stars that fall off silicon. The final tally amounts to $\sim 6.8 \times 10^5$ stars (after applying the factor-of-10 correction).

We now come back to the issue of the sample contamination by RC stars. Indeed, 34 % of stars in the above sample are RC stars (see Fig. 1). We apply an additional cut on the asteroseismic observable $\Delta\nu$, namely, $\Delta\nu > 10 \mu\text{Hz}$, in order to mitigate this contamination effect. This effectively leads to the removal of all RC stars and is the equivalent to raising the lower $\log g$ cut to $\log g \gtrsim 2.9$. We take the resulting sample of bona fide LLRGB stars as our final, parent (synthetic) stellar population, which comprises $\sim 3.0 \times 10^5$ stars (after applying the factor-of-10 correction).

2.2 Population characteristics

We now look at the main characteristics of this parent (synthetic) stellar population. Figure 2 shows the overall stellar radius, mass and metallicity distributions (light red). Distributions are also shown for a subset of stars lying farther away from the Galactic plane, i.e., $|b| > 10^\circ$ (light blue; comprising 68 % of the parent population). The Galactic latitude of a target strongly influences the likelihood of it giving rise

to an astrophysical false positive. For $|b| < 10^\circ$, the density of background stars is very high, meaning that any observed eclipse is more likely to be from a background eclipsing binary, whereas for $|b| > 20^\circ$ planets should represent a majority over false positives² (Sullivan *et al.*, 2015).

Panel (a) of Fig. 3 shows the V -band magnitude distribution of the parent population (peaking at $V \sim 13$ –14), while panel (b) shows a luminosity-color diagram that may be used to inform target selection. We will likely need to apply a stricter magnitude cut depending on the actual oscillations detectability limit. We also assessed the fraction of stars in the parent population observed over a median of 1 TESS sector (78 %) and 1 or 2 TESS sectors (93 %). Furthermore, we note that for 26 % of the stars ν_{max} is above the Nyquist frequency.

3 Automated transit detection: performance assessment

3.1 Artificial light curves

We generated artificial light curves for the $\sim 30,000$ unique, bona fide LLRGB stars in our parent (synthetic) stellar population (see Sect. 2). Generation of the light curves is performed originally in the frequency domain, after which an inverse Fourier transform is applied (Kuszlewicz *et al.*, submitted). We consider only the 30-minute cadence of TESS FFIs and apply a window function to account for the data downlink occurring every spacecraft orbit.

We used a photometric noise model for TESS (Sullivan *et al.*, 2015; Campante *et al.*, 2016) to predict the rms noise per exposure time. A systematic term of 20 ppm $\text{hr}^{1/2}$ was included in this calculation. To model the granulation power spectral density, we adopted a scaled version (to predict TESS granulation amplitudes) of model F of Kallinger *et al.* (2014), which contains two Harvey-like components. No aliased granulation power was considered. Individual radial, (mixed) dipole and quadrupole modes were also modeled whenever $\nu_{\text{max}} < \nu_{\text{Nyq}}$. Panels (a) and (b) of Fig. 4 respectively display the power spectral density (PSD) and corresponding light curve of a $V = 10.3$ star observed for 27.4 days (or 1 TESS sector). The oscillation bump can be seen around $\nu_{\text{max}} \approx 167 \mu\text{Hz}$.

We generated model transit light curves using the PYTHON package BATMAN³. Assuming circular planetary orbits, we seeded one planet per star. We next drew orbital periods and planet radii from uniform distributions spanning the parameter space of interest (0.5 to 27.4 days and 4 to 22 R_\oplus , respectively). Orbital periods were redrawn until no systems were left within the Roche limit or the stellar envelope. We assume that all planets transit and draw the impact parameter from a uniform distribution defined over the half-open interval $[0, 1[$. Input to BATMAN includes the time of inferior conjunction, orbital period, planet radius, semi-major axis, and orbital inclination. A quadratic limb darkening law is used and its coefficients set to fixed values (see Barclay *et al.*, 2015). We further account for the long integration time by supersampling the model 11 times per cadence then integrating over these subsamples.

²Note that these remarks were made with reference to planet detections around TESS target stars.

³<https://www.cfa.harvard.edu/~lkreidberg/batman/>

¹<https://heasarc.gsfc.nasa.gov/docs/tess/proposal-tools.html#tvguide>

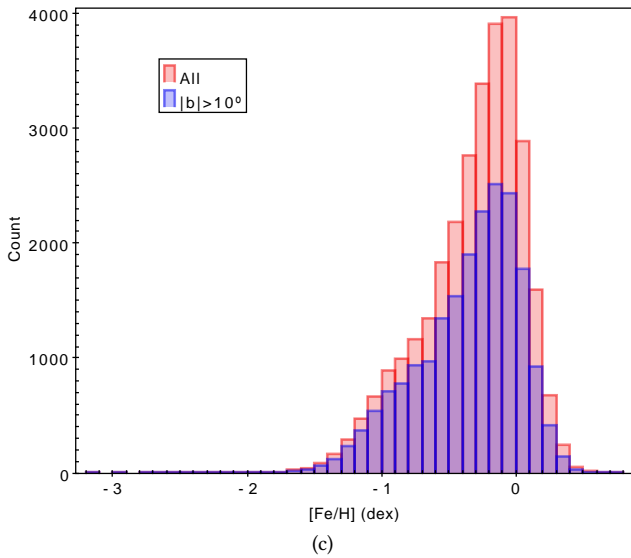
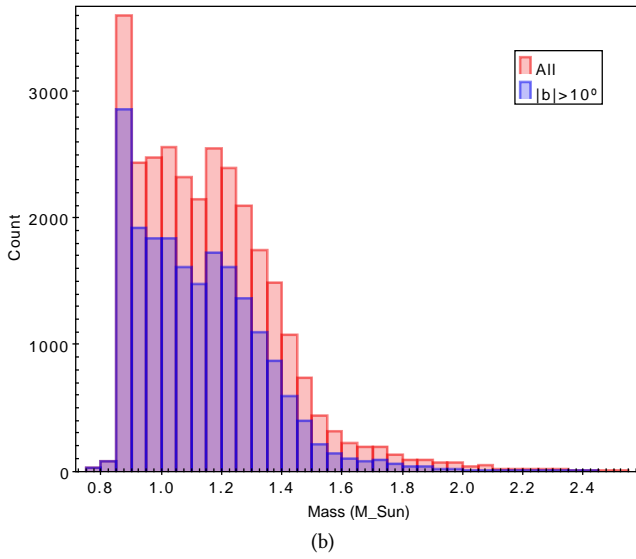
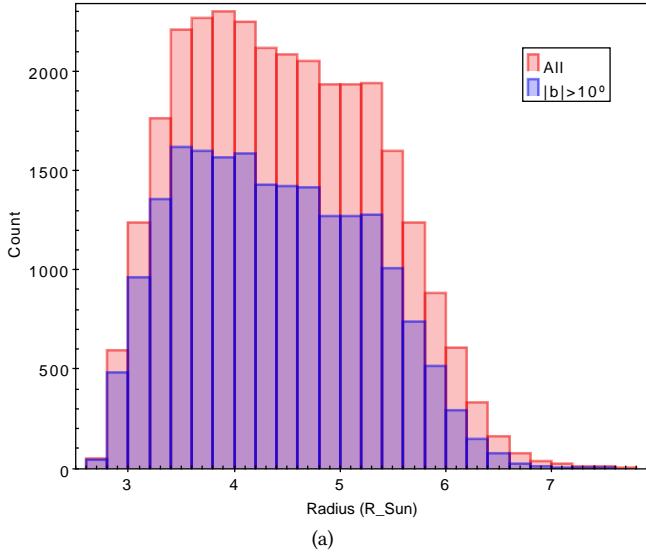


Figure 2: Stellar (a) radius, (b) mass and (c) metallicity distributions of the parent stellar population.

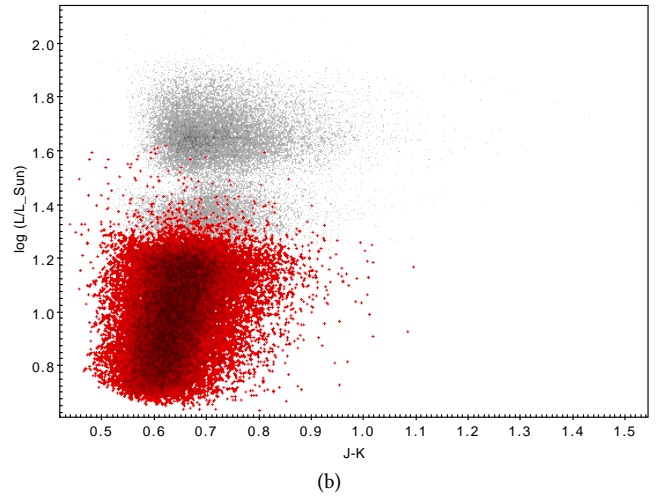
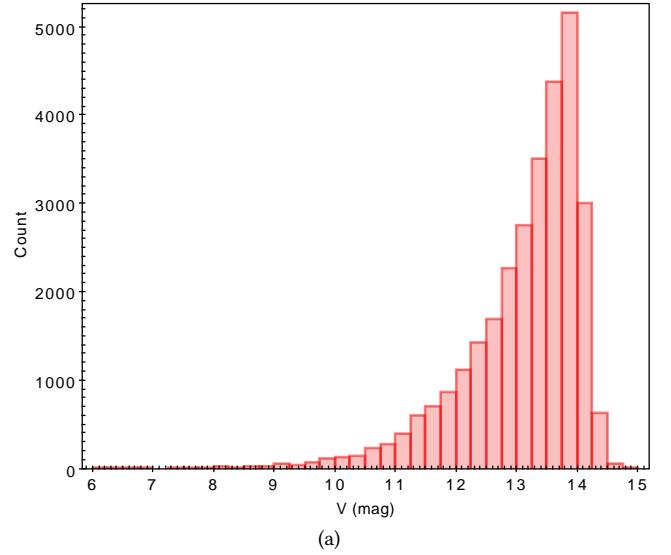


Figure 3: V -band magnitude distribution (a) and luminosity-color diagram (b) of the parent stellar population. The luminosity-color diagram of the parent stellar population (red) is superimposed on that of the initial sample (gray; see Sect. 2.1).

3.2 BLS algorithm

We search for transits using an updated version of the pipeline presented in Barros *et al.* (2016), which makes use of a PYTHON implementation⁴ of the BLS algorithm originally introduced by Kovács *et al.* (2002). The search is made over periods ranging from 1 day to 70 % of the light curve duration and over fractional transit durations ranging from 0.001 to 0.3 with $nb = 200$ phase bins. Frequency sampling is optimized according to $\delta\nu = 1/(P_{\max} \cdot nb)$, where P_{\max} is the maximum period searched for. Using the periods and epochs found by the BLS algorithm, each light curve is phase-folded and the signal detection efficiency (SDE ; Kovács *et al.*, 2002) computed.

The pipeline searches $npass$ transits per light curve and sorts them according to the SDE . Results for all the light curves are also sorted according to the maximum SDE reported for each light curve. It also tests for the following features: possibility of a secondary transit/eclipse, sinusoidal behavior, and mono transit (or an effective number of transits that is less than the total number of transits). Provided

⁴<https://github.com/dfm/python-bls>

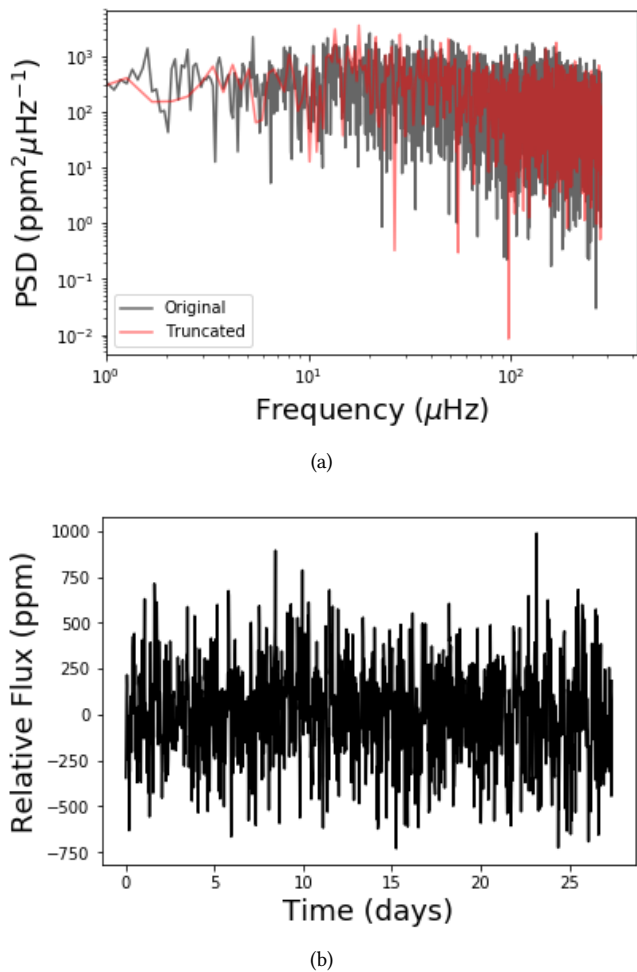


Figure 4: Power spectral density (a) and corresponding light curve (b) of a $V = 10.3$ star observed for 27.4 days (or 1 TESS sector). The star has $\nu_{\max} \approx 167 \mu\text{Hz}$. The resolution of the truncated spectrum (red) is defined by the duration of the light curve, whereas the original spectrum (black) is oversampled. No window function has yet been applied to the light curve to account for the data downlink.

the fitted depth is positive, the pipeline produces a series of plots for each candidate. Figure 5 shows the pipeline output for the same artificial star considered in Fig. 4 (after transit injection). The two injected transits are correctly recovered. Another example can be found in Fig. 6, where the pipeline output is shown for a $V = 13.1$ star observed for 54.8 days (or 2 TESS sectors). All three injected transits are correctly recovered.

3.3 Statistical false positives

We began by running the pipeline on the $\sim 30,000$ generated artificial light curves (prior to transit injection) in order to assess the rate of statistical false positives. Results are shown in Fig. 7. From panel (a), we find that an SDE threshold of 8.88 produces approximately one statistical false positive over the $\sim 30,000$ light curves or a rate of 0.003%. Table 1 provides SDE thresholds as a function of the statistical false positive rate (0.1%, 1% and 5%). Panels (b) and (c) of Fig. 7 emphasize the dependence of the SDE on the sample’s limiting magnitude and light curve duration, respectively. The dependence on the latter is particularly obvious, with shorter light curves giving rise to lower SDE values (cf. Kovács *et al.*, 2002). Therefore, Table 1 also provides SDE thresholds when only 27.4- (1 sector) and 54.8-day-long (2 sectors) light curves

Table 1: SDE threshold as a function of the statistical false positive rate.

Rate	SDE (All)	SDE (1 sector)	SDE (2 sectors)
0.1 %	7.29	6.75	7.22
1 %	6.34	5.98	6.39
5 %	5.61	5.33	5.73

are considered.

3.4 Detection sensitivity

We now run the pipeline on the same $\sim 30,000$ artificial light curves after transit injection in order to assess the detection sensitivity (or survey completeness). Figure 8 shows the relative difference between the injected (P_{in}) and recovered (P_{out}) orbital periods as a function of the SDE . Vertical dashed lines mark (from right to left) the SDE thresholds corresponding to 0.1%, 1% and 5% statistical false positive rates when considering 1 TESS sector.

The pipeline sometimes recovers half (0.5 ordinate) or double (-1 ordinate) the injected period, which tends to happen close to the aforementioned thresholds. The former subset mostly corresponds to the case of two injected transits with an orbital period exceeding 70% of the light curve duration, whereas the latter, less numerous subset mostly corresponds to those injected transits with periods shorter than 1 day. These injected periods are out of bounds with respect to the search parameters of the algorithm. The remaining, rarer cases seem to be genuine statistical misidentifications prompted by the low SDE .

Figure 9 shows the transit detection sensitivity based on 27.4 days of data and a minimum of two transit events, where we have assumed a statistical false positive rate of 5%. The resulting contours demonstrate that even for the most pessimistic case of 27.4 days coverage it will be possible to detect close-in, inflated Jupiters for over 80% of stars, Jupiter-size planets for 70–80% of stars, and large Neptunes only in the most favorable cases.

3.4.1 Out-of-transit flux modulation

We test for the presence of sinusoidal behavior in the light curve by first fitting a sine function to the binned, folded light curve (on the best period). We then perform a linear regression (after linearization) to obtain the coefficient of determination, r^2 . Here, we apply this procedure to an artificial system resembling the confirmed Kepler-91 planetary system, for which ellipsoidal variations have been measured that are caused by a close-in giant planet (Lillo-Box *et al.*, 2014; Barclay *et al.*, 2015). We injected a sine function with period $P_{\text{in}}/2$ and phase 0 at minimum into the light curve (Pfahl *et al.*, 2008), having varied its amplitude. The star is observed for 27.4 days only (or 1 TESS sector). A notional amplitude of 100 ppm of the ellipsoidal modulation (see Fig. 10) leads to $r^2 = 0.22$, above the $r^2 = 0.1$ threshold adopted within the pipeline to flag potential sinusoidal behavior. This amplitude is slightly higher than the approximate upper limit of 75 ppm measured for Kepler-91. We conclude by noting that it is usually necessary to remove instrumental trends and rotational modulation due to spots/plages prior to any transit search. Such detrending of the light curve may (depending on the period) also remove part of the out-of-transit flux modulation, thus making detection of ellipsoidal variations more challenging.

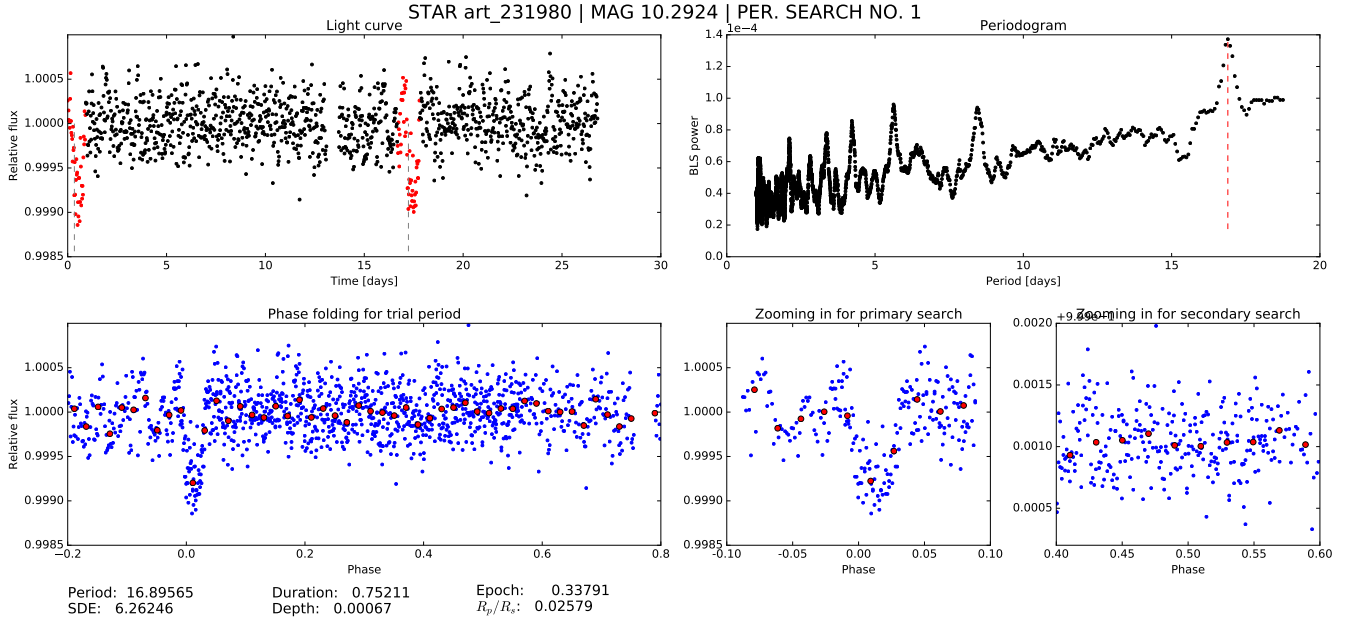


Figure 5: Pipeline output for the same artificial star considered in Fig. 4. The light curve is shown in the top left panel with both (correctly) recovered transits in red. Notice the gap at ≈ 13.7 d due to the data downlink. The BLS periodogram is shown in the top right panel with the vertical dashed line indicating the best period, as determined by the algorithm. The bottom left panel displays the phase-folded light curve using the best period (blue) and a binned version of it (red). The bottom middle and right panels simply zoom in on the phase-folded light curve at the locations of the primary and possible secondary, respectively.

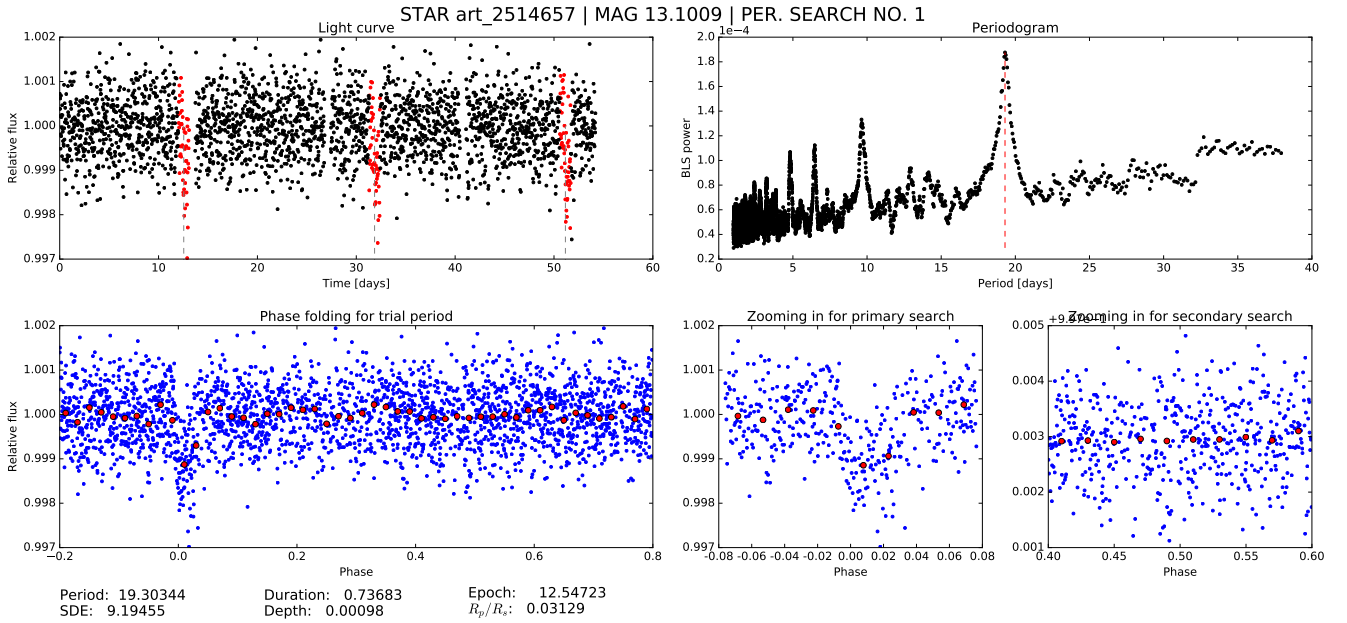
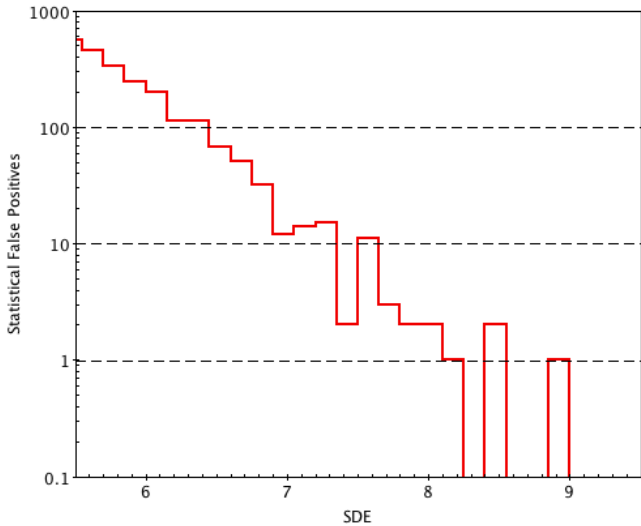


Figure 6: Pipeline output for a $V = 13.1$ star observed for 54.8 days (or 2 TESS sectors). Panels are the same as in Fig. 5.

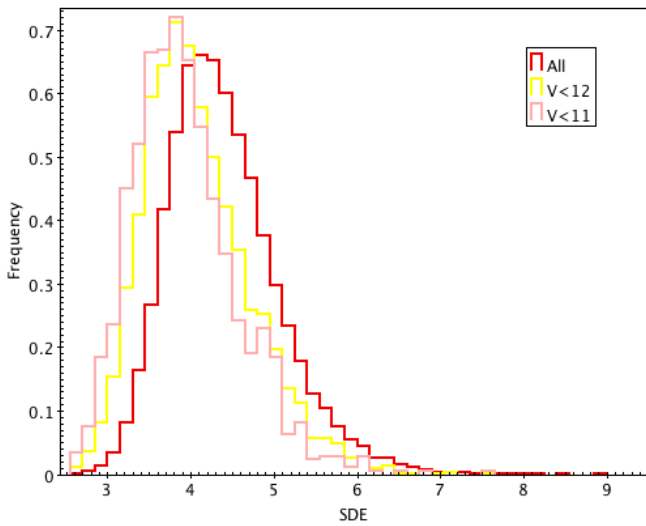
3.4.2 Secondary transit/eclipse

We test for the presence of a secondary transit/eclipse by (i) looking for two closely spaced periods (at most 0.5 days apart) and (ii) assessing whether the corresponding depths differ by at least 10% (relative to the larger of the two depths). The former step also involves checking if a newly detected period matches the 1st overtone of the primary (i.e., $P_{\text{out}}^{\text{primary}}/2$). Figure 11 illustrates this procedure. The light

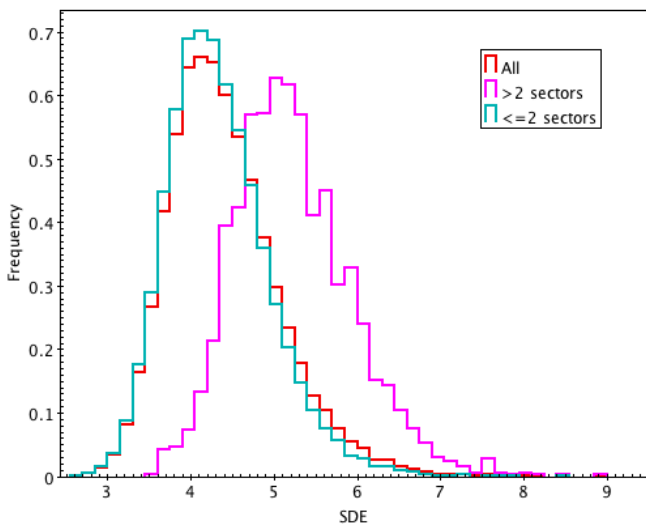
curve has nine injected primary transits. These are easily recovered during the first period search despite the clear presence of a secondary at phase 0.5 (see top panel). A planet-to-star flux ratio of 5×10^{-4} was assumed when simulating the secondary. Such an arbitrarily large planet-to-star flux ratio – as a term of comparison, the occultation depth of Kepler-91b is ~ 50 ppm – is employed here for illustrative purposes only and no attention has been paid to determine whether or



(a)



(b)



(c)

Figure 7: Statistical false positive rates. Panel (a): Number of statistical false positives as a function of the SDE . Horizontal dashed lines indicate 1, 10, and 100 false positives. Panels (b) and (c): Normalized histograms of the SDE depicting its dependence on limiting magnitude and light curve duration, respectively.

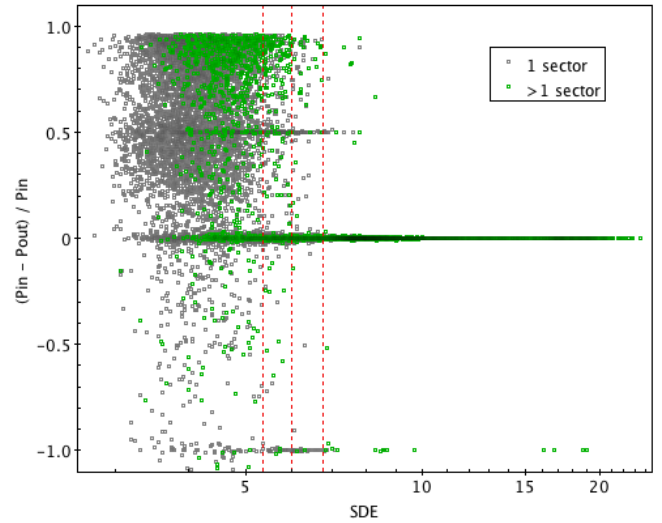


Figure 8: Relative difference between injected (P_{in}) and recovered (P_{out}) orbital periods as a function of the SDE . Only systems with two or more detected transits and measured positive depths have been displayed. Vertical dashed lines mark (from right to left) the SDE thresholds corresponding to 0.1%, 1% and 5% statistical false positive rates when considering 1 TESS sector (see Table 1). Systems observed over 1 (gray) or more (green) TESS sectors are identified.

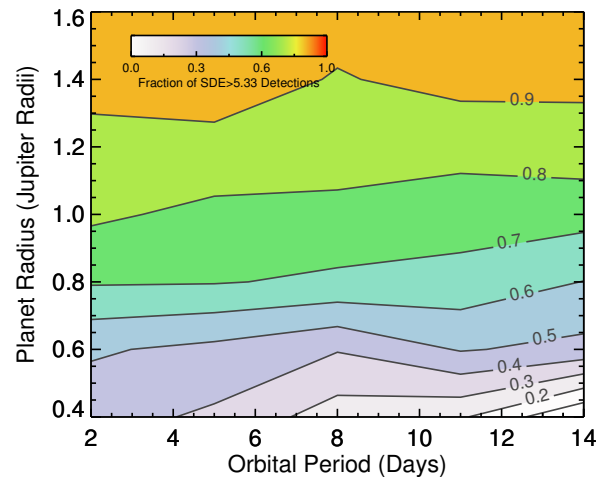
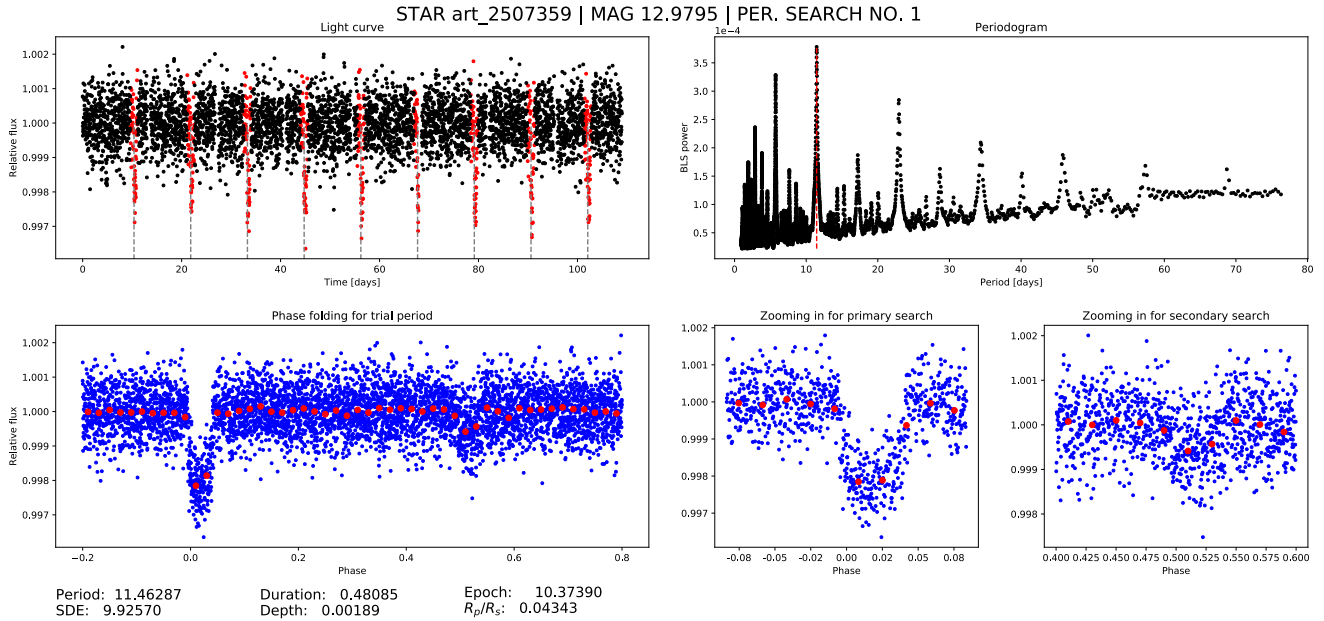


Figure 9: Transit detection sensitivity based on 27.4 days of data and a minimum of two transit events. We assume a statistical false positive rate of 5% (or an SDE threshold of 5.33; see Table 1).

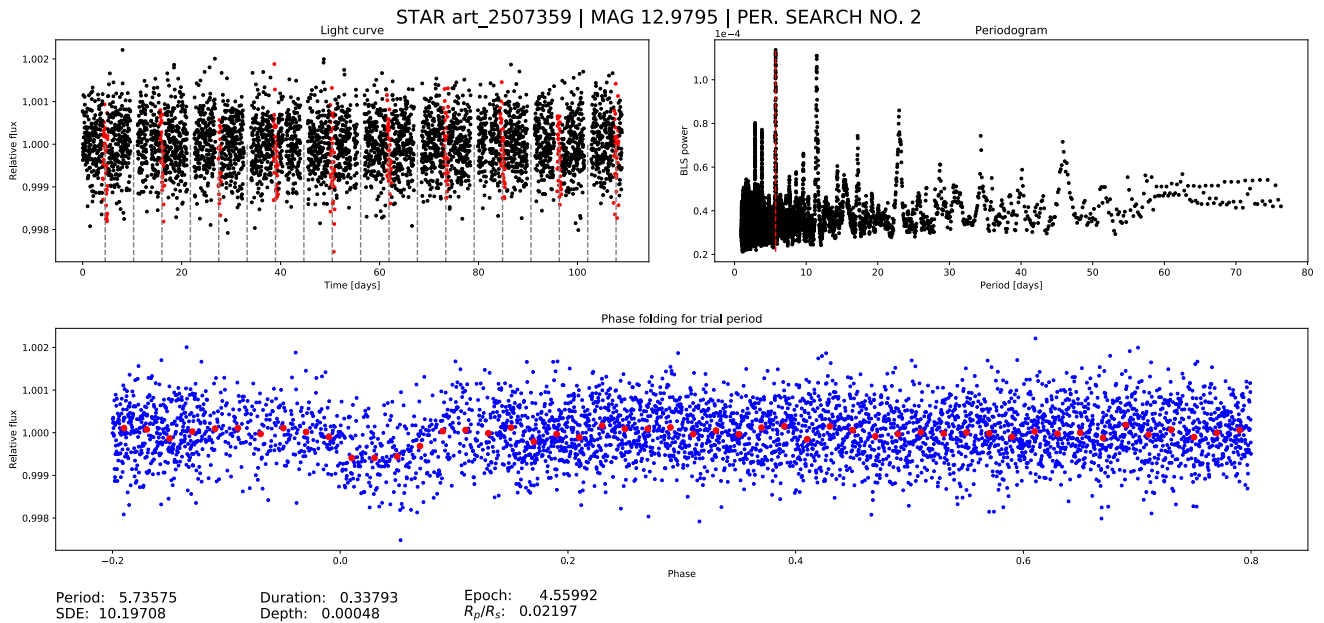
not it is physically sound. When performing the second period search (after masking out all primary transits from the light curve), we notice how the BLS power associated with the true period is brought down relative to that of its 1st overtone (see bottom panel). Detection of the 1st overtone of the primary (after its removal) is a telltale sign of the presence of a secondary transit/eclipse.

4 Candidate vetting

We will be using TESS photometry alone in a first attempt to separate the expected transit-like signals from (the notionally many) astrophysical false positives (due, e.g., to EBs and BEBs) and systematic false alarms. A clear distinction will not always be possible, of course, but this will aid retaining those cases we think are due to genuine transits of short-



(a)



(b)

Figure 11: Pipeline output for a $V = 13.0$ star observed for 109.6 days (or 4 TESS sectors). A primary (at phase 0) and a secondary (at phase 0.5) were injected into the light curve. Panels refer to the first (a) and second (b) period search. Notice that all primary transits have been masked out from the light curve prior to the second period search.

period gas giants around LLRGB stars.

Housekeeping operations such as retaining only those systems with two or more detected transits as well as measured positive depths will be implemented. An SDE threshold will be adopted, its value depending on the number of TESS sectors (cf. Sect. 3.3). Any measured transit depth in excess of 1% will be attributed to an eclipsing binary (a $2 R_J$ planet transiting an LLRGB will at most cause a $\sim 0.5\%$ flux reduction). As illustrated in the previous section, the pipeline tests for a number of additional features, namely, the presence of a

secondary transit/eclipse and sinusoidal behavior. Although not decisive with respect to the vetting procedure, these flags provide useful information that can be used during the detailed transit fitting. We will pay particular attention, at low SDE , to possible period misidentifications (due to an out-of-bounds true period or else statistical in nature) by carefully inspecting the light curve and associated BLS periodogram. We also advocate for the independent fitting – as part of the detailed transit fitting – of the phase-folded odd and even primary transits, since a significant difference in their depths

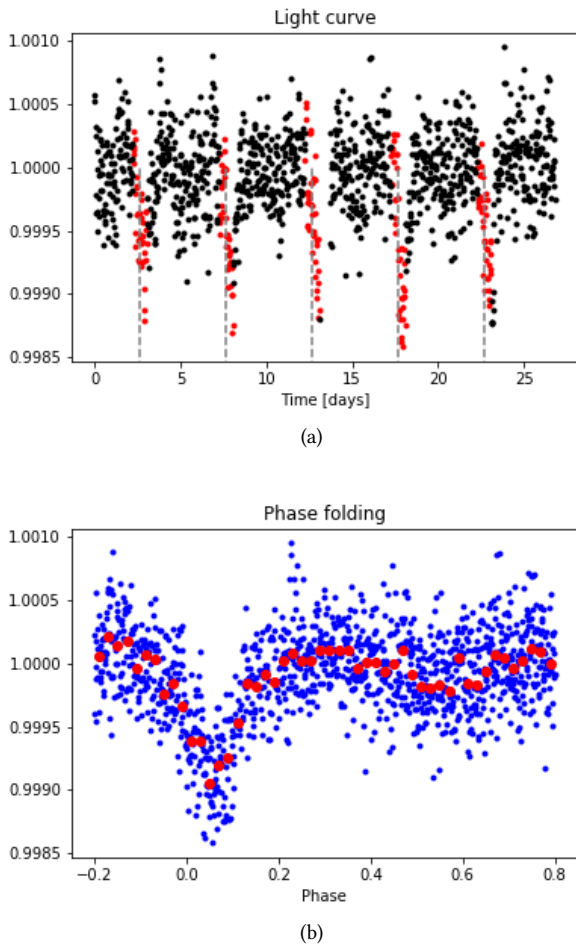


Figure 10: Light curve (a) and phase-folded light curve (b) of an artificial system resembling the Kepler-91 system. The star is observed for 27.4 days (or 1 TESS sector). A sine function with period $P_{\text{in}}/2$ and phase 0 at minimum has been injected into the light curve. Its amplitude was set to 100 ppm. An ellipsoidal modulation is clearly seen in the right panel.

may indicate we are dealing with an eclipsing binary.

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