

Epoch.	Observer.	θ_0	θ_c	$\theta_0 - \theta_c$	ρ_0	ρ_c	$\rho_0 - \rho_c$
1875.20	Dunér	311°42	310°58	+ 0°84	0"52	0"60	- 0"08
1876.89	Dembowski	314°92	315°85	- 0°93	0.522	0.55	- 0.028
1878.90	Doberck	327°71	323°92	+ 3°79	0.47	0.49	- 0.02
1883.50	Engelmann	348°29	349°41	- 1°12	0.311	0.40	- 0.089
1888.142	Schiaparelli	33°60	22°33	+ 11°27	0.35	0.39	- 0.04
1888.171	Tarrant	17°23	22°54	- 5°31	0.43	0.39	+ 0.04
1888.696	"	21°31	26°10	- 4°79	0.40	0.40	0.00
1888.778	"	21°54	26°61	- 5°07	0.46	0.40	+ 0.06
1888.931	Schiaparelli	25°73	27°63	- 1°90	0.36	0.40	- 0.04
1889.018	"	30°25	28°23	+ 2°02	0.35	0.40	- 0.05
1889.021	Young	34°56	28°24	+ 6°32	0.41	0.40	+ 0.01
1889.570	Tarrant	26°97	31°76	- 4°79	0.37	0.41	- 0.04

The recent measures by Professors Schiaparelli and Young and Mr. Tarrant were kindly communicated to me by private letter.

According to the above orbit the distance between the components will gradually increase during the next few years, up to a maximum of about 0''55, and then diminish again as the companion approaches the periastron. The minimum distance will not be reached till the position-angle is nearly 180° (after the periastron passage), when the components will probably be separated by less than 0''2. Careful measures during the next sixteen years will be very valuable for correction of the provisional elements.

On the supposition that the combined mass of the components is equal to the mass of the Sun, the "hypothetical" parallax would be

$$\pi = \alpha P - \frac{2}{3} = 0''05$$

The binary lies a little preceding 62 *Andromedæ*, and its position for 1890.0 is approximately :—

$$\text{R.A. } 2^{\text{h}} 6^{\text{m}} 59^{\text{s}}, \text{ Decl. } +46^{\circ} 58'.4$$

The magnitudes of the components are about 6.7 and 7.6.

A Method of Recording the Transits of Stars by Photography.

By W. E. Wilson.

I wish to bring before the notice of the Society a method by which the transits of stars can be recorded by photography and the personal errors eliminated. If a sensitive photographic plate is placed in the focus of a transit instrument close behind the wires, and the image of a star of suitable magnitude allowed to transit across it, the result is a straight black line on developing the

I used for these experiments the 4-inch finder of my 2-foot reflector. The object-glass was not corrected for photographic rays, and the star trail on the plate was therefore not as fine as it would be with a suitable objective.

Catalogue of Bright Meteors observed at Bristol during the Years 1877 to 1889 inclusive. By William F. Denning.

The following catalogue includes 217 fine meteors, chiefly as bright as *Jupiter* or *Venus*. The radiant points affixed may be generally relied upon, but it is necessary to say that in some cases there is uncertainty, notwithstanding strong presumptive evidence afforded by the flight-directions and appearance of the meteors, that they diverged from the centres assigned.

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