



Study for g-mode oscillations in the Sun using solar neutrino with Super-Kamiokande

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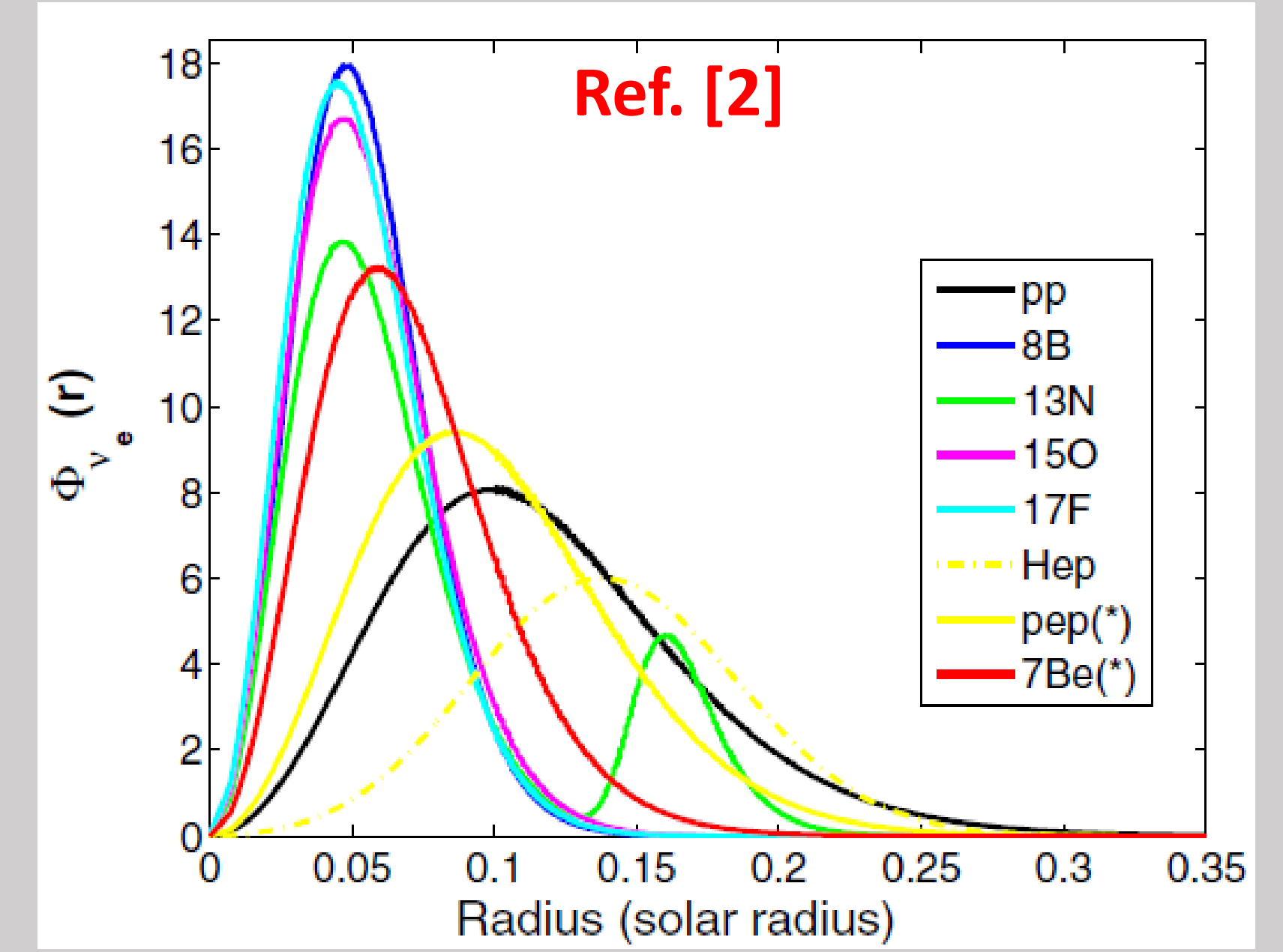
1. Introduction

Standard Solar Model (SSM) [1].

- The SSM has been constructed by inputting several observation results assuming:
 - (1) Spherically symmetric,
 - (2) Hydrostatic equilibrium with no macroscopic motion except for the convection,
 - (3) No mass loss, no mass accretion, no rotation and no magnetic field.
- The model well predicts the production and energy spectrum of solar neutrinos.

Periodic variations in the Sun.

- 11-years periodic change of the sunspot at the surface. This leads to the change of the magnetic field.
 - Some mechanisms propose **conversions of neutrino** ($\nu_L \rightarrow \nu_R$) if neutrino has a large magnetic moment [3-5].
- Solar oscillations around its equilibrium state because of the Sun's restoring force [6].
 - **Acoustic oscillation (p-mode)** due to compressibility [7-9] or **gravity oscillation (g-mode)** due to buoyancy [10].



2. G-mode oscillation and neutrino

Introduction about g-mode oscillations [11, 12].

- G-mode oscillations are described as spherical harmonics function.
- They may **affect both electron density and temperature** in the core of the Sun.
- **These modes have not been detected** because their amplitudes are low at the surface.
- Typical frequencies are about **a few hours (~100-300 μHz)** [11].

⁸B neutrino production rate.

- G-modes are trapped under the convective zone, where ⁸B solar neutrino produced.
- Production rate depends on temperature. It is proportional to T^{24-25} ($T \sim 10^6$ K) [13].
 - Its production rate is **amplified by a factor of 170** [14].
 - It may affect the propagation of solar ν , MSW effect [15, 16], thus survival probability of ν_e [12].

Ref. [11]	n	l	ν (μHz)	Period (hr)	Energy ^a (ergs)	$\delta T/T^a$
1	1	1	267.4	1.04	3.4×10^{31}	6.7×10^{-8}
1	2	2	297.1	0.94	2.1×10^{31}	4.4×10^{-8}
1	3	3	340.2	0.82	4.3×10^{31}	7.7×10^{-8}
2	1	1	193.0	1.44	1.3×10^{33}	4.6×10^{-7}
2	2	2	259.7	1.07	3.7×10^{31}	7.2×10^{-8}
2	3	3	299.2	0.93	4.4×10^{31}	8.8×10^{-8}
3	1	1	153.7	1.81	1.7×10^{33}	6.9×10^{-7}
3	2	2	224.0	1.24	8.7×10^{31}	1.3×10^{-7}
3	3	3	264.7	1.05	3.3×10^{31}	8.1×10^{-8}
4	1	1	128.3	2.17	1.7×10^{33}	7.4×10^{-7}
4	2	2	195.5	1.42	1.6×10^{32}	2.0×10^{-7}
4	3	3	240.5	1.15	2.9×10^{31}	7.4×10^{-8}
5	1	1	109.9	2.53	1.7×10^{33}	8.7×10^{-7}
5	2	2	171.8	1.62	3.1×10^{32}	3.1×10^{-7}
5	3	3	218.9	1.27	6.8×10^{31}	1.3×10^{-7}

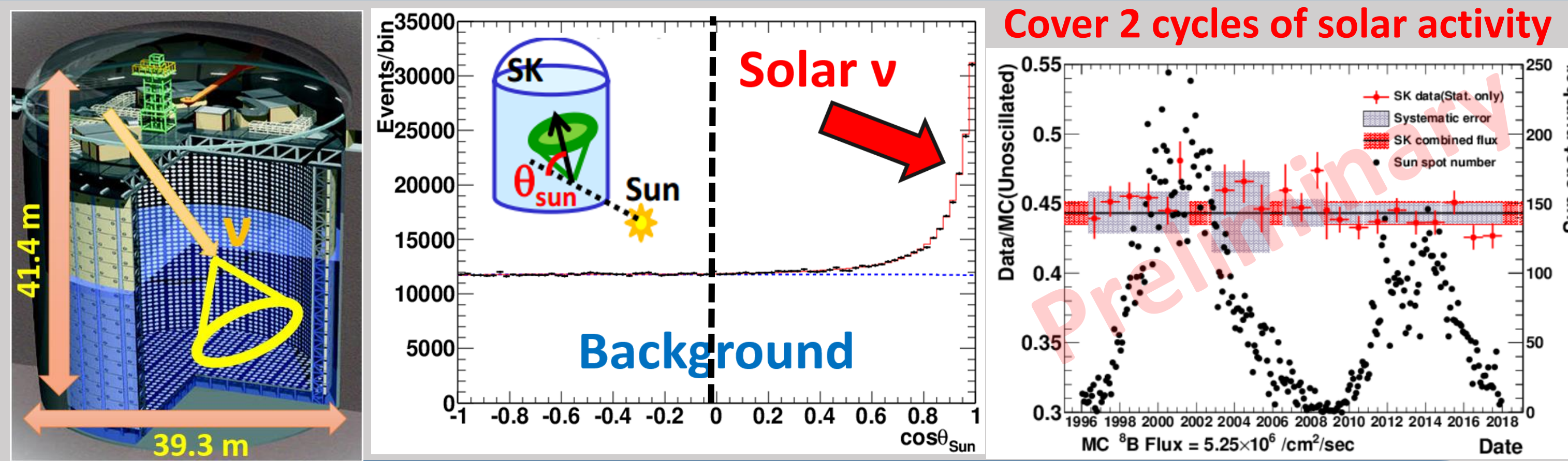
^a Normalized to a velocity of 1 cm s⁻¹ at the solar surface.

3. Super-Kamiokande & Methods

Super-Kamiokande [17] and solar neutrino observation [18-21].

- Cherenkov light produced via $\nu - e$ elastic scattering.
- Recoil electrons **preserve the direction of incident neutrinos**.

Two methods to search for periodical signals.

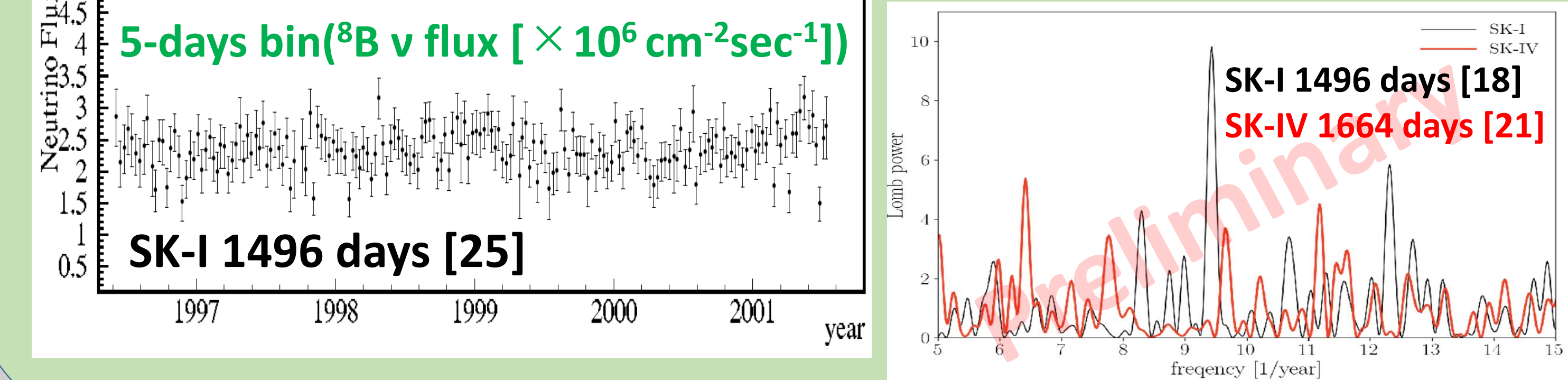


Generalized Lomb-Scargle method (Binned analysis)

Search for periodic signals in uncontinuous data set [22-24].

$$\hat{p}(\omega) = \frac{1}{\sum_i y_i^2} \left\{ \frac{[\sum_i y_i \cos \omega(t_i - \hat{\tau})]^2}{\sum_i \cos^2 \omega(t_i - \hat{\tau})} + \frac{[\sum_i y_i \sin \omega(t_i - \hat{\tau})]^2}{\sum_i \sin^2 \omega(t_i - \hat{\tau})} \right\}$$

y_i : flux of i -th bin
 t_i : time of i -th bin
 ω : angular frequency
 τ : offset

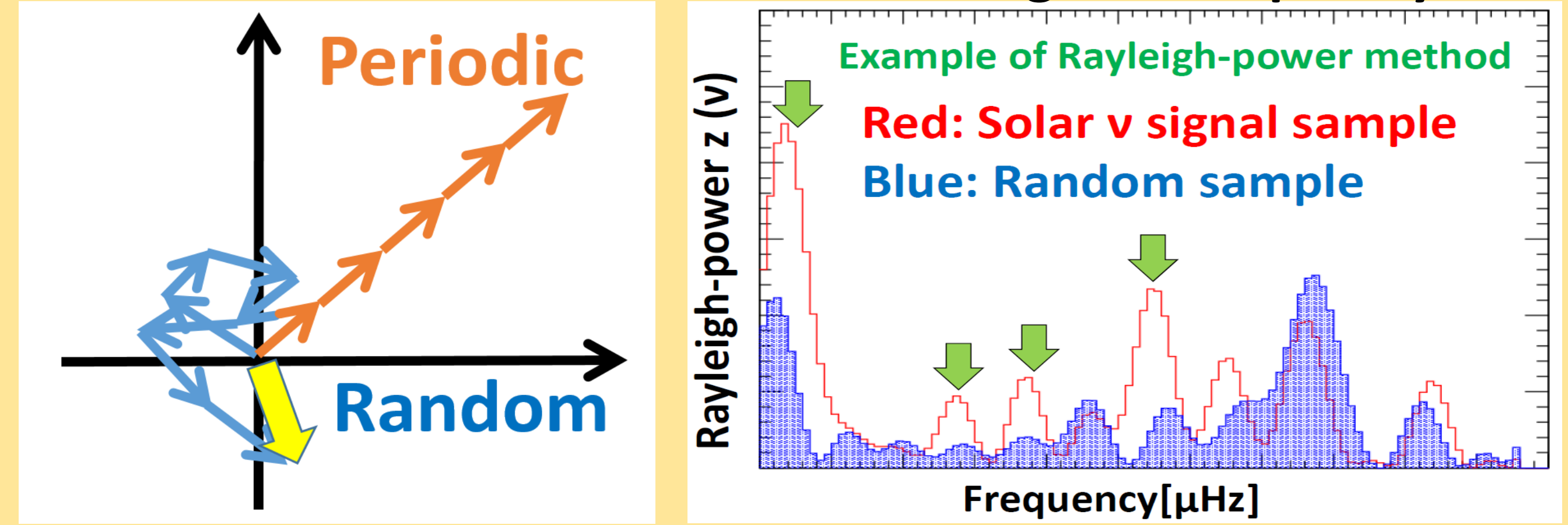


Rayleigh power method (Unbinned analysis)

Power spectrum considering timing of observed events [26, 27].

$$z(\nu) = \frac{1}{N} \times \left[\left(\sum_i \sin 2\pi\nu t_i \right)^2 + \left(\sum_i \cos 2\pi\nu t_i \right)^2 \right]$$

N : number of total event
 t_i : time of i -th event
 ν : given frequency



4. Progress, Future prospect and Summary

Progress and future prospects

- Dead-time due to calibrations affects the sensitivity in Rayleigh power method.
- We have started to develop a randomly timed MC simulation considering operation-time/dead-time of SK.
- We will open the data soon after the MC production and compare with the former results from SNO collaboration [28-30].

Summary

- The SSM has been constructed for 50 years. However, the Sun oscillates around its equilibrium such as p-mode or g-mode.
- G-mode oscillation may amplify the solar ν production rate by a factor of 170 because of fluctuation of electron density.
- SK has a chance to search for g-mode oscillations by using its timing information of the observed solar ν .

Reference: [1] *Rev. Mod. Phys.* 60 (1988) 297. [2] *Astrophys. J.* 765 (2013) 14. [3] *Astrophys. Spa. Sci.* 10 (1971) 87. [4] *Phys. Rev. Lett.* 45 (1980) 963. [5] *Sov. Phys. JETP* 64 (1986) 446. [6] *Mon. Not. R. Astron. Soc.* 101 (1941) 367. [7] *Astrophys. J.* 135 (1962) 474. [8] *Astrophys. J.* 162 (1970) 993. [9] *Solar Phys.* 44 (1975) 371. [10] *Astrophys. J.* 613 (2004) L169. [11] *Astrophys. J.* 409 (1993) L73. [12] *Astrophys. J.* 588 (2003) L65. [13] *Phys. Rev. D* 53 (1996) 4202. [14] *Astrophys. J. Lett.* 792 (2014) L35. [15] *Sov. J. Nucl. Phys.* 42 (1985) 913. [16] *Phys. Rev. D* 17 (1978) 2369. [17] *Nucl. Inst. Meth. A* 501 (2003) 418. [18] *Phys. Rev. D* 73 (2006) 112001. [19] *Phys. Rev. D* 78 (2008) 032002 [20] *Phys. Rev. D* 83 (2011) 052010. [21] *Phys. Rev. D* 94 (2016) 052010. [22] *Astrophys. Spa. Sci.* 49 (1976) 447. [23] *Astrophys. J.* 263 (1982) 835. [24] *A&A* 496 (2009) 577. [25] *Phys. Rev. D* 68 (2003) 092002. [26] *Astrophys. J* 397 (1992) 584. [27] *Mon. Not. R. Astron. Soc.* 268 (1994) 709. [28] *Astrophys. J.* 710 (2010) 540. [29] *Phys. Rev. D* 72 (2005) 052010. [30] *Phys. Rev. D* 75 (2007) 013010.