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ACCURACY OF ESTIMATING HIGHLY ECCENTRIC BINARY BLACK HOLE PARAMETERS WITH GRAVITATIONAL-WAVE DETECTIONS

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ABSTRACT

Mergers of stellar-mass black holes on highly eccentric orbits are among the targets for ground-based gravitational-wave detectors, including LIGO, VIRGO, and KAGRA. These sources may commonly form through gravitational-wave emission in high velocity dispersion systems or through the secular Kozai-Lidov mechanism in triple systems. Gravitational waves carry information about the binaries' orbital parameters and source location. Using the Fisher matrix technique, we determine the measurement accuracy with which the LIGO-VIRGO-KAGRA network could measure the source parameters of eccentric binaries using a matched filtering search of the repeated burst and eccentric inspiral phases of the waveform. We account for general relativistic precession and the evolution of the orbital eccentricity and frequency during the inspiral. We find that the signal-to-noise ratio and the parameter measurement accuracy may be significantly higher for eccentric sources than for circular sources. This increase is sensitive to the initial pericenter distance, the initial eccentricity, and component masses. For instance, compared to a $30\,M_\odot - 30\,M_\odot$ non-spinning circular binary, the chirp mass and sky localization accuracy can improve for an initially highly eccentric binary by a factor of ~ 129 (38) and ~ 2 (11) assuming an initial pericenter distance of $20\,M_{\rm tot}$ ($10\,M_{\rm tot}$).

Keywords: black hole physics – gravitational waves

1. INTRODUCTION

The Advanced Laser Interferometer Gravitational Wave Observatory[1] (aLIGO) detectors (Aasi et al. 2015) and Advanced Virgo[2] (AdV) (Acernese et al. 2015) have made the first six detections of GWs from approximately circular inspiraling binaries (Abbott et al. 2016c,d, 2017a,b,c,d), and opened a new window through which to observe the universe. These advanced gravitational-wave (GW) detectors together with upcoming instruments KAGRA[3] (Somiya 2012) and LIGO-India[4] (Iyer et al. 2011; Abbott et al. 2016f) are expected to continue to detect GW sources in the upcoming years (Abbott et al. 2016f). The orbital eccentricity was neglected in the analysis of the detected GW sources, but a preliminary upper limit was claimed to be $e \lesssim 0.1$ at 10 Hz (Abbott et al. 2016a,e, 2017a,e). In this paper we estimate the future potential of the aLIGO-AdV-KAGRA network of advanced GW detectors to measure the orbital eccentricity and other physical parameters of initially highly eccentric sources.

Initially highly eccentric black hole (BH) binaries are inspiraling systems, which have orbital eccentricities beyond $e_0 \geq 0.9$ when their peak GW frequency (Wen 2003) enters the sensitive frequency band of advanced Earth-based GW detectors. The orbital eccentricity decreases in the inspiral phase from this value until the last stable orbit (LSO) (Peters 1964). Such systems can form in multiple ways including single-single encounters due to GW emission (Kocsis et al. 2006b; O'Leary et al. 2009; Gondán et al. 2017) in dense, high-velocity-dispersion environments; dynamical multibody interactions (Gültekin et al. 2006; O'Leary et al. 2006; Kushnir et al. 2013; Amaro-Seoane & Chen 2016; Antonini & Rasio 2016; Rodriguez et al. 2017); the secular Kozai-Lidov

mechanism (Wen 2003; Thompson 2011; Aarseth 2012; Antonini & Perets 2012; Antognini et al. 2014; Antonini et al. 2014, 2016; Breivik et al. 2016; Rodriguez et al. 2016a,b; VanLandingham et al. 2016; Hoang et al. 2017; Petrovich & Antonini 2017; Silsbee & Tremaine 2017; Randall & Xianyu 2018) in hierarchical triples, or the binary-single interaction (Samsing et al. 2014; Samsing & Ramirez-Ruiz 2017; Samsing 2017; Samsing et al. 2017). Eccentric BH binaries offer promising new detection candidates.

Previous parameter estimation studies of stellar-mass compact binaries have mostly focused on circular binaries (see Finn 1992; Finn & Chernoff 1993; Marković 1993; Cutler & Flanagan 1994; Jaranowski & Krolak 1994; Kokkotas et al. 1994; Królak et al. 1995; Poisson & Will 1995 for the first papers and Chatziioannou et al. 2014; Favata 2014; Mandel et al. 2014; O'Shaughnessy et al. 2014; Rodriguez et al. 2014; Canizares et al. 2015; Veitch et al. 2015; Berry et al. 2015; Miller et al. 2015; Farr et al. 2016; Moore et al. 2016; Lange et al. 2017; Vitale et al. 2017 for recent developments) due to their predicted high detection rates. The current detections constrain the merger rate density of BH-BH mergers in the Universe to $12 - 213\,{\rm Gpc}^{-3}{\rm yr}^{-1}$ (Abbott et al. 2017a), which corresponds to a detection rate between $400 - 7,000\,{\rm yr}^{-1}$ for a typical 2 Gpc detection range for aLIGOs design sensitivity. Furthermore, see Abadie et al. (2010) for a partial list of historical compact binary coalescence rate predictions, and Dominik et al. (2013); Kinugawa et al. (2014); Abbott et al. (2016a,b,h,i,g); Belczynski et al. (2016); Rodriguez et al. (2016a); Bartos et al. (2017); McKernan et al. (2017); Hoang et al. (2017); Stone et al. (2017) and references therein for recent rate estimates.

However, several theoretical studies have shown that the detection rates of highly eccentric BH binaries may be non-negligible. For sources formed by GW-emission in galactic nuclei (GNs), the expected aLIGO detection rate at design sensitivity may be higher than $\approx 100\,{\rm yr}^{-1}$ if the BH mass func-

tion extends to masses above $25\,M_\odot$ (O'Leary et al. 2009; Kocsis & Levin 2012). Recently, such heavy BHs have been observed in several LIGO/VIRGO detections (Abbott et al. 2016b, 2017a,c). Additionally, the expected merger rate densities in the Kozai-Lidov channel are $1-1.5\,\mathrm{Gpc^{-3}\,yr^{-1}}$ for BH binaries forming in nuclear star clusters without supermassive BHs (SMBHs) through multi-body interactions (Antonini & Rasio 2016) and $0.14-6.1\,\mathrm{Gpc^{-3}\,yr^{-1}}$ in isolated triple systems (Silsbee & Tremaine 2017). Of order $1-5\,\mathrm{Gpc^{-3}\,yr^{-1}}$ merger rate density is expected for BH binaries forming via the Kozai-Lidov mechanism in globular clusters (Antonini et al. 2014, 2016; Rodriguez et al. 2016a) and in GNs (Antonini & Perets 2012; Hoang et al. 2017), and non-spherical nuclear star clusters may produce BH binary merger rates of up to $15\,\mathrm{Gpc^{-3}\,yr^{-1}}$ (Petrovich & Antonini 2017). Smaller size GNs with intermediate mass BHs may produce higher rates (VanLandingham et al. 2016). Binary-single gravitational interactions may greatly increase the rates (Samsing et al. 2014; Samsing & Ramirez-Ruiz 2017; Samsing 2017; Samsing et al. 2017). In a companion paper (Gondán et al. 2017), we have shown that GW capture sources in galactic nuclei, which appear to be circular to within $e < 0.2$ near the LSO, may be highly eccentric at the beginning of the detected waveform at 10 Hz, and that heavier BH binaries are expected to be systematically more eccentric in this channel. The ongoing development of detectors towards their design sensitivity at low frequencies may open the possibility of detecting eccentricity in such systems.

In this paper, we determine the expected accuracy with which a network of ground-based interferometric GW detectors may determine the physical parameters that describe highly eccentric BH binaries in comparison to circular sources. We investigate how signal-to-noise ratios (SNRs) and parameter measurement errors depend on the initial orbital parameters, particularly the initial pericenter distance and eccentricity. We examine if it is possible to measure the initial binary parameters (initial eccentricity and pericenter distance) at formation for sources that form in the GW frequency band of the instrument.

Previous GW parameter estimation accuracy studies for eccentric waveforms were carried out for extreme mass ratio (EMRI) sources around SMBHs for LISA (Barack & Cutler 2004; Porter & Sesana 2010; Cornish & Key 2010; Mikóczi et al. 2012; Nishizawa et al. 2016) and for low-eccentricity stellar-mass compact binaries for Earth-based GW detector network (Sun et al. 2015). The premerger localization accuracy of eccentric neutron star binary systems was determined by Kyutoku & Seto (2014), and the source localization accuracy was investigated for low-eccentricity binaries by Ma et al. (2017).

The parameter space of an eccentric spinning binary waveform is generally very large, 17-dimensional (Vecchio 2004; Cornish & Key 2010). Therefore, state-of-the-art methods such as Monte Carlo Markov Chain calculations (see O'Shaughnessy et al. 2014 and references therein) are numerically prohibitively expensive to explore the full range of source parameters for a large set of binaries. For Gaussian noise and a large SNR, the posterior distribution function of the measured parameters is generally well approximated by a multidimensional Gaussian, and the parameter measurement errors can be estimated accurately and very efficiently using the Fisher matrix method (Finn & Chernoff 1993; Cutler & Flanagan 1994; Cutler & Vallisneri 2007). Using this technique, we determine the physical parameters' measurement

accuracy.

We restrict this first study to waveforms introduced by Moreno-Garrido et al. (1994) and Moreno-Garrido et al. (1995), which account for part of the leading order post-Newtonian correction, the GR pericenter precession (hereafter simply precession) and neglect other first post-Newtonian and higher order corrections including those due to spins. Future extensions of this work should include higher order post-Newtonian and merger waveforms (see Levin et al. 2011; Csizmadia et al. 2012; East et al. 2013 for waveform generators, Damour et al. 2004; Memmesheimer et al. 2004; Königsdörffer & Gopakumar 2005, 2006; Yunes et al. 2009; Tessmer & Schäfer 2010, 2011; Huerta et al. 2014; Mikóczi et al. 2015; Moore et al. 2016; Tanay et al. 2016; Boetzel et al. 2017; Cao & Han 2017; Hinderer & Babak 2017; Huerta et al. 2017a,b; Loutrel & Yunes 2017 for analytic waveform models, and Hinder et al. 2008; East et al. 2012; Gold et al. 2012; Gold & Brügmann 2013; East et al. 2015; Paschalidis et al. 2015; East et al. 2016; Lewis et al. 2017 for waveforms of numerical relativity simulations of eccentric compact binary inspirals). In this paper, we focus on BH-BH binaries, but the method is also applicable for neutron star-neutron star and neutron star-black hole binaries on highly eccentric orbits as long as tidal interactions and matter exchange among the components are negligible (see Gold et al. 2012; East et al. 2015, 2016; Radice et al. 2016 and references therein). [5]

Once a large number of GW sources is detected, the correlations between the orbital eccentricity, binary total mass, reduced mass, and spins may be distinctive among different astrophysical mechanisms leading to BH mergers (O'Leary et al. 2009; Cholis et al. 2016; Rodriguez et al. 2016a,b; Chatterjee et al. 2017; Gondán et al. 2017; Samsing & Ramirez-Ruiz 2017; Silsbee & Tremaine 2017; Kocsis et al. 2017). Therefore, detections of eccentric BH binaries have a potential in constraining GW source populations.

However, detecting eccentric sources and recovering their physical parameters is very challenging. So far three search methods were developed to find the signals of stellar-mass eccentric BH binaries in data streams of GW detectors (Tai et al. 2014; Coughlin et al. 2015; Tiwari et al. 2016). All three methods achieve substantially better sensitivity for eccentric BH binary signals than existing localized burst searches or chirp-like template based search methods. Once a source is detected, different algorithms are used to recover its physical parameters. For compact binary coalescences, BAYESTAR (Singer & Price 2016) is an online fast sky localization algorithm that produces probability sky maps, LALINFERENCE (Veitch et al. 2015) is an offline full parameter estimation algorithm, and GSTLAL (Cannon et al. 2012; Privitera et al. 2014) is a low-latency binary BH parameter estimation algorithm. All three algorithms use waveform models of compact binaries on circular orbits. In addition, for short-duration GW "bursts" with poorly modeled or unknown waveforms, CO-HERENT WAVEBURST (Klimenko et al. 2016), BAYESWAVE (Cornish & Littenberg 2015), and LALINFERENCEBURST (Veitch et al. 2015) pipelines produce reconstructed waveforms with minimal assumptions on the waveform morphology. The development of algorithms recovering the parameters of compact binaries on eccentric orbits are currently underway. These algorithms will play an important role for the

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[5] Eccentric neutron star (NS) binaries (NS-NS or NS-BH) will also benefit from additional information if an electromagnetic counterpart is identified, which may lead to smaller parameter errors (Radice et al. 2016).

astrophysical interpretation of eccentric sources.

The paper is organized as follows. In Section 2, we summarize the basic formulae describing the time-domain and frequency-domain eccentric waveform model. In Section 3, we outline the properties of advanced detectors we use in the analysis. In Section 4, we describe the signal parameter measurement estimation method. In Section 5, we discuss which parameters of an eccentric binary can be measured through the binary's waveform. We present our main results in Section 6, and compare or results with previous papers. Finally, we summarize our conclusions in Section 7. Several details about our methodology is included in appendices. In Appendix A, we consider the values of source parameters in the circular limit. Next, in Appendix B, we introduce the geometric conventions we use to describe how the GWs interact with ground-based detectors. In Appendix C, we discuss the applicability of the assumptions of neglecting the Earth's rotation around its axis and Earth's motion around the Sun. In Appendix D, we derive numerically effective formulae to reduce the computational cost of numerical calculations of the SNR and the Fisher matrix. In Appendix E, we present numerical comparisons to validate our codes for both precessing and non-precessing waveforms.

We use $G = 1 = c$ units when referring to the initial orbital parameters, and when determining the phases of waveforms. We work in the observer frame assuming a binary at cosmological redshift z. In this frame, all of the formulae have redshifted mass parameters $m_z = (1+z)m$.[6]

2. ECCENTRIC WAVEFORM MODEL

In this section, we summarize the basic formulae describing the time-domain (Section 2.1) and frequency-domain (Section 2.2) eccentric waveform models including precession in the leading quadrupole-order radiation approximation using the Fourier-Bessel decomposition. Note that we neglect the radiation of higher multipole orders, which are typically subdominant at least in cases where the initial pericenter distance is not close to a grazing or zoom-whirl configuration and the initial velocity is much less than the speed of light (Davis et al. 1972; Berti et al. 2010; Healy et al. 2016).

2.1. *The waveform in time domain*

We adopt the waveform model of Moreno-Garrido et al. (1994) and Moreno-Garrido et al. (1995), which describes the quadrupole waveform emitted by a spinless binary on a Keplerian orbit undergoing slow precession. For a fixed semi-major axis a and orbital eccentricity e, the two polarization states of a GW, h_+ and $h_\times$, with component masses m_A and m_B, and at luminosity distance D_L, can be given in the observer's time-domain as (Moreno-Garrido et al. 1995):

$$h_+(t) = -\frac{h\sin^2\Theta}{2}\sum_{n=1}^{\infty} A_n\cos\Phi_n(t)$$
$$+\frac{h(1+\cos^2\Theta)}{2}\sum_{n=1}^{\infty}\left(B_n^+\cos\Phi_n^-(t)-B_n^-\cos\Phi_n^+(t)\right),\quad(1)$$

$$h_\times(t) = -h\cos\Theta\sum_{n=1}^{\infty}\left(B_n^-\sin\Phi_n^+(t)+B_n^+\sin\Phi_n^-(t)\right),\quad(2)$$

[6] Additional corrections are necessary if the binary has a peculiar velocity (Kocsis et al. 2006a).

where Θ is the angle between the orbital plane and the line-of-sight to the observer, $\Phi_n^\pm$ describe the orbital phase given below for the nth harmonic,

$$h = \frac{4M_{\text{tot},z}\mu_z}{aD_L},\quad(3)$$

where $M_{\text{tot},z} = (m_A + m_B)(1 + z)$ is the redshifted total binary mass at cosmological redshift z, and $\mu_z = (1+z)m_A m_B(m_A + m_B)^{-1}$ is the redshifted reduced mass. We can express the luminosity distance for a flat ΛCDM cosmology as a function of z as

$$D_L = \frac{(1+z)c}{H_0}\int_0^z \frac{dz'}{\sqrt{\Omega_M(1+z')^3+\Omega_\Lambda}},\quad(4)$$

where $H_0 = 68\,\text{km}\,\text{s}^{-1}\text{Mpc}^{-1}$ is the Hubble constant, and $\Omega_M = 0.304$ and $\Omega_\Lambda = 0.696$ are the density parameters for matter and dark energy, respectively (Planck Collaboration et al. 2014a,b).

The A_n and $B_n^\pm$ prefactors in Equations (1) and (2) are the linear combinations of Bessel function of the first kind, $J_n(x)$,

$$A_n = J_n(ne),\quad B_n^\pm = \frac{S_n \pm C_n}{2},\quad(5)$$

where e is the orbital eccentricity,

$$S_n = -\frac{(1-e^2)^{1/2}}{e}\frac{2}{n}J_n'(ne)+\frac{(1-e^2)^{3/2}}{e^2}2nJ_n(ne),\quad(6)$$

$$C_n = -\frac{2-e^2}{e^2}J_n(ne)+\frac{2(1-e^2)}{e}J_n'(ne),\quad(7)$$

and $J_n'(ne)$ is the first derivative of $J_n(ne)$ with respect to e, which satisfies

$$J_n'(ne) = \frac{n}{2}\left[J_{n-1}(ne)-J_{n+1}(ne)\right].\quad(8)$$

We will also need the second derivative of $J_n(ne)$ with respect to e when calculating the Fisher matrix, thus we introduce $J_n''(ne)$ as

$$J_n'' = \frac{n^2}{4}\left[J_{n-2}(ne)-2J_n(ne)+J_{n+2}(ne)\right]\quad\text{if } n\geq 2,\quad(9)$$

$$J_1'' = -\frac{J_1(e)}{2}-\frac{1}{4}\left[J_1(e)-J_3(e)\right].\quad(10)$$

The phase functions $\Phi_n(t)$ and $\Phi_n^\pm(t)$ in Equations (1) and (2) are

$$\Phi_n(t) = \Phi_c - 2\pi n\int_t^{t_c}\nu(t')dt',\quad(11)$$

$$\Phi_n^\pm(t) = \Phi_n(t)\pm 2\gamma(t),\quad(12)$$

where the second term in the right hand side is n times the mean anomaly expressed with the time-integral of the redshifted Keplerian mean orbital frequency ν, Φ_c is the phase extrapolated to coalescence time $t = t_c$, and γ is the azimuthal angle of the pericenter relative to the x axis of the coordinate system defined by the orbital plane. The redshifted Keplerian mean orbital frequency may be expressed with the dimensionless pericenter distance

$$\rho_p = \frac{a(1-e)}{M_{\text{tot},z}}\quad(13)$$

(where a is the semi-major axis in the observer frame) as

$$\nu(e,\rho_p) = \frac{(1-e)^{3/2}}{2\pi\rho_p^{3/2}M_{\text{tot},z}}.\quad(14)$$

For an inspiraling binary, both the eccentricity and the Keplerian orbital frequency evolve in time. Assuming quadrupole radiation and adiabatic evolution of orbital parameters, the equations of time evolution of e and ν, as seen at some cosmological redshift, can be given to leading order as (Peters 1964)

$$\dot{e} = -\frac{304}{15} \frac{e \mathcal{M}_z^{5/3}(2\pi\nu)^{8/3}}{(1-e^2)^{5/2}} \left(1 + \frac{121}{304}e^2\right), \qquad (15)$$

$$\dot{\nu} = \frac{48}{5\pi} \frac{\mathcal{M}_z^{5/3}(2\pi\nu)^{11/3}}{(1-e^2)^{7/2}} \left(1 + \frac{73}{24}e^2 + \frac{37}{96}e^4\right), \qquad (16)$$

where $\mathcal{M}_z = \mu_z^{3/5} M_{\text{tot},z}^{2/5}$ is the redshifted chirp mass, and the overdot denotes a redshifted time-derivative $\dot{x} \equiv dx/dt$. The fraction of the two equations

$$\frac{d\nu}{de} = -\frac{18}{19}\frac{\nu}{e}\frac{\left(1 + \frac{73}{24}e^2 + \frac{37}{96}e^4\right)}{(1-e^2)\left(1 + \frac{121}{304}e^2\right)}, \qquad (17)$$

may be integrated as (Peters 1964; Mikóczi et al. 2012)

$$\nu(e) = c_0/H(e) \qquad (18)$$

where we define

$$H(e) = e^{18/19}(1-e^2)^{-3/2}\left(1 + \frac{121}{304}e^2\right)^{\frac{1305}{2299}}, \qquad (19)$$

and c_0 is an integration constant set by the initial condition $\nu(e_0, \rho_{p0}) = \nu_0$ or the conditions at the LSO, $\nu(e_{\text{LSO}}, \rho_{\text{pLSO}}) = \nu_{\text{LSO}}$ (see Equations (21) and (23) below). Equation (18) shows that the product $c_0 = \nu(e)H(e)$ is conserved during the evolution. Similarly, it is straightforward to determine the evolution of the dimensionless pericenter distance

$$\rho_p(e) = \frac{c_1 e^{12/19}}{M_{\text{tot},z}^{2/3}(1+e)}\left(1 + \frac{121}{304}e^2\right)^{\frac{870}{2299}}$$

$$= \rho_{p0}\frac{e^{12/19}(1+e)^{-1}\left[1 + (121/304)e^2\right]^{\frac{870}{2299}}}{e_0^{12/19}(1+e_0)^{-1}\left[1 + (121/304)e_0^2\right]^{\frac{870}{2299}}}$$

$$= \rho_{\text{pLSO}}\frac{e^{12/19}(1+e)^{-1}\left[1 + (121/304)e^2\right]^{\frac{870}{2299}}}{e_{\text{LSO}}^{12/19}(1+e_{\text{LSO}})^{-1}\left[1 + (121/304)e_{\text{LSO}}^2\right]^{\frac{870}{2299}}} \qquad (20)$$

(Peters 1964), where $c_1 = (2\pi c_0)^{-2/3}$ and in the second and third lines we expressed the evolution with the initial condition ρ_{p0} and e_0, or the "final condition" at the LSO, which satisfies

$$\rho_{\text{pLSO}} = \rho_p(e_{\text{LSO}}) = \frac{6 + 2e_{\text{LSO}}}{1 + e_{\text{LSO}}} \qquad (21)$$

in the leading order approximation in the test mass geodesic zero spin limit (Cutler et al. 1994). This shows that the evolution may be parameterized with the single parameter e_{LSO}, or the two parameters ρ_{p0} and e_0. Note that for any e, the orbital frequency depends only on the single parameter c_0, which is set uniquely by e_{LSO} and $M_{\text{tot},z}$ as

$$c_0 = \frac{1}{2\pi M_{\text{tot},z}}\frac{(1 - e_{\text{LSO}})^{3/2} H(e_{\text{LSO}})}{[\rho_{\text{pLSO}}(e_{\text{LSO}})]^{3/2}}. \qquad (22)$$

We restrict our interest to the repeated burst (O'Leary et al. 2009; Kocsis & Levin 2012) and eccentric inspiral phases of

the waveform model between $0 < e_{\text{LSO}} \leq e \leq e_0 \leq 1$ and neglect the merger and ringdown phases in this analysis. The repeated burst phase starts when the binary is formed with initial eccentricity $e_0 > 0.9$ and initial dimensionless pericenter distance ρ_{p0}, and the eccentric inspiral phase ends when the binary reaches the LSO with eccentricity e_{LSO}. Note that during the evolution e and ρ_p both shrink strictly monotonically in time.

Let us also note for further use, that the Keplerian redshifted orbital frequency at the end of the assumed eccentric inspiral waveform (i.e. at the LSO) is given by Equations (14) and (21) as

$$\nu_{\text{LSO}} = \nu(e_{\text{LSO}}) = \frac{1}{2\pi M_{\text{tot},z}}\left(\frac{1 - e_{\text{LSO}}^2}{6 + 2e_{\text{LSO}}}\right)^{3/2}. \qquad (23)$$

Precession leads to a time-dependent γ in Equation (12). Using the analysis in Mikóczi et al. (2012), we adopt pericenter precession from the classical relativistic motion, and assume that the adiabatic evolution of the orbital parameters are governed by Equations (15) and (16). The angle of precession for a single eccentric orbit in the test particle geodesic approximation around a Schwarzschild BH is

$$\Delta\gamma = \frac{6\pi M_{\text{tot}}}{a(1-e^2)}. \qquad (24)$$

Using an adiabatic approximation, we approximate the redshifted precession rate to be constant during the orbit with

$$\dot{\gamma} \approx \frac{\Delta\gamma}{T} = \frac{3(2\pi\nu)^{5/3}M_{\text{tot},z}^{2/3}}{1-e^2}. \qquad (25)$$

The phase functions given by Equations (11) and (12), can be calculated from Equations (15) and (18) as[7]

$$\Phi_n(t) = \Phi_c + 2\pi n \int_0^{e(t)} \frac{\nu(e')}{\dot{e}(\nu(e'), e')}de' \qquad (26)$$

(Cutler & Flanagan 1994). The phase functions, which arise due to precession, $\Phi_n^+(t)$ and $\Phi_n^-(t)$, follow from Equations (12), (18), (25) and (26)

$$\Phi_n^{\pm}(t) = \Phi_n(t) \pm 2\gamma_c \mp 2\int_t^{t_c} \dot{\gamma}(t')dt' \qquad (27)$$

$$= \Phi_c \pm 2\gamma_c + \int_0^{e(t)} \frac{2\pi n\nu(e') \pm 2\dot{\gamma}(\nu(e'), e')}{\dot{e}(\nu(e'), e')}de',$$

where γ_c is the angle of periapsis extrapolated to coalescence. Note that ρ_{p0}, t_c, γ_c, Φ_c, and e_{LSO} are free parameters of the waveform. Alternatively, we may use the corresponding initial values e_0, t_0, γ_0, Φ_0, and ρ_{p0}.

2.2. The waveform in frequency domain

Since the expressions defining the SNR and the Fisher matrix are both given in Fourier space (Section 4), we construct the Fourier transforms of the waveform [8]

$$\tilde{h}_{+,\times}(f) = \int_{-\infty}^{\infty} h_{+,\times}(t)e^{2\pi if}dt, \qquad (28)$$

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[7] For circular orbits the Fourier phase is conveniently parameterized by ν (Cutler & Flanagan 1994). However, for eccentric inspirals, since $\nu(e)$, $\dot{\nu}(e)$, and $\dot{e}(e)$ are given analytically in the PN approximation, the phase is more conveniently parameterized by e (O'Leary et al. 2009; Mikóczi et al. 2012).

[8] We find that modulations due to Earth's rotation around its axis and Earth's motion around the Sun can be neglected because the signal spends relatively short time in the advanced detectors' sensitive frequency band (see Appendix C).

where $h_+(t)$ and $h_\times(t)$ are given in Equations (1) and (2) as an infinite sum over orbital harmonics n. In the stationary phase approximation each frequency harmonic splits into a triplet due to precession (see Equation (25)) $\mathbf{f} \equiv (f_n, f_n^\pm)$ (Moreno-Garrido et al. 1995; Mikóczi et al. 2012), where

$$f_n = n\nu, \tag{29}$$

$$f_n^\pm = n\nu \pm \frac{\dot{\gamma}}{\pi}, \tag{30}$$

and the Fourier transform simplifies to

$$\tilde{h}_+(\mathbf{f}) = -\frac{h_0}{4}\sin^2\Theta \sum_{n=1}^\infty A_n \Lambda_n e^{i(\Psi_n - \pi/4)}$$
$$-\frac{h_0}{4}(1+\cos^2\Theta)\sum_{n=1}^\infty B_n^+ \Lambda_n^- e^{i(\Psi_n^- - \pi/4)}$$
$$+\frac{h_0}{4}(1+\cos^2\Theta)\sum_{n=1}^\infty B_n^- \Lambda_n^+ e^{i(\Psi_n^+ - \pi/4)}, \tag{31}$$

$$\tilde{h}_\times(\mathbf{f}) = -\frac{h_0}{2}\cos\Theta \sum_{n=1}^\infty B_n^- \Lambda_n^+ e^{i(\Psi_n^+ + \pi/4)}$$
$$-\frac{h_0}{2}\cos\Theta \sum_{n=1}^\infty B_n^+ \Lambda_n^- e^{i(\Psi_n^- + \pi/4)}, \tag{32}$$

where

$$h_0 = \frac{4\mathcal{M}_z^{5/3}(2\pi\nu)^{2/3}}{D_L}, \tag{33}$$

$$\Lambda_n = \frac{1}{\sqrt{|n\dot\nu|}}, \quad \Lambda_n^\pm = \frac{1}{\sqrt{|n\dot\nu \pm \ddot\gamma/\pi|}}, \tag{34}$$

$\dot\nu$ and $\ddot\gamma$ are given by Equations (16), (17) and (25), and $\mathcal{M}_z = (1+z)\mathcal{M}$. The Ψ_n and $\Psi_n^\pm$ phases and their first $(\dot\Psi_n, \dot\Psi_n^\pm)$ and second $(\ddot\Psi_n, \ddot\Psi_n^\pm)$ derivatives with respect to redshifted time t are (Mikóczi et al. 2012)

$$\Psi_n(e, f_n) = 2\pi f_n t_n - \Phi_n, \tag{35}$$
$$\Psi_n^\pm(e, f_n^\pm) = 2\pi f_n^\pm t_n^\pm - \Phi_n^\pm. \tag{36}$$

Here the $(t_n, t_n^\pm)$ parameters of the stationary phase approximation specify the times at which the orbital frequency satisfies Equations (29) and (30) for given $(f_n, f_n^\pm)$, see Appendices A and B in Mikóczi et al. (2012) for details. In Equations (35) and (36), $(\Phi_n, \Phi_n^\pm)$ are to be substituted from Equations (26) and (27). We eliminate $\nu(t)$ and $e(t)$ for $(f_n, f_n^\pm)$ using Equations (18), (22), and (25) together with Equations (29) and (30) to obtain the frequency-domain waveform.[9] The result depends on constant parameters t_c, Φ_c, γ_c, e_{LSO}, $\mathcal{M}_z$, and $M_{tot,z}$. Further, we note that if the precessing eccentric BH binary forms with ρ_{p0} and e_0, then the frequency-domain waveform is truncated at the corresponding minimum frequency $(f_{n,\min}, f_{n,\min}^\pm) = (n\nu_0, n\nu_0 \pm \dot\gamma_0/\pi)$. Furthermore, the waveform model becomes invalid after reaching the LSO (with ρ_{pLSO} and e_{LSO}), which corresponds to a maximum frequency for each harmonic $(f_{n,\max}, f_{n,\max}^\pm) = (n\nu_{LSO}, n\nu_{LSO} \pm \dot\gamma_{LSO}/\pi)$

[9] In practice, there are closed analytic expressions for the e-dependence of ν, $\dot\nu$, $\dot\gamma$, $\ddot\gamma$, Φ_n, $\Phi_n^\pm$, and hence also for f_n and $f_n^\pm$. We must invert these relations $f_n(e)$ and $f_n^\pm(e)$ to obtain the waveform in frequency domain.

Detector	East Long.	North Lat.	Orientation ψ
LIGO H	$-119.4°$	$46.5°$	$-36°$
LIGO L	$-90.8°$	$30.6°$	$-108°$
VIRGO	$10.5°$	$43.6°$	$20°$
KAGRA	$137.3°$	$36.4°$	$65°$

Table 1
Locations and orientations of considered GW detectors in the coordinate system defined in Appendix B. LIGO H marks the Advanced LIGO detector in Hanford, and LIGO L marks the Advanced LIGO detector in Livingston.

where this model is applicable. If we truncate the waveform at these maximum frequencies, this respectively introduces an explicit (ρ_{p0}, e_0) and (ρ_{pLSO}, e_{LSO}) parameter dependence in the waveform model. This is shown in Equation (D11) in Appendix D. Examples of the frequency-domain waveforms are shown in Kocsis & Levin (2012).

In principle, the number of spectral harmonics of an eccentric binary system is infinite. Note however that a large fraction of the signal power is accumulated in a finite number of harmonics. Therefore, in order to reduce the necessary computation time, we truncate n at $n_{\max}(e_0)$ (O'Leary et al. 2009; Mikóczi et al. 2012)

$$n_{\max}(e_0) = \left\{ 5\frac{(1+e_0)^{1/2}}{(1-e_0)^{3/2}} \right\}, \tag{37}$$

which accounts for 99% of the signal power (Turner 1977). Here the bracket {} denotes the floor function. In Appendix D, we discuss other technical details to optimize the calculation of the SNR and the Fisher matrix.

To test our calculations, we examine the limiting cases of no precession ($\dot\gamma \to 0$) and circular orbits ($e \to 0$), respectively. In Appendix D, we discuss numerical and analytic tricks to optimize the calculation and discuss results for the precessing (*Prec*) and precession-free (*NoPrec*) waveform model (i.e. $\dot\gamma \equiv 0$).

3. GW DETECTORS USED IN THE ANALYSIS

Here we summarize the GW detectors and the assumed properties of the detector noise in our analysis.

The aLIGO and AdV detectors completed their first two observing run, and made the first six detections of GWs (Abbott et al. 2016c,d, 2017a,c,d,b). Two additional GW detectors are planned to join the network of aLIGO and AdV; (i) the Japanese KAGRA is under construction with baseline operations beginning in 2018 (Somiya 2012); while (ii) the proposed LIGO-India is expected to become operational in 2022 (Iyer et al. 2011; Abbott et al. 2016f). LIGO-India was approved by the government of India and a study has already suggested site location and orientations of arms for the detector based on scientific figures of merit (Raffai et al. 2013). These parameters, however, have not been finalized yet, and because of this we omit LIGO-India from the analysis.

Due to the expected similarities of design sensitivities of the aLIGO, AdV, and KAGRA detectors within the frequency range of BH inspiral waveforms, for simplicity we adopt the design sensitivities of the two aLIGO (Abbott et al. 2016f) for AdV (Abbott et al. 2016f) and for KAGRA (Somiya 2012) detectors. Table 1 gives the locations and orientations of these detectors, which we used to calculate the response functions. For each detector, we define the detector's orientation angle, ψ, as the angle measured clockwise from North between the x-arm of the detector (see Appendix B for the geometric conventions of detectors) and the meridian that passes through

6

the position of the detector.

We assume that the noise in each detector is stationary colored Gaussian with zero mean, and that it is uncorrelated between different detectors. In reality, detector noise arises from a combination of instrumental, environmental, and anthropomorphic sources that are difficult to characterize precisely (Aasi et al. 2012; Aso et al. 2013; Aasi et al. 2015), and non-Gaussian noise transients (glitches) may arise as well (Blackburn et al. 2008). However, there are existing techniques to identify and remove glitches from GW strain channels and to reduce the level of these artifacts (Littenberg & Cornish 2010; Prestegard et al. 2012; Biswas et al. 2013; Powell et al. 2015; Bose et al. 2016; Torres-Forné et al. 2016; George et al. 2017; Mukund et al. 2017; Powell et al. 2017; Shen et al. 2017). Furthermore, correlated noise between widely separated detectors can arise from so-called Schumann resonances (predicted in Schumann (1952b,a) and observed soon thereafter (Schumann & König 1954; Balser & Wagner 1960)), as well as from other EM phenomena such as solar storms, currents in the van Allen belt (Rycroft 2006), and anthropogenic emission (see Shvets et al. 2010; Thrane et al. 2013, and references therein). Note however, that Schumann resonances mostly affect the stochastic GW background searches (Thrane et al. 2013), and a strategy against such noise artifact already exists (Thrane et al. 2014). Our simplifying assumptions on uncorrelated Gaussian noise are therefore partly justified.

4. OVERVIEW OF THE FISHER MATRIX FORMALISM

In this section, we provide a brief overview of the Fisher matrix method to estimate the measurement errors of physical parameters characterizing a precessing eccentric BH binary source, and refer the reader to Finn (1992) and Cutler & Flanagan (1994) for further details.

The output of a GW detector, $s(t)$, is a combination of a signal, $h(t)$, and a noise term, $n(t)$; i.e.

$$s(t) = h(t) + n(t). \tag{38}$$

We assume the noise of a detector to be stationary, zero mean, and Gaussian, where the different Fourier components of the noise are uncorrelated, i.e.

$$\langle \tilde{n}(f)\tilde{n}^*(f') \rangle = \frac{1}{2}\delta(f-f')S_n(f) \tag{39}$$

(Nissanke et al. 2010), where $\langle \cdot \rangle$ denotes the average, $S_n(f)$ is the one sided noise power spectral density of the detector, the * superscript denotes complex conjugate. With these assumptions, the probability for the noise to have some realization $n_0(t)$ is given as

$$p(n \equiv n_0) \propto e^{-(n_0|n_0)/2} \tag{40}$$

(Finn 1992), where $p(n)$ is the probability distribution function of the noise to assume a value n, and $(\dots|\dots)$ denotes the following inner product between any two functions of frequency, e.g. x(f) and y(f):

$$(x \mid y) \equiv 4 \int_0^\infty \frac{\tilde{x}(f)\tilde{y}^*(f)}{S_n(f)} df. \tag{41}$$

The optimal SNR is given by the standard expression

$$\frac{S}{N} = \sqrt{(h \mid h)} = \sqrt{4 \int_0^\infty \frac{|\tilde{h}(f)|^2}{S_n(f)} df}. \tag{42}$$

Here the signal waveform, $h(f)$, depends on the parameter set $\{\lambda_p \mid p \in \{1\dots P\}\}$, which characterizes the

source. For a large SNR, the parameter estimation errors $\Delta\lambda = \{\Delta\lambda_p \mid p \in \{1,\dots,P\}\}$ defined as the measured value minus the true value have the Gaussian probability distribution for a given signal

$$p(\Delta\lambda) = \mathcal{N} \exp\left(-\frac{1}{2}\Gamma_{ij}\Delta\lambda_i\Delta\lambda_j\right) \tag{43}$$

(Finn 1992), where $\mathcal{N}$ is a normalization constant. In Equation (43), we assume summation over repeated indices, and Γ_{ij} is the Fisher information matrix defined as

$$\Gamma_{ij} \equiv (\partial_i h \mid \partial_j h) = 4 \int_0^\infty \frac{\Re(\partial_i\tilde{h}^*(f)\partial_j\tilde{h}(f))}{S_n(f)} df, \tag{44}$$

where $\partial_i h = \partial h/\partial\lambda_i$ and $\Re$ labels the real part.

Following Cutler & Flanagan (1994), we define the combined signal-to-noise ratio of a network of detectors (SNR$_{tot}$) as an uncorrelated superposition of individual SNRs

$$\left(\frac{S}{N}\right)_{tot}^2 = \sum_{k=1}^{N_{det}} \left(\frac{S}{N}\right)_k^2, \tag{45}$$

where the number of detectors in the network is denoted by N_{det}, and $(S/N)_k$ denotes the SNR in the kth detector.

Similarly, for uncorrelated Gaussian noise, the Fisher matrix of a network of detectors is the sum of the Fisher matrices of individual detectors,

$$\Gamma_{ij,tot} = \sum_{k=1}^{N_{det}} \Gamma_{ij,k}. \tag{46}$$

The covariance matrix is defined with the inverse of the Fisher matrix:

$$\Sigma_{ij} = (\Gamma_{ij,tot})^{-1} = \langle \Delta\lambda_i\Delta\lambda_j \rangle, \tag{47}$$

where the angle brackets denote an average over the probability distribution function in Equation (43). The root-mean-square parameter measurement error σ_i in the parameters λ_i marginalized over all other parameters is

$$\sigma_i = \langle(\Delta\lambda_i)^2\rangle^{1/2} = \sqrt{\Sigma_{ii}}. \tag{48}$$

The off-diagonal elements of Σ_{ij} give the cross-correlation coefficients between parameters λ_i and λ_j.

Parameters can be measured independently if the corresponding Fisher matrix $\Gamma_{ij,tot}$ is nonsingular. Otherwise, if the Fisher matrix is singular, then the eigenvector(s) corresponding to the zero-eigenvalue(s) of the Fisher matrix represent the linear combination(s) of the parameters, which cannot be measured by the network.

We derive efficient formulae to compute the SNR and the Fisher matrix in Appendix D.

5. MEASURING THE PARAMETERS OF PRECESSING ECCENTRIC BLACK HOLE BINARIES

In this section, we identify the parameters of a precessing eccentric binary that can be extracted from the detected waveform for the signal model introduced in Section 2. We set the parameters in our calculations and measure their errors as follows.

- D_L : We set $D_L = 100\,\text{Mpc}$, and measure its relative error $\langle\Delta D_L^2\rangle^{1/2}/D_L = \langle(\Delta\ln D_L)^2\rangle^{1/2}$. This choice is arbitrary, smaller than the nearest circular BH-BH merger detection to date, $340 \pm 140\,\text{Mpc}$ (Abbott et al. 2017b).

The Fisher matrix method gives accurate results for the parameter measurement errors for high SNR. For moderately larger distances, the errors scale as $\propto D_L$.

- θ_N and ϕ_N: We generate an isotropic random sample of the sky position angles θ_N and ϕ_N by drawing $\cos\theta_N$ and ϕ_N from a uniform distribution between $[-1,1]$ and $[0,2\pi]$, and calculate the parameter estimation covariance for each sample. The errors of the sky position is described by a localization ellipse. We characterize the sky localization accuracy either by the corresponding proper angular length of the semi-major and semi-minor axes of the sky-localization error ellipsoid given by Lang & Hughes (2006), (a_N, b_N), or its proper solid angle $\Delta\Omega_N = \sqrt{\pi a_N b_N}$. The calculated results are valid if $a_N \ll 1$ radian and $b_N \ll 1$ radian.

- θ_L and ϕ_L: We draw the angular momentum vector direction angles from an isotropic distribution and construct their error ellipsoids or solid angles similar to that given for θ_N and ϕ_N.

- m_A and m_B: We fix the fiducial component masses to $m_A = m_B = 30\,\mathrm{M}_\odot$, consistent with the first discovered source GW150914 (Abbott et al. 2016d). Such high mass sources are expected in galactic nuclei since mass segregation helps to increase their numbers relative to the lower mass binaries, and the SNR is also higher for these binaries (O'Leary et al. 2009). Since we neglect additional post-Newtonian corrections of the GW phase, we restrict the measurement error estimation to $\mathcal{M}_z$ for calculations evaluated for comparison in which we neglect precession. However, generally the assumed precessing eccentric waveform model depends on two independent combinations of component masses: $\mathcal{M}_z$ sets the inspiral rate, and $M_{\mathrm{tot},z}$ sets both the apsidal precession rate and the final frequency at the LSO. We calculate the relative errors for both of these mass parameters and for the precessing eccentric waveform model $\langle\Delta\mathcal{M}_z^2\rangle^{1/2}/\mathcal{M}_z = \langle(\Delta\ln\mathcal{M}_z)^2\rangle^{1/2}$, and similarly for $M_{\mathrm{tot},z}$.

- t_c, Φ_c, and γ_0: These parameters only enter in the complex phase of the waveform through Ψ_n and $\Psi_n^\pm$ (see Equations (35) and (36)), but do not affect the SNR. Since these parameters are responsible for an overall phase shift of the waveform, we do not randomize their values but assume the fiducial value $t_c = \Phi_c = \gamma_c = 0$ for each binary in the Monte Carlo sample.

- e_{LSO}: The adopted eccentric inspiral waveform model depends explicitly on the final eccentricity at the LSO, see Equation (21). This quantity parameterizes the evolutionary path of the binary during its eccentric inspiral in the (ρ_p, e) plane as shown in Equations (20) and (21); see also Figure 3 in Kocsis & Levin (2012) for illustration. In fact, any segment of the evolutionary path $\rho_p(e)$ specifies the value of e_{LSO} uniquely. Conversely, e_{LSO} specifies $\rho_p(e)$, which sets a constraint on the possible values of (ρ_{p0}, e_0), if the post-Newtonian binary inspiral model is extrapolated backwards in time. Indeed, in some cases this is the only indirect information we may have on the formation parameters (ρ_{p0}, e_0). In particular, $e_0 \le 1$ puts an upper bound on ρ_{p0} for a given

e_{LSO}.[10]

- e_0: We choose several e_0 values from the highly eccentric ($e_0 \ge 0.9$) limit when discussing the e_0 dependence of measurement errors (see Section 6.2). However, we restrict to $e_0 = 0.9$ for calculations of a large survey of binaries.[11]

- ρ_{p0}: We examine two values for the dimensionless initial pericenter distance $\rho_{p0} = \{10, 20\}$, and the circular limit corresponds to $\rho_{p0} \to \infty$ (O'Leary et al. 2009). These values are likely for sources that form through the GW capture mechanism in high velocity dispersion environments such as GN as shown in O'Leary et al. (2009); Gondán et al. (2017) or the core collapsed regions of star clusters without a central massive black hole (Kocsis et al. 2006b; Antonini & Rasio 2016).

If the peak GW frequency of the initial orbit is large enough to be in the detectors' sensitive frequency band then ρ_{p0} and e_0 are directly measurable due to the truncation of the time-domain waveform for times when $e < e_0$ and $\rho_p > \rho_{p0}$. In the opposite case only a lower limit may be given for ρ_{p0}, which corresponds to $e_0 \to 1$ (Kocsis & Levin 2012).

In summary, we use the following free parameters in the Fisher matrix analysis:

$$\boldsymbol{\lambda}_{\mathrm{Prec}} = \{\ln(D_L), \ln(\mathcal{M}_z), \ln(M_{\mathrm{tot},z}), \theta_N, \phi_N, \theta_L, \phi_L, e_0,$$
$$e_{\mathrm{LSO}}, t_c, \Phi_c, \gamma_c\}. \tag{49}$$

Given these parameters, other parameters' marginalized measurement errors may be determined by linear combinations of the covariance matrix based on Equation (48). For example, ρ_{p0} is given by e_0 and e_{LSO} using Equations (20) and (21). Its measurement error is

$$\langle(\Delta\rho_{p0})^2\rangle = \left(\frac{\partial\rho_p(e_0, e_{\mathrm{LSO}})}{\partial e_0}\right)^2 \langle(\Delta e_0)^2\rangle$$
$$+ \left(\frac{\partial\rho_p(e_0, e_{\mathrm{LSO}})}{\partial e_{\mathrm{LSO}}}\right)^2 \langle(\Delta e_{\mathrm{LSO}})^2\rangle$$
$$+ 2\frac{\partial\rho_p(e_0, e_{\mathrm{LSO}})}{\partial e_0}\frac{\partial\rho_p(e_0, e_{\mathrm{LSO}})}{\partial e_{\mathrm{LSO}}}\langle\Delta e_0\Delta e_{\mathrm{LSO}}\rangle. \tag{50}$$

The parameter estimation errors of individual component masses or the mass ratio can be estimated similarly using ΔM_{tot} and $\Delta\mathcal{M}$ after inverting $M_{\mathrm{tot}}(m_a, m_b)$ and $\mathcal{M}(m_a, m_b)$.

6. RESULTS

The measurement errors depend on the sky position of the source with respect to the detectors and on the relative orientation of the binary. We generate random Monte Carlo samples of ~ 4500 binaries by drawing from isotropic distributions of the sky position and of the binary orientation normal vector. We present the results for the $\mathrm{SNR}_{\mathrm{tot}}$ for detecting precessing highly eccentric BH binaries with the GW detector network described in Table 1 and for the expected parameter measurement errors.

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[10] When studying the measurement errors for non-precessing eccentric binaries, the waveform depends explicitly on a single combination of e_{LSO} and $M_{\mathrm{tot},z}$ parameters c_0 (Section 2.1). Therefore, we use c_0 for the *NoPrec* model to avoid a singularity of the Fisher matrix.

[11] We note that the e_0 dependence of the waveform is due to the truncation of the time-domain waveform for times when $e < e_0$.
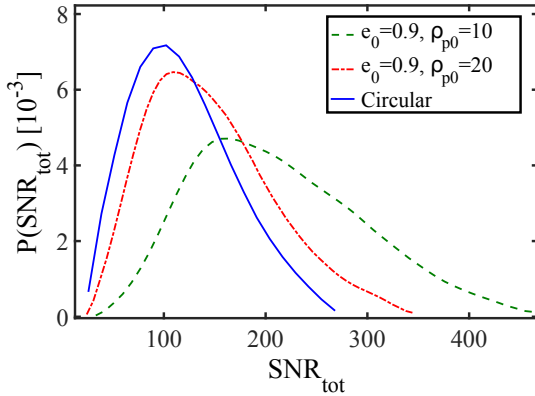
Figure 1. Smoothed probability density function of the total network signal-to-noise ratio (SNR_{tot}) of gravitational-wave detection from $30 M_\odot - 30 M_\odot$ precessing eccentric BH binaries with initial eccentricity $e_0 = 0.9$ and dimensionless pericenter distance $\rho_{p0} = 10$ and 20 (green dashed and red dash-dot) and similar binaries in the circular limit (blue solid) at luminosity distance $D_L = 100\,$Mpc with a random source direction and orientation. Distributions correspond to a Monte Carlo sample of 4500 binaries. Parameters of the assumed detector network are given in Table 1. The medians of SNR_{tot} distributions are 108.7, 202.7, and 137.3 in the circular limit and for binaries with $\rho_{p0} = 10$ and 20, respectively. Systematically higher SNR_{tot} values for precessing highly eccentric BH binaries implies that they are detectable to a larger distance compared to precessing eccentric BH binaries in the circular limit.

6.1. *Signal-to-noise ratio distributions*

Figure 1 displays the distribution of the SNR_{tot} for precessing highly eccentric BH binaries detected with the detector network described in Table 1, assuming binary parameters $m_A = m_B = 30 M_\odot$, $e_0 = 0.9$, $\rho_{p0} = \{10, 20\}$, and similar binaries in the circular limit (see Appendix A for details). Generally, similar to the results of O'Leary et al. (2009) (see Figure 11 therein, which corresponds to a single aLIGO detector), the SNR_{tot} is systematically higher for binaries with $\rho_{p0} = 10$ than for binaries with $\rho_{p0} = 20$. We find that increasing the initial eccentricity from $e_0 = 0.9$ to 0.97 for fixed ρ_{p0} does not change the SNR_{tot} significantly (see Table 3 below), hence we expect that the distribution of SNR_{tot} for fixed ρ_{p0} converges in the $e_0 \to 1$ limit. This is expected since Figure 10 in O'Leary et al. (2009) shows that a low amount of the SNR_{tot} accumulates near $e_0 \approx 1$ for low to moderately high BH masses.

Figure 2 shows that the SNR_{tot} increases rapidly with ρ_{p0} for low ρ_{p0}, has a maximum between $\rho_{p0} \sim 9$ and 20, and decreases for higher ρ_{p0} approaching the circular binary limit for $\rho_{p0} \to \infty$. These findings may be understood qualitatively as follows. Within $\rho_{p0} \leq 40.9 (M_{\mathrm{tot},z}/20 M_\odot)^{-2/3}$ the binary forms with a characteristic frequency above $10\,$Hz in the detector band (see Equation (59) in Gondán et al. 2017). The rapid decrease of the SNR_{tot} for decreasing $\rho_{p0} < 9$ is due to the fact that we neglect the GWs of the first hyperbolic encounter (Kocsis et al. 2006b). The decrease of the SNR_{tot} at high ρ_{p0} is due to the fact that part of the GW spectrum falls outside of the detectors' sensitive frequency band. For very large $\rho_{p,0}$, the binary becomes circular by the time it enters the detectors' sensitive frequency band, and a significant fraction of the SNR_{tot} accumulates only in the $n = 2$ harmonic (Figure 9), which explains the flat asymptotics for high ρ_{p0}. Decreasing ρ_{p0} from high to moderate values higher harmonics start to contribute to the SNR_{tot} (Figure 9), which explains the increase of the SNR_{tot}. A combination of these arguments leads to the peak of the SNR_{tot} at an intermediate ρ_{p0} value seen in Figure 2. However, note that in addition to ne-
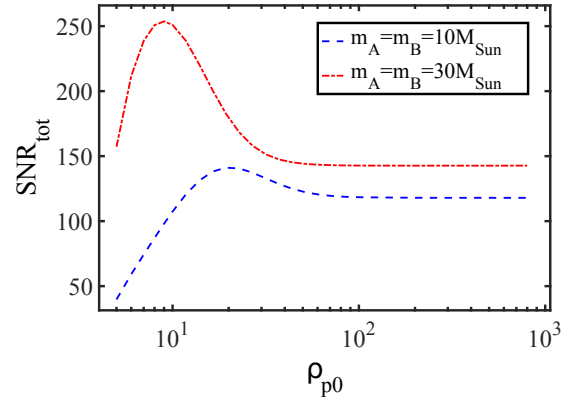
Figure 2. The SNR_{tot} of precessing eccentric BH binaries as a function of their initial dimensionless pericenter distance ρ_{p0}. Parameters of the assumed detector network are given in Table 1. Here the luminosity distance is $D_L = 100\,$Mpc, the initial eccentricity is $e_0 = 0.9$, the source direction and orientation angular parameters are fixed at $\theta_N = \pi/2$, $\phi_N = \pi/3$, $\theta_L = \pi/4$, and $\phi_L = \pi/5$. Depending on the binary mass, the SNR_{tot} has a maximum at ρ_{p0} between 9 and 20, and converges asymptotically to the value of precessing eccentric binaries in the circular limit for high ρ_{p0}. Note that we find similar trends with ρ_{p0} for other random choices of binary direction and orientation (not shown).

glecting the initial hyperbolic encounter and the final coalescence/ringdown segments of the signal, our waveform model also neglects contributions of spherical moments beyond the quadrupole-order and deviations from a precessing Keplerian orbit (Davis et al. 1972; Berti et al. 2010; Healy et al. 2016). This approximation may not be valid for low ρ_{p0} (particularly for $\rho_{p0} \leq 10$). The SNR_{tot} is expected to be underestimated in this region in Figure 5.

6.2. *Parameter measurement errors*

We present the measurement accuracy for the final eccentricity at the LSO for parameters grouped as

$$\boldsymbol{\lambda}_{\mathrm{slow}} = \{\ln(D_L), \theta_N, \phi_N, \theta_L, \phi_L\}, \qquad (51)$$

$$\boldsymbol{\lambda}_{\mathrm{fast}} = \{\Phi_c, t_c, \ln(\mathcal{M}_z), \ln(M_{\mathrm{tot},z}), e_{\mathrm{LSO}}, \gamma_c\} \qquad (52)$$

(Kocsis et al. 2007). The $\boldsymbol{\lambda}_{\mathrm{fast}}$ fast parameters are related to the high frequency GW phase, while the $\boldsymbol{\lambda}_{\mathrm{slow}}$ slow parameters appear only in the slowly-varying amplitude of the GW signal. Slow parameters are mostly determined by a comparison of the GW signals measured by the different detectors in the network. For the polar angles (θ_N, ϕ_N) describing the source direction, we calculate the minor and major axes (a_N, b_N) of the corresponding 2D sky location error ellipse and its area $(\Omega_N = \pi a_N b_N)$, and we do the same for the binary orientation error ellipse (a_L, b_L) and its area $(\Omega_N = \pi a_L b_L)$.

Figures 3 and 4 show the distribution of the measurement errors for randomly chosen source sky position and binary orientation for $\rho_{p0} = 10$ and 20, and for the circular limit $\rho_{p0} \to \infty$ (see Appendix A), while Table 2 shows the 10%, 50%, and 90% quantiles of the error distributions. Compared to a $\rho_{p0} = 20$ binary, a $\rho_{p0} = 10$ binary is more eccentric throughout its evolution, which leads to a higher SNR_{tot}, and most of its measurement errors are smaller. There are, however, exceptions to this finding: the fast parameters such as the mass parameters and the eccentricity have higher errors for $\rho_{p0} = 10$ than for $\rho_{p0} = 20$ (see discussion below).

Many of the binaries in galactic nuclei[12] form with very

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[12] particularly the heavy BHs therein (Gondán et al. 2017)

ρ_{p0}	10	10	10	20	20	20	Circular	Circular	Circular
quantile	10%	50%	90%	10%	50%	90%	10%	50%	90%
$\Delta(\ln D_L)$	1.13(−2)	2.49(−2)	7.21(−2)	1.89(−2)	4.36(−2)	0.154	8.81(−3)	5.86(−2)	0.311
$\Delta\Omega_N$ [sr]	1.13(−5)	1.46(−4)	7.23(−4)	2.21(−4)	7.28(−4)	2.89(−3)	5.65(−5)	1.58(−3)	6.67(−3)
$\Delta\Omega_L$ [sr]	7.89(−5)	3.85(−3)	0.13	2.21(−5)	1.13(−2)	0.85	4.36(−4)	2.21(−2)	9.6
a_N [deg]	0.14	0.62	2.42	0.57	1.36	4.07	0.28	2.01	5.16
b_N [deg]	0.11	0.22	0.48	0.26	0.52	0.98	0.41	0.76	1.44
a_L [deg]	0.96	2.78	20.04	1.44	4.65	45.48	0.78	6.98	2.02(+2)
b_L [deg]	0.55	1.33	7.51	0.84	2.46	20.58	0.32	3.41	50.63
$\Delta\Phi_c$ [rad]	9.27(−2)	0.23	0.71	0.29	0.66	2.44	0.36	0.72	57.91
Δt_c [ms]	4.32(−2)	8.41(−2)	0.181	9.28(−2)	0.167	0.311	9.41(−2)	1.582	3.011
$\Delta(\ln\mathcal{M}_z)$	3.53(−5)	6.17(−5)	1.71(−4)	1.04(−5)	1.81(−5)	3.42(−5)	1.36(−3)	2.34(−3)	4.27(−3)
$\Delta(\ln M_{\mathrm{tot},z})$	5.42(−4)	9.51(−4)	2.43(−3)	2.82(−4)	4.81(−4)	9.18(−4)	5.88(−3)	1.13(−2)	1.81(−2)
Δe_{LSO}	1.18(−4)	2.16(−4)	5.83(−4)	2.39(−5)	3.19(−5)	5.88(−5)	–	–	–
Δe_0	1.44(−3)	2.16(−3)	3.95(−3)	1.72(−3)	2.91(−3)	5.79(−3)	–	–	–
$\Delta\rho_{p0}$	6.64(−3)	1.08(−2)	2.28(−2)	1.33(−2)	2.29(−2)	4.58(−2)	–	–	–
$\Delta\gamma_c$ [rad]	0.04	0.11	0.47	0.14	0.34	1.51	–	–	–

Table 2
The 10%, 50%, and 90% quantile of measurement errors for parameters of $30\,\mathrm{M}_\odot - 30\,\mathrm{M}_\odot$ precessing eccentric BH binaries with initial eccentricity $e_0 = 0.9$ and dimensionless pericenter distance $\rho_{p0} = 10$ and 20, and circular binaries at distance $D_L = 100\,\mathrm{Mpc}$, random source sky location and orientation using the detector network in Table 1. Here (Ω_N, a_N, b_N) are respectively the area, semi-major and semi-minor axes of the 2D error ellipse corresponding to the source's sky direction, and similarly for (Ω_L, a_L, b_L) describing the source's orbital plane orientation (i.e. angular momentum vector direction). Note that $e_{\mathrm{LSO}} = 0.187$ and 0.059 for $\rho_{p0} = 10$ and 20, respectively, if $e_0 = 0.9$ is assumed. In the circular limit the binary forms outside of the sensitive frequency band of the detector network, and $\Delta e_0 \to \infty$ and $\Delta\rho_{p0} \to \infty$. We adopt the following notation in the table: $1.13(−2) = 1.13 \times 10^{-2}$.

| ρ_{p0}	10	10	10	20	20	20
e_0	0.9	0.95	0.97	0.9	0.95	0.97
e_{LSO}	0.1872	0.1932	0.1956	5.89(−2)	6.07(−2)	6.15(−2)
$\mathrm{SNR}_{\mathrm{tot}}$	251.1	257.7	260.4	179.4	182.1	182.9
$\Delta(\ln D_L)$	3.92(−2)	3.79(−2)	3.76(−2)	6.41(−2)	6.31(−2)	6.27(−2)
$\Delta\Omega_N$ [sr]	8.37(−5)	7.78(−5)	7.53(−5)	4.06(−4)	3.91(−4)	3.85(−4)
$\Delta\Omega_L$ [sr]	9.68(−3)	9.07(−3)	8.84(−3)	2.74(−2)	2.65(−2)	2.61(−2)
a_N [deg]	0.501	0.484	0.477	0.994	0.977	0.970
b_N [deg]	0.174	0.168	0.165	0.427	0.418	0.415
a_L [deg]	4.31	4.18	4.13	6.96	6.85	6.81
b_L [deg]	2.35	2.26	2.23	4.12	4.04	4.01
$\Delta\Phi_c$ [rad]	0.331	0.313	0.308	1.102	1.049	1.031
Δt_c [ms]	9.68(−2)	9.08(−2)	8.88(−2)	0.198	0.192	0.189
$\Delta(\ln\mathcal{M}_z)$	4.61(−5)	1.89(−5)	1.39(−5)	1.38(−5)	6.03(−6)	4.33(−6)
$\Delta(\ln M_{\mathrm{tot},z})$	7.09(−4)	5.62(−4)	4.85(−4)	3.67(−4)	2.87(−4)	2.38(−4)
Δe_{LSO}	1.61(−4)	1.28(−4)	1.11(−4)	2.43(−5)	1.93(−5)	1.62(−5)
Δe_0	1.95(−3)	1.74(−3)	1.69(−3)	2.25(−3)	2.17(−3)	2.14(−3)
$\Delta\rho_{p0}$	8.33(−3)	7.45(−3)	7.11(−3)	1.71(−2)	1.61(−2)	1.56(−2)
$\Delta\gamma_c$ [rad]	0.169	0.162	0.159	0.556	0.529	0.520

Table 3
Measurement errors for parameters of $30\,\mathrm{M}_\odot - 30\,\mathrm{M}_\odot$ precessing eccentric BH binaries with initial eccentricities $e_0 = 0.9$, 0.95, and 0.97 for initial dimensionless pericenter $\rho_{p0} = 10$ and 20, luminosity distance $D_L = 100\,\mathrm{Mpc}$, and arbitrarily fixed source direction $(\theta_N, \phi_N) = (\pi/2, \pi/3)$ and orientation $(\theta_L, \phi_L) = (\pi/4, \pi/5)$ for the detector network specified in Table 1.

high e_0, close to unity, in single-single encounters due to GW emissions (O'Leary et al. 2009; Gondán et al. 2017). However, similarly to the finding that the $\mathrm{SNR}_{\mathrm{tot}}$ does not increase significantly for $1 > e_0 \geq 0.9$, we find that $\boldsymbol{\lambda}_{\mathrm{slow}}$ parameter errors do not improve due to the early very eccentric evolutionary period beyond $e > 0.9$ (repeated burst phase) compared to waveforms with $e_0 = 0.9$ as shown in Table 3. However, some $\boldsymbol{\lambda}_{\mathrm{fast}}$ parameters' measurement errors improve more significantly with $e_0 > 0.9$. In particular, the measurement errors of the mass parameters $(\mathcal{M}_z, M_{\mathrm{tot},z})$ improve by a factor of ~ 2, and the measurement error of e_{LSO} improves by $\sim 50\%$ if increasing e_0 from 0.9 for 0.97. This difference is due to the fact that eccentricity modifies the GW phase significantly, which affects the determination of $\boldsymbol{\lambda}_{\mathrm{fast}}$ parameters only.

We calculate the parameter measurement errors for pre-cessing highly eccentric BH binaries as a function of ρ_{p0} for some arbitrarily fixed binary direction and orientations. For one such binary direction and orientation Figure 5 shows the ρ_{p0} dependence of $\Delta D_L/D_L$, $\Delta\mathcal{M}_z/\mathcal{M}_z$, Δe_0, $\Delta\rho_{p0}$, Δe_{LSO}, semi-major and semi-minor axes of the sky position error ellipse (a_N, b_N), and semi-major and semi-minor axes of the error ellipse for the binary orbital plane normal vector direction (a_L, b_L). Note that we find similar trends with ρ_{p0} for other random choices of binary direction and orientation. We find that measurement errors systematically decrease with decreasing ρ_{p0} for precessing highly eccentric binaries relative to similar binaries in the circular limit (Appendix A), and the errors have a minimum in the range $8 < \rho_{p0} < 80$ and deteriorate rapidly for $\rho_{p0} < 8$. The latter is due to the rapid decrease of the $\mathrm{SNR}_{\mathrm{tot}}$ in that range. The ρ_{p0} dependence of
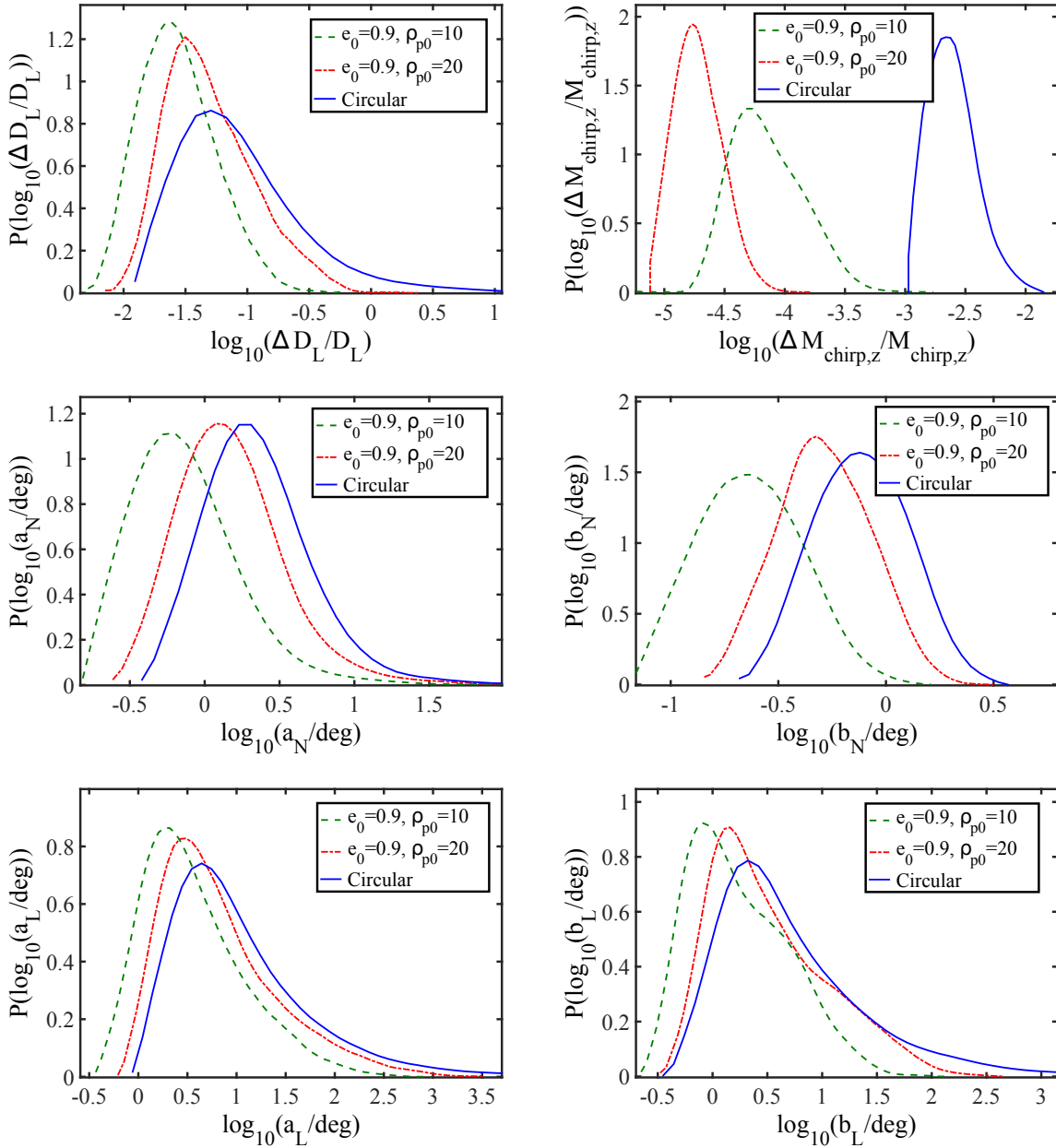
Figure 3. Smoothed distribution of the measurement errors of parameters measured for $30\,M_\odot - 30\,M_\odot$ precessing eccentric BH binaries with initial eccentricity $e_0 = 0.9$ and dimensionless pericenter distance $\rho_{p0} = 10$ and 20 (green dashed and red dash-dot) and for similar binaries in the circular limit (blue solid) at luminosity distance $D_L = 100\,\text{Mpc}$ with a random source direction and binary orientation. We have assumed the detector network specified in Table 1. *Top row:* Distribution of the relative measurement error of luminosity distance, $\Delta D_L/D_L = \Delta(\ln D_L)$, and redshifted chirp mass, $\Delta\mathcal{M}_z/\mathcal{M}_z = \Delta(\ln\mathcal{M}_z)$. *Middle row:* Distribution of semi-major axis of the sky position error ellipse, a_N, and its semi-minor axis, b_N. *Bottom row:* Distribution of the semi-major axis of the binary's orbital plane orientation error ellipse, a_L, and its semi-minor axis, b_L.

$\Delta D_L/D_L$ and the principal axes of the sky position and binary orientation error ellipses (a_N, b_N) and (a_L, b_L) (i.e. quantities derived from slow parameters, see Section 5) are qualitatively similar to that of $1/\text{SNR}_\text{tot}$ in the complete range of ρ_{p0} (i.e. they decrease rapidly with ρ_{p0} for low ρ_{p0}, have a minimum at moderate ρ_{p0}, and converge asymptotically to the value of precessing highly eccentric binaries in the circular limit), see Figure 6 for details. However, Figure 5 shows that the chirp mass errors have a minimum at much higher ρ_{p0}, i.e. between $50-60$ and $20-40$ for $10\,M_\odot - 10\,M_\odot$ and $30\,M_\odot - 30\,M_\odot$ precessing highly eccentric BH binaries, respectively. The main reason for the different behavior of the chirp mass from

the distance and angular errors is the fact that the chirp mass is a fast parameter, while the distance and angular parameters are slow parameters. Slow parameters are insensitive to the GW phase perturbations, and depend on the GW amplitude, which is set by the SNR_tot. The SNR of the early part of the waveform near the low-frequency noise wall of the detector is small. However, fast parameters depend sensitively on the GW phase, and the GW phase accumulates mostly at low frequencies, since the residence time (i.e. $\nu/\dot{\nu}$) is largest at low orbital frequencies. Thus, the fast parameters' errors are minimized for binaries which form with $e_0 \sim 1$ with a ρ_{p0} value for which the GW characteristic frequency is near the detectors'
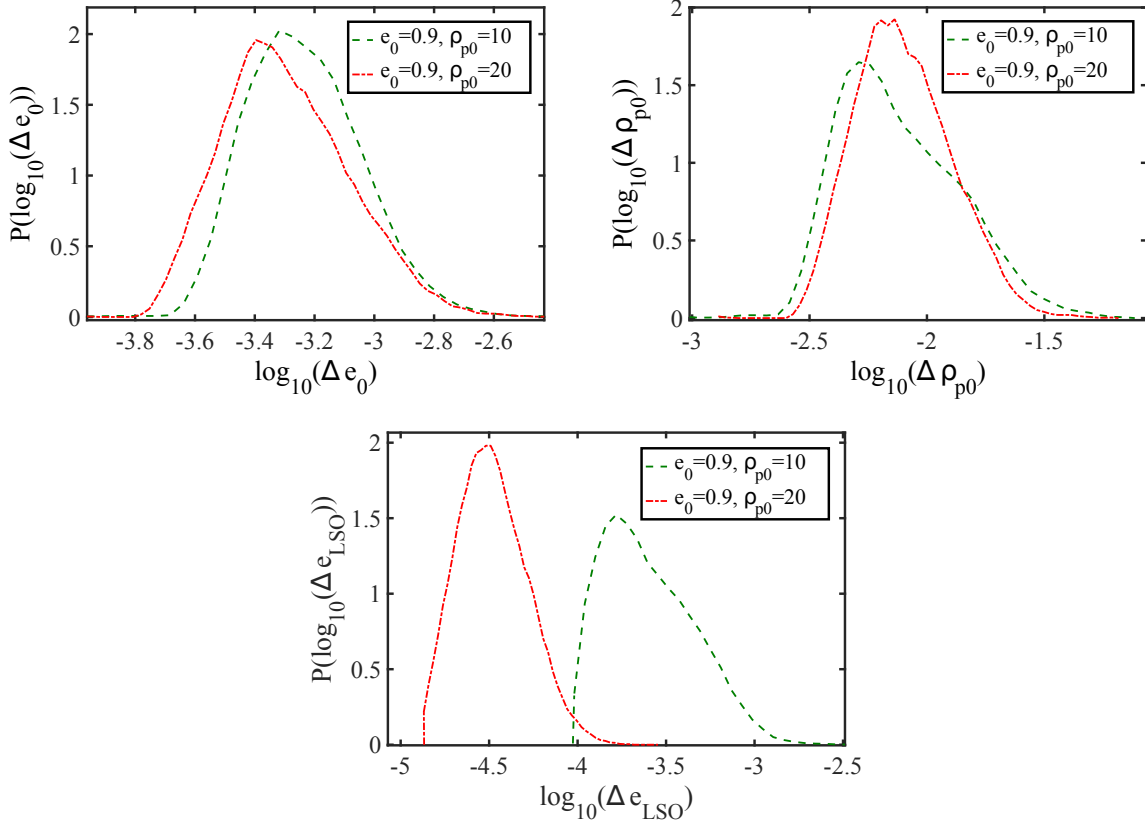
Figure 4. Smoothed distribution of the measurement errors of various parameters measured only for precessing eccentric binaries. Similar to Figure 3, these distributions correspond to a Monte Carlo sample of 4500 binaries with random source direction and binary orientation for $30\,M_\odot - 30\,M_\odot$ precessing eccentric BH binaries with initial eccentricity $e_0 = 0.9$ and dimensionless pericenter distance $\rho_{p0} = 10$ and 20 (green dashed and red dash-dot). The source distance is also fixed at $D_L = 100\,\text{Mpc}$, and the detector network is specified in Table 1. *Top left:* Distribution of the measurement error in initial orbital eccentricity, e_0. *Top right:* Distribution of the measurement error in initial dimensionless pericenter distance, ρ_{p0}. *Bottom:* Distribution of the measurement error in the eccentricity at the last stable orbits, e_{LSO}.

minimum frequency. The peak of the spectrum is initially at $f_{\min}$ if $\rho_{p0} = 40.9[M_{\text{tot}}/(20\,M_\odot)]^{-2/3}(f_{\min}/10\,\text{Hz})^{-2/3}$ (Gondán et al. 2017). A slightly lower value of $f_{\min} \sim 7\,\text{Hz}$ leads to values that represent the minimum of the fast parameters. This also leads to the result observed in Figure 3 that these parameters have higher errors for $\rho_{p0} = 10$ than for $\rho_{p0} = 20$.

In Figure 5, note that Δe_{LSO} errors are relatively small for relatively high ρ_{p0} up to $\rho_{p0} \sim 200$. At high ρ_{p0}, the orbital eccentricity approaches zero when it enters the aLIGO band, and Δe_{LSO} increases. We note that the posterior probability distribution function of e_{LSO} is well-defined even in the circular limit $\rho_{p0} \to \infty$, and Δe_{LSO} is finite for a given confidence region. However, the Fisher matrix algorithm becomes invalid in this regime as the signal is not approximated well by its linear Taylor expansion with respect to the Δe_{LSO} parameter, since its first e_{LSO} derivative vanishes in the circular limit. Therefore the true asymptotic value of Δe_{LSO} for high ρ_{p0} cannot be recovered with the Fisher matrix technique used in this paper. Further, note that Δe_0 and $\Delta \rho_{p0}$ also increase rapidly with ρ_{p0} for high ρ_{p0}. This is due to the fact that for these parameters the binary forms with a pericenter frequency smaller than the minimum frequency of the detector network, and the information on e_0 and ρ_{p0} is limited to higher harmonics with small power. Thus, these parameters indeed have a very high error and become indeterminate in the circular limit. The fact that the relative error of e_0 and ρ_{p0} can be less than $\sim 5\%$ percent in the range $5 < \rho_{p0} < 50$ ($6 < \rho_{p0} < 100$)

for $30\,M_\odot - 30\,M_\odot$ ($10\,M_\odot - 10\,M_\odot$) precessing highly eccentric BH binaries implies that the GW detections might have the potential to constrain the formation environment of these system (O'Leary et al. 2009; Cholis et al. 2016; Rodriguez et al. 2016a,b; Chatterjee et al. 2017; Gondán et al. 2017; Kocsis et al. 2017; Samsing & Ramirez-Ruiz 2017; Silsbee & Tremaine 2017).

Furthermore, we found from numerical investigations that Δe_0 does not correlate significantly with other parameters' errors, which is due to the fact that e_0 is measured from the truncation of the signal for $e > e_0$ at the start of the waveform, while other parameters of a precessing eccentric binary are measured from the inspiral rate (Section 5). However, $\Delta \rho_{p0}$ behaves differently from Δe_0 in this regard, which is due to the fact that ρ_{p0} is determined by e_{LSO} in Equation (50), and e_{LSO} depends on the mass parameters.

6.3. *Comparison with previous results*

In this paper, we have determined the SNR and the expected accuracy with which the aLIGO-AdV-KAGRA detector network may determine the parameters that describe highly eccentric BH binaries, and investigated how these quantities depend on the initial pericenter distance ρ_{p0} and initial eccentricity e_0. There are some previous studies that also made similar investigations for eccentric compact binaries with significant differences (Yunes et al. 2009; Kyutoku & Seto 2014; Sun et al. 2015; Ma et al. 2017). They considered different
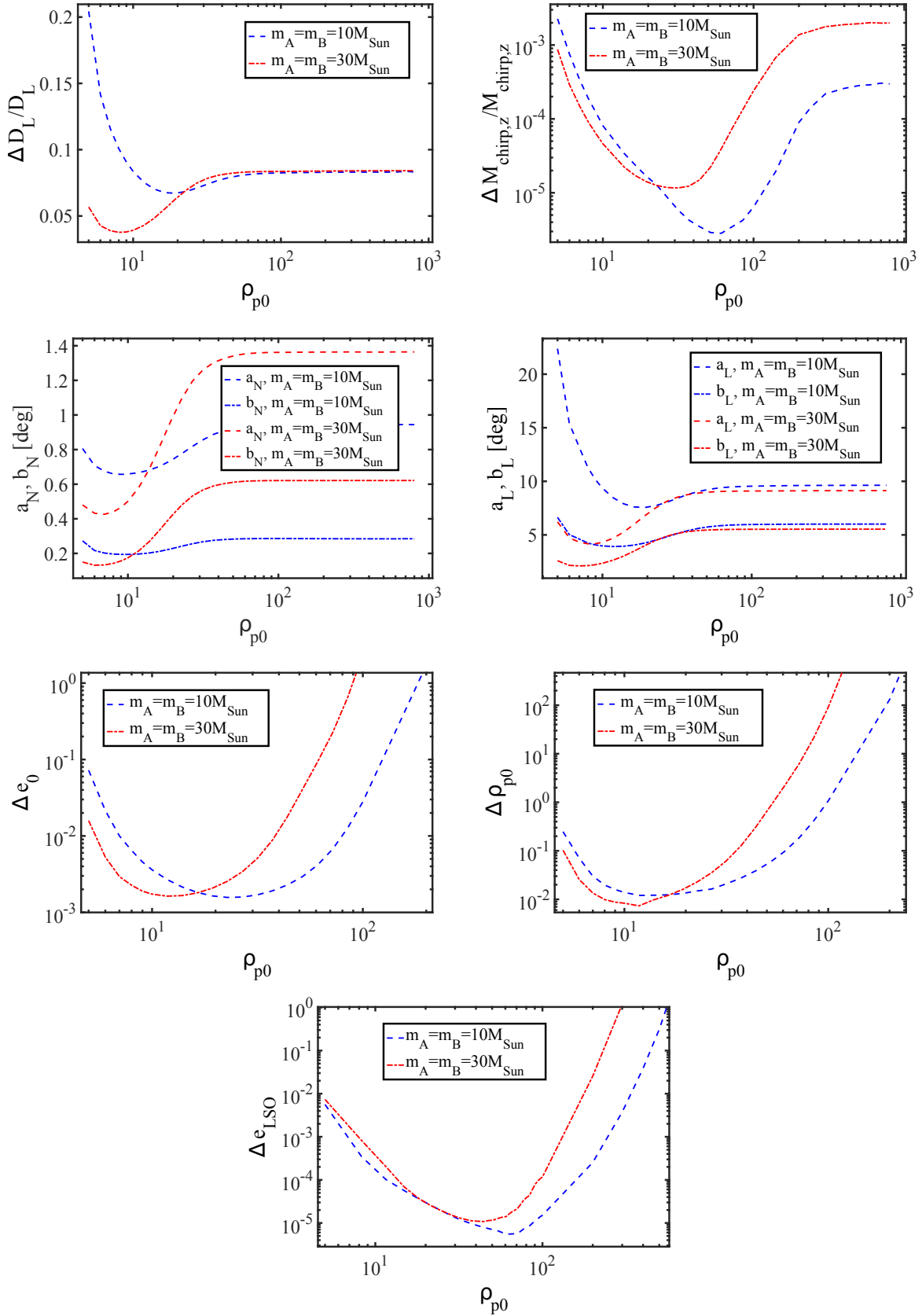
Figure 5. The measurement error of source parameters as a function of initial dimensionless pericenter distance with all other binary parameters fixed as in Figure 2. *First row left:* luminosity distance $\Delta D_L/D_L = \Delta(\ln D_L)$. *First row right:* redshifted chirp mass, $\Delta \mathcal{M}_z/\mathcal{M}_z = \Delta(\ln \mathcal{M}_z)$. *Second row left:* Semi-major and semi-minor axes of the sky position error ellipse a_N and b_N. *Second row right:* Semi-major and semi-minor axes of the error ellipse for the binary orbital plane normal vector direction, a_L and b_L. *Third row left:* Initial orbital eccentricity, Δe_0. *Third row right:* Initial dimensionless pericenter distance, $\Delta \rho_{p0}$. *Fourth row:* Eccentricity at the last stable orbit, Δe_{LSO}. The measurement error of parameters $(\ln(D_L), \ln(\mathcal{M}_z), a_N, b_N, a_L, b_L)$ converge asymptotically to the value of precessing eccentric binaries in the circular limit for high ρ_{p0}, and the measurement error of parameters $(\Delta e_0, \Delta \rho_{p0}, \Delta e_{LSO})$ increase rapidly with ρ_{p0} for high ρ_{p0}. We find similar trends with ρ_{p0} for other random choices of binary direction and orientation (not shown). Note that the measurement error of Δe_{LSO} is undetermined for high ρ_{p0} because the Fisher matrix algorithm becomes invalid for this parameter in this regime, see Section 6.2 for details.
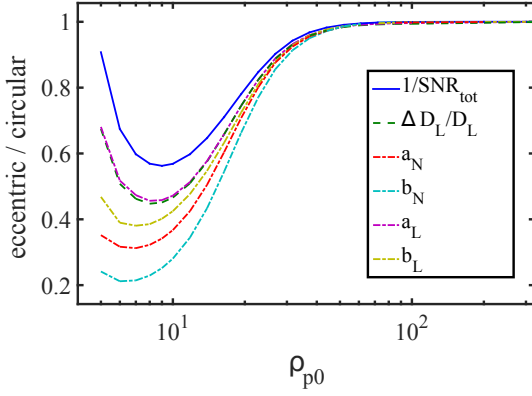
Figure 6. The initial dimensionless pericenter distance dependence of $\Delta D_L/D_L$, the principal axes of the sky localization and binary orientation error ellipses (a_N, b_N) and (a_L, b_L), and $1/\mathrm{SNR_{tot}}$ for $30\,\mathrm{M_\odot} - 30\,\mathrm{M_\odot}$ precessing highly eccentric BH binaries with all other binary parameters fixed as in Figure 2. The values are with respect to those of circular binaries and match those in Figure 5. Note that we find similar trends with ρ_{p0} and e_0 for other random choices of binary direction and orientation (not shown).

detector networks, applied different waveform models, and used different definitions for e_0 and ρ_{p0}. As a consequence, only a qualitative comparison is possible with those results, which we discuss in this section. At the end of this section, we compare our results for the measurement errors in the circular limit with those presented in previous studies.

We first compare our results with a previous study for the ρ_{p0} dependence of the $\mathrm{SNR_{tot}}$. Our result for the ρ_{p0} dependence of the $\mathrm{SNR_{tot}}$ (Figure 2) is qualitatively in agreement with the result of Figure 2 in Kyutoku & Seto (2014), i.e. the $\mathrm{SNR_{tot}}$ increases rapidly with ρ_{p0} for low ρ_{p0}, peaks at a moderate ρ_{p0}, and converges asymptotically to the value of highly eccentric binaries in the circular limit for high ρ_{p0}.

In order to compare our results for the e_0 dependence of the $\mathrm{SNR_{tot}}$ with Yunes et al. (2009) and Sun et al. (2015), we also set the lower bound of advanced GW detectors' sensitive frequency band to $f_{\min} = 20\,\mathrm{Hz}$. We define $e_{20\mathrm{Hz}}$ to be the eccentricity at which the peak GW frequency of the binary defined in Wen (2003) is $f_{\mathrm{GW}} = 20\,\mathrm{Hz}$, and evaluate ρ_{p0} corresponding to $e_0 = e_{20\mathrm{Hz}}$ from Equation (37) in Wen (2003) as

$$\rho_{20\mathrm{Hz}} = \left[(1 + e_{20\mathrm{Hz}})^{0.3046} f_{\mathrm{GW}} \pi M_{\mathrm{tot},z}\right]^{-2/3} . \quad (53)$$

We recalculate the distribution of the $\mathrm{SNR_{tot}}$ for $10\,\mathrm{M_\odot} - 10\,\mathrm{M_\odot}$ precessing eccentric compact binaries with $e_{20\mathrm{Hz}} = (0.1, 0.2, 0.3, 0.4)$. The top panel of Figure 7 shows that the $\mathrm{SNR_{tot}}$ is roughly the same for different $e_{20\mathrm{Hz}}$, which is consistent with results presented in Figure 2 in Sun et al. (2015) and in the left panel of Figure 8 in Yunes et al. (2009). Moreover, we find that the $\mathrm{SNR_{tot}}$ increases weakly with $e_{20\mathrm{Hz}}$, which is in agreement with results in the left panel of Figure 8 in Yunes et al. (2009). Note that this result disagrees with Table 5 in Sun et al. (2015). The $\mathrm{SNR_{tot}}$ does not depend significantly on $e_{20\mathrm{Hz}}$ in the range of $[0.1, 0.4]$ for $10\,\mathrm{M_\odot} - 10\,\mathrm{M_\odot}$ for $\rho_{p0} = \rho_{20\mathrm{Hz}} \sim 28$ as seen in the bottom panel of Figure 7.[13] For $10\,\mathrm{M_\odot} - 10\,\mathrm{M_\odot}$ binaries with $e_0 < 0.4$ and $\rho_{p0} \sim 28$, we find that $\rho_{20\mathrm{Hz}}$ is high enough to fall into the range of ρ_{p0} where the

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[13] For binaries with relatively high ρ_{p0}, binaries are well-circularized by the time their peak GW frequency enters the sensitive frequency band of advanced ground-based GW detectors, thus the information about the initial eccentricity vanishes from the detectable part of the waveform. This explains the very weak e_0 dependence of the $\mathrm{SNR_{tot}}$ and of parameter measurement errors for high ρ_{p0} in the bottom panel of Figure 7 and in Figure 8.
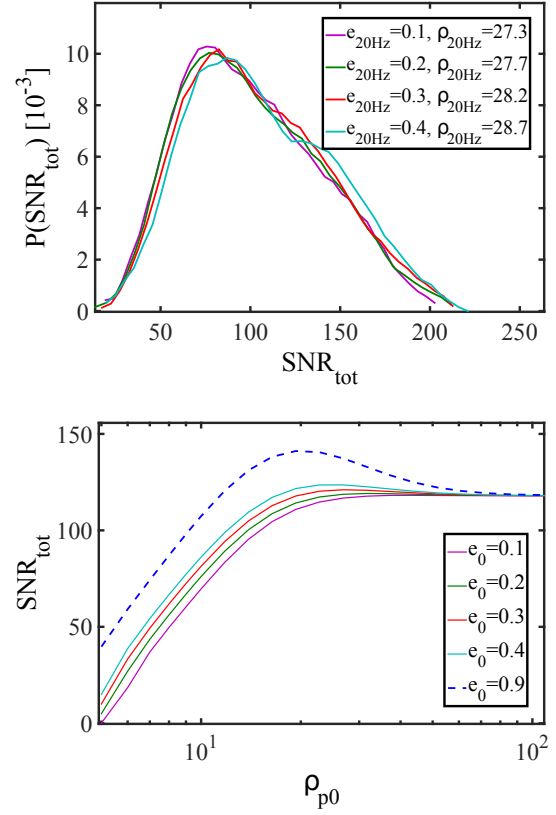
Figure 7. *Top panel*: Smoothed probability density function of the $\mathrm{SNR_{tot}}$ for $10\,\mathrm{M_\odot} - 10\,\mathrm{M_\odot}$ precessing eccentric BH binaries. In these calculations we set the lower bound of the detectors' sensitive frequency band to $20\,\mathrm{Hz}$. Here $e_{20\mathrm{Hz}}$ and $\rho_{20\mathrm{Hz}}$ represent the initial orbital eccentricity and initial dimensionless pericenter distance at which the peak GW frequency (Wen 2003) of the binary is $f_{\mathrm{GW}} = 20\,\mathrm{Hz}$. Other details of the calculations are the same as in Figure 1. *Bottom panel:* The same as in Figure 2 but for $10\,\mathrm{M_\odot} - 10\,\mathrm{M_\odot}$ precessing eccentric BH binaries with initial eccentricities $e_0 = (0.1, 0.2, 0.3, 0.4, 0.9)$ as a function of ρ_{p0}. We also represent the $e_0 = 0.9$ curve in this plot in order to show that $\mathrm{SNR_{tot}}$ is a strictly monotonically increasing function of e_0 over the full range of ρ_{p0}. We find similar trends with ρ_{p0} and e_0 for other random choices of binary direction and orientation (not shown).

$\mathrm{SNR_{tot}}$ depends on e_0 at the $\sim 10\%$ level for $e_{20\mathrm{Hz}} < 0.4$. The influence of $e_{20\mathrm{Hz}}$ on the $\mathrm{SNR_{tot}}$ increases with M_{tot} since in this case $\rho_{20\mathrm{Hz}}$ is lower as shown by Equation (53). Thus, the influence of $e_{20\mathrm{Hz}}$ on the distribution of $\mathrm{SNR_{tot}}$ is more significant for higher-mass low-eccentricity binaries. Examples for this characteristic of the $\mathrm{SNR_{tot}}$ are seen in Figure 8 in Yunes et al. (2009).

Finally, we compare our results with previous studies for the e_0 dependence of measurement errors of parameters describing eccentric binaries. Sun et al. 2015; Ma et al. 2017 set the initial orbital parameters to be $e_{20\mathrm{Hz}}$ ($e_{10\mathrm{Hz}}$) and $\rho_{20\mathrm{Hz}}$ ($\rho_{10\mathrm{Hz}}$). For various values of $e_{20\mathrm{Hz}}$ ($e_{10\mathrm{Hz}}$) in the range $[0.1, 0.2, 0.3, 0.4]$ and the corresponding values of $\rho_{20\mathrm{Hz}}$ ($\rho_{10\mathrm{Hz}}$), they determined the measurement accuracies for various parameters of eccentric binaries. Since they applied different waveform models and different parameters describing the eccentric binaries, we resort to a qualitative comparison. We repeated the analysis of Figure 5 for $e_0 = (0.1, 0.2, 0.3, 0.4)$ and determined the measurement error of parameters as a function of ρ_{p0} as shown in Figure 8. We find qualitative agreement with Sun et al. (2015); Ma et al. (2017). The measurement accuracies of parameters increase strictly monotonically with e_0.
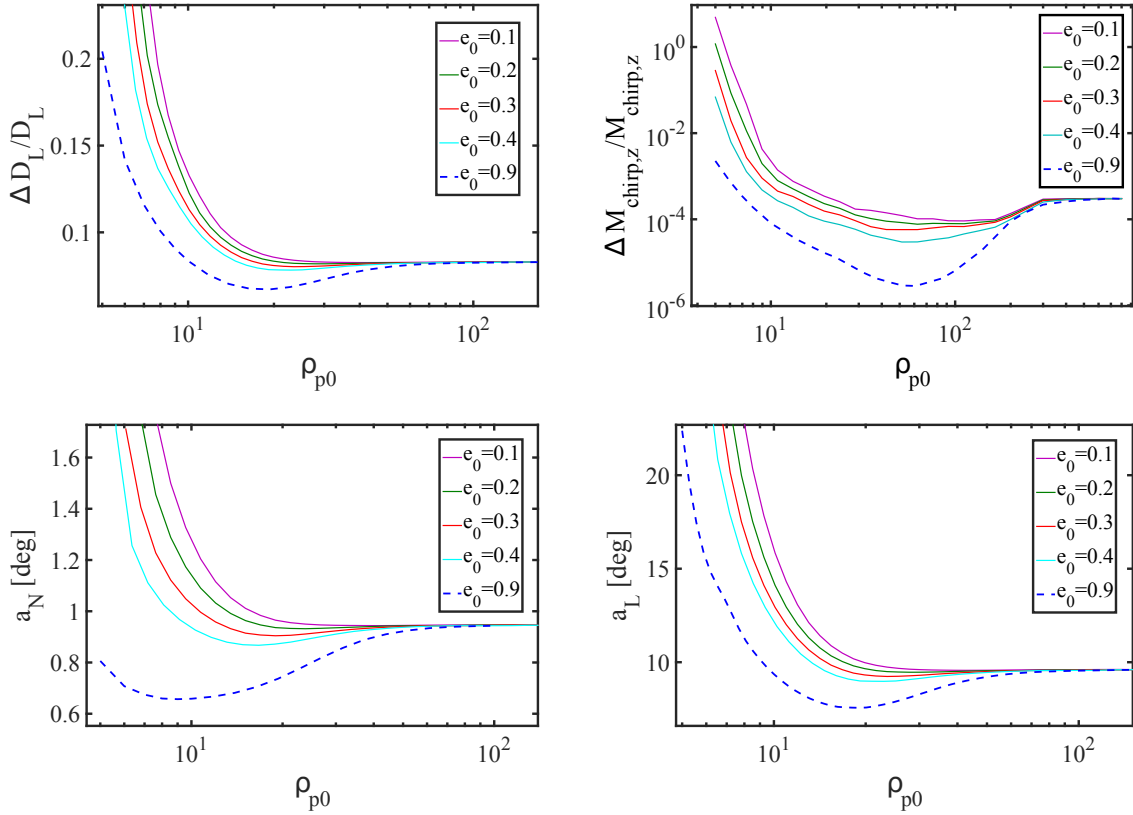
Figure 8. The initial pericenter distance and eccentricity dependence of the measurement errors of various parameters for $10\,\mathrm{M}_\odot - 10\,\mathrm{M}_\odot$ precessing eccentric BH binaries for a specific sky position and inclination chosen as in Figure 2. Different lines show different initial eccentricities as labeled. *Top left:* luminosity distance $\Delta D_\mathrm{L}/D_\mathrm{L} = \Delta(\ln D_\mathrm{L})$. *Top right:* redshifted chirp mass, $\Delta\mathcal{M}_z/\mathcal{M}_z = \Delta(\ln\mathcal{M}_z)$. *Bottom left:* Semi-major axis of the sky position error ellipse, a_N. *Bottom right:* Semi-major axis of the error ellipse for the binary orbital plane normal vector direction, a_L. We find similar trends with ρ_{p0} and e_0 for other random choices of binary direction and orientation (not shown). We also show the $e_0 = 0.9$ curve in this plot to illustrate that measurement errors of binary parameters are a strictly monotonically increasing functions of e_0 over the full range of ρ_{p0}.

Previous papers have investigated the M_tot dependency of the measurement errors for $\{t_c, \Phi_c, \ln(\mathcal{M}_z), \ln(\eta)\}$ by using different PN order waveform models for non-spinning inspiraling binaries for a fixed SNR in a single aLIGO type detector. Previous results showed that the measurement accuracy of these parameters decreases with increasing M_tot for $2.8\,\mathrm{M}_\odot \leq M_\mathrm{tot} \leq 20\,\mathrm{M}_\odot$ provided that the SNR accumulated in one GW detector is fixed, see Table 1 in Arun et al. (2005) and references therein. Therefore, we determined the measurement errors in the circular limit for $\{t_c, \Phi_c, \ln(\mathcal{M}_z), \ln(\eta)\}$ for a qualitative comparison.[14] To calculate the measurement error of the $\ln(\eta)$ parameter we use the fact that $\eta = (\mathcal{M}_z M_{\mathrm{tot},z}^{-1})^{5/3}$ and so

$$\frac{\langle\Delta\eta^2\rangle}{\eta^2} = \frac{25}{9}\frac{\langle\Delta\mathcal{M}_z^2\rangle}{\mathcal{M}_z^2} + \frac{25}{9}\frac{\langle\Delta M_{\mathrm{tot},z}^2\rangle}{M_{\mathrm{tot},z}^2} + \frac{50}{9}\frac{\langle\Delta\mathcal{M}_z\Delta M_{\mathrm{tot},z}\rangle}{\mathcal{M}_z M_{\mathrm{tot},z}}. \tag{54}$$

In agreement with the 1PN order case in Arun et al. (2005), we find that Δt_c, $\Delta\Phi_c$, $\Delta\mathcal{M}_z/\mathcal{M}_z$, and $\Delta\eta/\eta$ increase with M_tot for fixed $\mathrm{SNR}_\mathrm{tot}$ (Table 4). Such a qualitative agreement is expected since the adopted precessing eccentric waveform approximates the full 1PN waveform in its most important features, and the M_tot-dependent trends of error distributions do not depend on the number of detectors or on the sky position or angular momentum unit vectors of the source.

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[14] A quantitative agreement is not expected since our precessing waveform model differs from the waveform models in those studies.

$m_A - m_B$	Δt_c	$\Delta\Phi_c$	$\Delta\mathcal{M}_z/\mathcal{M}_z$	$\Delta\eta/\eta$
$10\,\mathrm{M}_\odot - 10\,\mathrm{M}_\odot$	0.30	0.32	3.6×10^{-4}	3.4×10^{-3}
$15\,\mathrm{M}_\odot - 15\,\mathrm{M}_\odot$	0.55	0.39	8.7×10^{-4}	6.5×10^{-3}
$20\,\mathrm{M}_\odot - 20\,\mathrm{M}_\odot$	0.94	0.49	1.5×10^{-3}	10^{-2}
$25\,\mathrm{M}_\odot - 25\,\mathrm{M}_\odot$	1.38	0.57	2.2×10^{-3}	1.3×10^{-2}
$30\,\mathrm{M}_\odot - 30\,\mathrm{M}_\odot$	1.84	0.62	2.8×10^{-3}	1.5×10^{-2}

Table 4
Errors in t_c (msec), Φ_c (rad) and for the relative errors in $\mathcal{M}_z$ and η in the circular limit for equal-mass binaries for a specific sky position and inclination $\theta_N = \pi/2$, $\phi_N = \pi/3$, $\theta_L = \pi/4$, and $\phi_L = \pi/5$ in each case, we have assumed detection with the detector network introduced in Table 1, and errors correspond to a fixed $\mathrm{SNR}_\mathrm{tot} = 100$. We find similar trends for other random choices of binary directions and orientations (not shown).

7. SUMMARY AND CONCLUSION

We carried out a Fisher-matrix-type study to determine the accuracy with which the parameters of highly eccentric BH binaries may be measured using the aLIGO-AdV-KAGRA GW detector network. Eccentricity changes the GWs of binaries compared to circular binaries, in several ways. In time-domain, the gravitational waveform of eccentric binaries is quasiperiodic but not sinusoidal. Relativistic precession adds a slow amplitude modulation to the waveform for each polarization. Eccentricity also changes the inspiral rate at which the binary separation and period shrink. We take all of these effects into account using the stationary phase approximation (Moreno-Garrido et al. 1994; Mikóczi et al. 2012). In con-

trast to circular binaries, the waveform of eccentric binaries includes several prominent orbital frequency harmonics, general relativistic precession causes each harmonic to split into three frequencies for both GW polarizations respectively, and the eccentric inspiral creates a spectrum, which is different from the $\tilde{h} \propto f^{-7/6}$ waveform of circular inspiral sources for each harmonic. These features in the waveform make it possible to accurately determine the eccentricity and angle of periapsis, and the modulated inspiral rate improves the measurement accuracy of mass parameters for eccentric inspirals.

The main parameters that describe eccentric inspiraling binaries are the initial pericenter distance ρ_{p0} when the eccentricity is close to unity and the final eccentricity at the last stable orbit e_{LSO}. These parameters are systematically different for different formation channels. Thus their measurement may have important implications on the astrophysical origin of the sources (O'Leary et al. 2009; Cholis et al. 2016; Rodriguez et al. 2016a,b; Chatterjee et al. 2017; Gondán et al. 2017; Samsing & Ramirez-Ruiz 2017; Silsbee & Tremaine 2017). Based on a survey with $10\,M_\odot - 10\,M_\odot$ and $30\,M_\odot - 30\,M_\odot$ precessing highly eccentric BH binaries at $100\,\mathrm{Mpc}$ using the planned aLIGO-AdV-KAGRA detector network, our results are be summarized as follows.

1. The $\mathrm{SNR_{tot}}$ improves by a factor of $\sim 1 - 1.7$ (depending on ρ_{p0}, the component masses, and the sky position and binary orientation angles, see Figure 1) for $30\,M_\odot - 30\,M_\odot$ precessing highly eccentric BH binaries compared to similar binaries in the circular limit [15] with the same masses and distance. The volume in the Universe for a fixed maximum S/N is $\langle (S/N)^3 \rangle \sim 5\times$ ($2\times$) larger for eccentric inspiraling binaries with $\rho_{p0} = 10$ ($\rho_{p0} = 20$) than for similar binaries in the circular limit.

2. We determined how the parameters' measurement accuracies depend on the initial dimensionless pericenter distance (ρ_{p0}) for precessing highly eccentric BH binaries. The smallest errors are obtained for small $\rho_{p0} < 10$ values for the sky position and angular momentum and $\rho_{p0} < 20$ for the luminosity distance and ρ_{p0}. However, the errors for fast parameters, which are sensitive to the GW phase, like the chirp mass, the initial eccentricity, and the eccentricity at the last stable orbit improve most significantly for a higher ρ_{p0} between 10 and 80 (Figure 5).

3. The parameter estimation errors can improve significantly for highly eccentric precessing BH binaries compared to similar binaries in the circular limit by a factor of (depending on ρ_{p0}, the component masses, and the sky position and binary orientation angles, see Figures 3 and 5)

   - $\sim 1 - 200$ for the mass errors,
   - $\sim 1 - 4.5$ for the semi-major and semi-minor axes of the sky localization ellipse,
   - $\sim 1 - 2$ for the distance errors,
   - $\sim 1 - 3$ for the semi-major and semi-minor axes of the error ellipse for the binary orientation.

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[15] We adopted the leading order stationary phase approximation waveform for circular sources.

4. For initially highly eccentric BH binaries at $D_L = 100\,\mathrm{Mpc}$, the measurement errors for parameters specific to precessing highly eccentric BH binary sources may be as low as of order (depending on ρ_{p0}, the component masses, and the sky position and binary orientation angles, see Figures 4 and 5)

   - 10^{-5} for the final eccentricity errors at LSO,
   - 10^{-4} for the initial eccentricity errors,
   - 10^{-3} for the initial pericenter distance.

For initially moderately eccentric to low eccentricity binaries, the parameter measurement errors and SNRs improve by a smaller amount for low to intermediate ρ_{p0} (Figure 8).

Note that the eccentricity errors are remarkably low, which is not surprising given that eccentricity is encoded in several measurable features of the waveform including the orbital harmonics, the splitting of each frequency harmonic into triplets, the frequency evolution of harmonics (the "chirp"), the frequency evolution of the distance between the spectral triplets, the low frequency cutoff of the signal at the initial pericenter frequency, and the eccentricity dependence of the last stable orbit where the inspiral transitions into a rapid coalescence.

However, there are several factors which may significantly increase the measurement errors in more typical cases. First, more typical sources are expected to be at much larger distances than $100\,\mathrm{Mpc}$. Assuming crudely that the eccentricity errors scale with D_L, the median measurement errors for a $30\,M_\odot - 30\,M_\odot$ precessing highly eccentric BH binary at $\sim 410\,\mathrm{Mpc}$ (similar to GW150914) are expected to be $\Delta e_{LSO} \sim 8.8 \times 10^{-4}$ (1.3×10^{-4}) for the final eccentricity at the last stable orbit if $\rho_{p0} = 10$ (20) for the design sensitivity of the aLIGO/AdV/KAGRA instruments. In these cases, the expected median initial eccentricity error is $\Delta e_0 \sim 8.9 \times 10^{-3}$ (1.2×10^{-2}), and the median initial pericenter distance is $\Delta \rho_{p0} \sim 4.4 \times 10^{-2}$ (9.4×10^{-2}; see Table 2).

Another important simplifying assumption, which may have skewed the errors to lower values, was to neglect higher order post-Newtonian (PN) corrections which depend on the spin of the merging objects. The spins of the two binary components introduce 6 additional parameters, which may become partially degenerate with all other parameters thereby increasing their errors. On the other hand, spin precession breaks degeneracies between the binary orientation and other slow parameters (Lang & Hughes 2006; Kocsis et al. 2007; Chatziioannou et al. 2014). However, the eccentricity-induced orbital harmonics enter at the Newtonian order, the frequency-triplets due to GR precession enter at the low 1PN order. Therefore, the eccentricity-related spectral features are dominant already at the early stages of the inspiral when higher order PN corrections are negligible. For this reason, the estimated Δe_0 and $\Delta \rho_{p0}$ errors are expected to be robust. On the other hand, most of the SNR accumulates at late times for stellar-mass BH binaries where the high order PN corrections are significant. We leave an estimate of parameter estimation errors for spinning binaries to future work.

Finally, an important simplification is the Fisher matrix method itself, which is valid only if the waveform model is a faithful representation of the GW signal and the noise is Gaussian and sufficiently small that the SNR is sufficiently large that the parameter error region is an ellipsoid in parameter space and when the parameter derivate of the waveform is non-vanishing. For smaller SNR, the parameter error region geometry is more complex and the uncertainties are

generally higher (see Cornish & Littenberg 2015, and reference therein). The parameter estimation errors may also be affected by theoretical uncertainties of the waveform model (Cutler & Vallisneri 2007), which may be especially important for highly eccentric binaries, where the post-Newtonian expansion is slowly convergent (Kocsis & Levin 2012).

Such low eccentricity errors may give the aLIGO-AdV-KAGRA GW detector network the capability to distinguish among different astrophysical formation channels. In a companion study (Gondán et al. 2017), we illustrate the expected distribution of eccentricities and other physical parameters for single-single GW capture sources in galactic nuclei. Similar studies for other astrophysical formation channels are underway.

Future multi-waveband searches of eccentric inspiral sources with *LISA* and aLIGO-AdV-KAGRA (Kocsis & Levin 2012; Sesana 2016) have good prospects for even more accurate measurements of the physical parameters well beyond the level reported here. In that case the GW frequency range is much wider. Since eccentricity decreases due to GW emission, eccentricity may be expected to be much higher at lower frequencies in the *LISA* band. This leads to a much larger total GW phase shift caused by eccentricity. The better measurement of relativistic precession may more efficiently break degeneracies between mass and other parameters. The modulation caused by the orbit of the instrument around the Sun and Earth's spin can help break degeneracies among source direction, orientation, and other parameters. Accounting for eccentricity for third generation Earth based (e.g. Einstein telescope), deci-Hertz to mHz space-based instruments will be essential (Chen & Amaro-Seoane 2017).

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APPENDIX

A. THE CIRCULAR LIMIT

The circular limit follows from the limit $e_0 \to 0$ for arbitrary ρ_{p0} or $\rho_{p0} \to \infty$ for arbitrary $e_0 \leq 1$ (see Figure 5 in O'Leary et al. (2009)). In practice we set $e_0 = 10^{-4}$, $\rho_{p0} = 1000$, and omit the parameters from the Fisher matrix that become degenerate or unconstrained in the circular limit: γ_c, e_0, and e_{LSO}. Thus, the parameters in the circular limit are

$$\boldsymbol{\lambda}_{\text{Prec,circ}} = \{t_c, \Phi_c, \ln(D_L), \ln(\mathcal{M}_z), \ln(M_{\text{tot},z}), \theta_N, \phi_N, \theta_L, \phi_L\}, \tag{A1}$$

where $M_{\text{tot},z}$ arises due to precession even in the circular limit (see Equation (25)). In the $\rho_{p0} \to \infty$ limit, $\Delta\gamma_0 \to \infty$ and $\Delta e_0 \to \infty$, however the Fisher matrix algorithm becomes invalid for Δe_{LSO} in this limit (Section 6.2).

We have also examined how the measurement errors depend on the total mass of the binary in the circular limit, and qualitatively compared our results to those of previous parameter estimation studies in Section 6.3. We have found an excellent match between our results and the results of previous studies.

B. RESPONSE OF AN INDIVIDUAL GROUND-BASED DETECTOR

Here we describe the measured signal of individual ground-based detectors, and introduce the adopted coordinate system.

We define the Cartesian coordinate system with basis vectors **i**, **j**, **k** and a spherical coordinate system (θ, ϕ) fixed relative to the center of the Earth, such that **k** and $\theta = 0$ is along the North geographic pole and $\phi = 0$ is along the prime meridian. We denote the unit vector pointing from the center of Earth to the binary's sky position as **N**, and define **L** to be the normal vector parallel with the binary's orbital angular momentum,

$$\mathbf{N} = \sin\theta_N \cos\phi_N \mathbf{i} + \sin\theta_N \sin\phi_N \mathbf{j} + \cos\theta_N \mathbf{k}, \tag{B1}$$

$$\mathbf{L} = \sin\theta_L \cos\phi_L \mathbf{i} + \sin\theta_L \sin\phi_L \mathbf{j} + \cos\theta_L \mathbf{k}. \tag{B2}$$

We denote the unit vectors parallel with the arms of the k-th detector as $\mathbf{x}_k$ and $\mathbf{y}_k$, and set $\mathbf{z}_k = \mathbf{x}_k \times \mathbf{y}_k$. As $\mathbf{x}_k$ and $\mathbf{y}_k$ are parallel with the surface of Earth for all detectors, $\mathbf{z}_k$ points from the center of Earth toward the geographical location of the kth detector. Let the coordinates (θ_k, ϕ_k) denote the location of the kth detector, thus the unit vectors along the arms can be expressed as

$$\begin{aligned} \mathbf{x}_k = &(\cos\psi_k \sin\phi_k - \sin\psi_k \cos\phi_k \cos\theta_k)\mathbf{i} \\ &+ (-\cos\psi_k \cos\phi_k - \sin\psi_k \sin\phi_k \cos\theta_k)\mathbf{j} \\ &+ (\sin\psi_k \sin\theta_k)\mathbf{k}, \end{aligned} \tag{B3}$$

$$\begin{aligned} \mathbf{y}_k = &(-\sin\psi_k \sin\phi_k - \cos\psi_k \cos\phi_k \cos\theta_k)\mathbf{i} \\ &+ (\sin\psi_k \cos\phi_k - \cos\psi_k \sin\phi_k \cos\theta_k)\mathbf{j} \\ &+ (\cos\psi_k \sin\theta_k)\mathbf{k}, \end{aligned} \tag{B4}$$

$$\mathbf{z}_k = \sin\theta_k \cos\phi_k \mathbf{i} + \sin\theta_k \sin\phi_k \mathbf{j} + \cos\theta_k \mathbf{k} \tag{B5}$$

(Creighton & Anderson 2011), where the orientation angle of the kth detector, ψ_k, is defined in Section 3.

These vectors define the response tensor for the kth detector:

$$D_k^{ij} = \frac{1}{2}\left(x_k^i x_k^j - y_k^i y_k^j\right) \tag{B6}$$

(Finn & Chernoff 1993), where x_k^i and y_k^i are the ith Cartesian components of $\mathbf{x}_k$ and $\mathbf{y}_k$.

We adopt the basis vectors following the conventions of previous studies (Finn & Chernoff 1993; Cutler & Flanagan 1994; Anderson et al. 2001; Dalal et al. 2006; Nissanke et al. 2010)

$$\mathbf{X} = \frac{\mathbf{N} \times \mathbf{L}}{|\mathbf{N} \times \mathbf{L}|}, \quad \mathbf{Y} = \frac{\mathbf{X} \times \mathbf{N}}{|\mathbf{X} \times \mathbf{N}|} \tag{B7}$$

with preferred polarization basis tensors

$$e_{ij}^+ = X_i X_j - Y_i Y_j, \tag{B8}$$

$$e_{ij}^\times = X_i Y_j + Y_i X_j, \tag{B9}$$

where i and j are Cartesian components. Thus, the transverse-traceless metric perturbation describing the GW is written as

$$h_{ij} = h_+ e_{ij}^+ + h_\times e_{ij}^\times, \tag{B10}$$

where h_+ and $h_\times$ are given in Equations (1) and (2).

The response of the kth detector to a GW with frequency f can be given in time-domain by

$$h_k = e^{i\Delta\Phi_k} D_k^{ij} h_{ij} = e^{i\Delta\Phi_k} \left(h_+ F_{+,k} + h_\times F_{\times,k} \right) \qquad \text{(B11)}$$

(Nissanke et al. 2010), where $\mathbf{r}_k = R_\oplus \mathbf{z}_k$ is the position of the kth detector, the factor $-\mathbf{N} \cdot \mathbf{r}_k$ measures the light travel time between the kth detector and the coordinate origin, thus the factor $\Delta\Phi_k = -2\pi f \mathbf{N} \cdot \mathbf{r}_k$ measures the phase shift between the kth detector and the coordinate origin. In Equation (B11), $F_{+,k}$ and $F_{\times,k}$ are the antenna factors

$$F_{+,k} = e_{ij}^+ D_k^{ij}, \quad F_{\times,k} = e_{ij}^\times D_k^{ij}. \qquad \text{(B12)}$$

In our calculations, Earth is taken to be a sphere with radius of $R_\oplus = 6,370\,\text{km}$.

If the time that the GW signal spends in the detectors' sensitive frequency band is negligible compared to the rotation period of Earth, then the measured waveform in frequency domain for the kth detector is

$$\tilde{h}_k(f) = \left[F_{+,k}\tilde{h}_+(f) + F_{\times,k}\tilde{h}_\times(f) \right] e^{-2\pi i f \mathbf{N} \cdot \mathbf{r}_k}, \qquad \text{(B13)}$$

where $\tilde{h}_+(f)$ and $\tilde{h}_\times(f)$ are the Fourier-transformed expressions of h_+ and $h_\times$ at Earth's center, and $F_{+,k}$ and $F_{\times,k}$ are given by the (practically time-independent) orientation of the detectors shown in Table 1.

Similarly, using the frequency harmonic triplets for eccentric precessing inspiraling binaries in the stationary phase approximation, the measured waveform for the kth detector is

$$\tilde{h}_k(\mathbf{f}) = F_{+,k}\tilde{h}_{+,k}(\mathbf{f}) + F_{\times,k}\tilde{h}_{\times,k}(\mathbf{f}), \qquad \text{(B14)}$$

where $\tilde{h}_{+,k}(\mathbf{f})$ and $\tilde{h}_{\times,k}(\mathbf{f})$ can be derived from $\tilde{h}_+(\mathbf{f})$ and $\tilde{h}_\times(\mathbf{f})$ by multiplying each term of $\mathbf{f}$ with the phase shift factors $e^{-2\pi i f_n \mathbf{N} \cdot \mathbf{r}_k}$ and $e^{-2\pi i f_n^\pm \mathbf{N} \cdot \mathbf{r}_k}$ for each harmonic, respectively. More specifically, the measured signal's Fourier phase in Equations (35) and (36) in the kth detector is shifted respectively according to

$$\Psi_{n,k} = \Psi_n - 2\pi i f_n \mathbf{N} \cdot \mathbf{r}_k, \qquad \text{(B15)}$$
$$\Psi_{n,k}^\pm = \Psi_n^\pm - 2\pi i f_n^\pm \mathbf{N} \cdot \mathbf{r}_k. \qquad \text{(B16)}$$

C. TIME EVOLUTION OF THE ORBIT

In this section, we derive the time evolution of different harmonics in the detectors' sensitive frequency band, and for each orbital harmonic we determine the eccentricity at which the signal enters the detectors' sensitive frequency band. These formulae will be utilized in Appendix D.

The time-dependent GW signal of a precessing eccentric BH binary as measured by the kth detector can be given as

$$\begin{aligned} h_k(t) = &\, h_+(t) F_{+,k}[\alpha_N(t), \beta_N, \alpha_L, \beta_L] \\ &+ h_\times(t) F_{\times,k}[\alpha_N(t), \beta_N, \alpha_L, \beta_L], \end{aligned} \qquad \text{(C1)}$$

where $h_+(t)$ and $h_\times(t)$ are given in Equations (1) and (2), and $F_{+,k}$ and $F_{\times,k}$ are quantified by Equation (B12). We neglect spins in this study, therefore the angular momentum vector direction (α_L, β_L) is conserved during the eccentric inspiral (Cutler & Flanagan 1994). The polar angle of the source α_N relative to the detector depends on the rotation phase of Earth during the day. We neglect Earth's rotation, since the total duration of an eccentric inspiral from $e = 0.9$ to merger is of

order $[(\rho_{\text{p0}}/40)^4(4\eta)^{-1}M_{\text{tot}}/(20\,\text{M}_\odot)\,\text{min}]$ as shown in Figure 3 in O'Leary et al. (2009). [16]

The waveform of an eccentric binary in the stationary phase approximation is a sum over harmonics n for each component of the frequency triplet (f_n, f_n^+, f_n^-) with different reference times (t_n, t_n^+, t_n^-) (see Section 2.2). During the evolution $e(t)$ (and $\rho_p(t)$) shrinks strictly monotonically in time (Peters 1964), therefore its inverse function $t(e)$ is well-defined and determines t_n, t_n^+, and t_n^-. For inspiraling circular binaries, time t can be expressed using the frequency of the emitted GW signal $f = 2\nu$ as

$$t(f) = t_c - \int_f^\infty \frac{df'}{\dot{f}'} = t_c - 5(8\pi f)^{-8/3}\mathcal{M}^{-5/3} \qquad \text{(C2)}$$

(see equations (2.13) and (2.19) in Cutler & Flanagan (1994) for details), where the constant of integration, t_c, is defined by the requirement that $t \to t_c$ as $f \to \infty$. We generalize Equation (C2) for eccentric inspirals by changing the integration variable from f to e in Equation (C2),

$$t(e) = t_c + \int_0^e \frac{de'}{\dot{e}(e')} = t_c - \tau I_t(e), \qquad \text{(C3)}$$

where $\dot{e}$ is given by Equation (15), and we introduced

$$\begin{aligned} \tau &= \frac{15}{304}\mathcal{M}^{-5/3}(2\pi c_0)^{-8/3} \\ &= \frac{15}{304}\frac{M_{\text{tot}}^{8/3}}{\mathcal{M}^{5/3}}\frac{[(1-e_{\text{LSO}})\rho_{\text{pLSO}}(e_{\text{LSO}})]^4}{H(e_{\text{LSO}})^{8/3}}, \end{aligned} \qquad \text{(C4)}$$

and substituted c_0 using Equation (22). Here $\rho_{\text{pLSO}}(e_{\text{LSO}})$ and $H(e_{\text{LSO}})$ are given by Equations (21) and (19), and $I_t(e)$ in Equation (C3) is of form

$$I_t(e) = \int_0^e \frac{x^{29/19}\left(1 + \frac{121}{304}x^2\right)^{\frac{1181}{2299}}}{(1-x^2)^{3/2}} dx \qquad \text{(C5)}$$

(see Mikóczi et al. 2012, for an analytic result). Using Equation (C3) we get the total duration of the nth harmonic in the detector's sensitive frequency band

$$T_n = t(e_{\text{min,n}}) - t(e_{\text{max,n}}) = [I_1(e_{\text{max,n}}) - I_1(e_{\text{min,n}})]\tau, \qquad \text{(C6)}$$
$$T_n^+ = t(e_{\text{min,n}}^+) - t(e_{\text{max,n}}^+) = [I_1(e_{\text{max,n}}^+) - I_1(e_{\text{min,n}}^+)]\tau, \qquad \text{(C7)}$$
$$T_n^- = t(e_{\text{min,n}}^-) - t(e_{\text{max,n}}^-) = [I_1(e_{\text{max,n}}^-) - I_1(e_{\text{min,n}}^-)]\tau. \qquad \text{(C8)}$$

Here $e_{\text{min,n}}$ ($e_{\text{min,n}}^+, e_{\text{min,n}}^-$) refers to the eccentricity at which the harmonic f_n (f_n^+, f_n^-) reaches the LSO or when it exists the detectable highest frequency for the given detector, and $e_{\text{max,n}}$ ($e_{\text{max,n}}^+, e_{\text{max,n}}^-$) refers to the eccentricity at which the signal related to f_n (f_n^+, f_n^-) first enters the detector's sensitive frequency band or when it forms within the band. Thus,

$$e_{\text{min,n}} = \max(e_{\text{LSO}}, e_{\text{det,n}}^{\text{min}}), \qquad \text{(C9)}$$
$$e_{\text{max,n}} = \min(e_0, e_{\text{det,n}}^{\text{max}}), \qquad \text{(C10)}$$

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[16] Earth's rotation may be relevant for highly eccentric low mass compact objects with large $\rho_{\text{p0}} \gtrsim 40$, such as neutron star binaries. For black holes, if $\rho_{\text{p0}} \gg 40$, then the signal mostly circularizes before it enters the detectors' sensitive frequency band, and the amount of time it spends in the band with a significant SNR_{tot} is limited to less than a minute. For an illustration of the accumulation of the SNR_{tot} with time we refer the reader to Figure 10 of O'Leary et al. (2009) and Figure 7 of Kocsis & Levin (2012).

where

$$e_{\mathrm{det},n}^{\min} = \nu_n^{-1}(f_{\mathrm{det,max}}), \tag{C11}$$

$$e_{\mathrm{det},n}^{\max} = \nu_n^{-1}(f_{\mathrm{det,min}}), \tag{C12}$$

$$\nu_n(e) = n\nu(e). \tag{C13}$$

Here $\nu(e)$ is given analytically by Equation (18), $\nu_n^{-1}(\cdot)$ denotes the inverse function of $\nu_n(e)$, $f_{\mathrm{det,min}}$ and $f_{\mathrm{det,max}}$ are the lower and upper limits of the detector's sensitive frequency band (typically 10 Hz and 10^4 Hz, respectively), and e_{LSO} is determined by Equation (21). Similarly, we define the parameters corresponding to f_n^+ and f_n^- as

$$e_{\min,n}^+ = \max(e_{\mathrm{LSO}}, e_{\mathrm{det},n+}^{\min}), \tag{C14}$$

$$e_{\max,n}^+ = \min(e_0, e_{\mathrm{det},n+}^{\max}), \tag{C15}$$

$$e_{\min,n}^- = \max(e_{\mathrm{LSO}}, e_{\mathrm{det},n-}^{\min}), \tag{C16}$$

$$e_{\max,n}^- = \min(e_0, e_{\mathrm{det},n-}^{\max}), \tag{C17}$$

where

$$e_{\mathrm{det},n+}^{\min} = \nu_{n+}^{-1}(f_{\mathrm{det,max}}), \tag{C18}$$

$$e_{\mathrm{det},n+}^{\max} = \nu_{n+}^{-1}(f_{\mathrm{det,min}}), \tag{C19}$$

$$e_{\mathrm{det},n-}^{\min} = \nu_{n-}^{-1}(f_{\mathrm{det,max}}), \tag{C20}$$

$$e_{\mathrm{det},n-}^{\max} = \nu_{n-}^{-1}(f_{\mathrm{det,min}}), \tag{C21}$$

$$\nu_{n\pm}(e) = n\nu(e) \pm \frac{\dot{\gamma}(e)}{\pi}, \tag{C22}$$

where $\dot{\gamma}(e)$ is given by Equation (25), and $\nu_{n\pm}^{-1}(\cdot)$ is the inverse function of $\nu_{n\pm}(e)$ given by Equation (C22).

In practice, the second term is negligible in Equations (C6)-(C8). We find that T_n^+ and T_n^- are within $\leq 20\%$ of T_n for any fixed n. The top panel of Figure 9 shows the total time the GW signal spends in an aLIGO type detector's sensitive frequency band for different harmonics. Higher harmonics enter the aLIGO band earlier and that depending on ρ_{p0} the first 10 orbital harmonics spend between seconds to minutes in the detector's sensitive frequency band for a $30\,\mathrm{M}_\odot - 30\,\mathrm{M}_\odot$ precessing highly eccentric BH binary.

The bottom panel of Figure 9 shows the fraction of the squared SNR that accumulates in different orbital harmonics for various ρ_{p0} for aLIGO. For high ρ_{p0}, the signal effectively circularizes by the time it enters the detector's sensitive frequency band and the $n=2$ harmonic dominates. However, the contribution of $n \neq 2$ is significant for $\rho_{p0} \lesssim 20$ for a $30\,\mathrm{M}_\odot - 30\,\mathrm{M}_\odot$ precessing highly eccentric BH binary.

## D. CALCULATING THE SIGNAL-TO-NOISE RATIO AND THE FISHER MATRIX

In this section, we derive numerically efficient formulae to calculate the SNR and the Fisher matrix for individual detectors. We first neglect pericenter precession, then extend the calculations for precessing eccentric sources.

D.1. *Signal-to-noise ratio*

D.1.1. *Eccentric inspirals without precession*

The *NoPrec* signal measured by a detector at position r is given in Fourier space from Equations (31) and (32) as

$$\tilde{h}_{\mathrm{NoPrec}} = \sum_{n=1}^{\infty} L_n(e, f_n) \Theta_{\mathrm{H}}(e_0 - e) e^{i\Psi_n(e, f_n)}, \tag{D1}$$
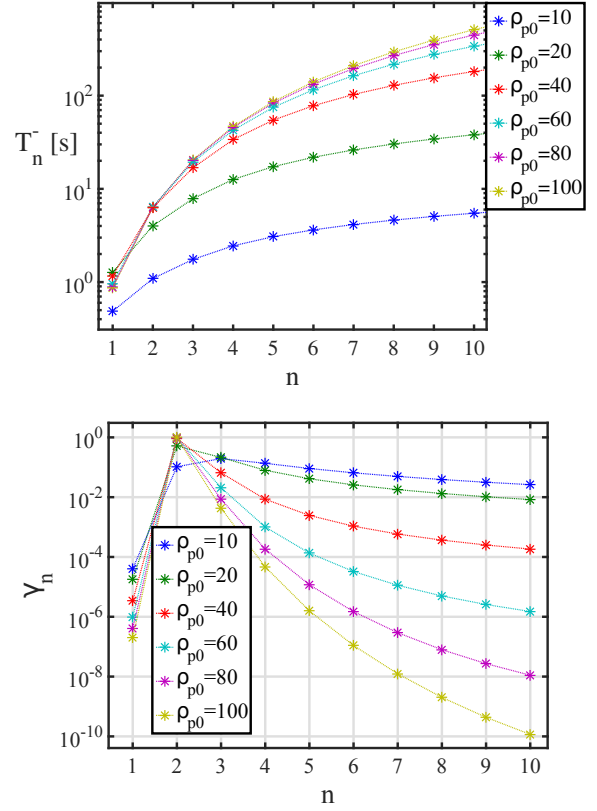
Figure 9. *Top panel*: The time duration that the first 10 harmonics (specifically f_n^- here) spend in an aLIGO type detector's sensitive frequency band for $30\,\mathrm{M}_\odot - 30\,\mathrm{M}_\odot$ precessing eccentric BH binaries with initial eccentricity $e_0 = 0.9$ for various ρ_{p0} values between 10 and 100 as labeled. For this choice of masses T_n^- varies within 10% of its value shown for $e_0 = 0.9$ for $0.9 \leq e_0 < 1$ for $n \in \{1, 2, \ldots, 10\}$. *Bottom panel*: The fraction of squared signal-to-noise ratio ($\gamma_n = (\mathrm{SNR}_n^-)^2 \cdot (\mathrm{SNR}_{\mathrm{tot}})^{-2}$) in the first ten harmonics corresponding to f_n^- frequencies in aLIGO for $30\,\mathrm{M}_\odot - 30\,\mathrm{M}_\odot$ precessing eccentric BH binaries with initial eccentricity $e_0 = 0.9$ for different ρ_{p0} values as labeled. We show results for the f_n^- components of the frequency triplet (f_n, f_n^+, f_n^-) because the dominant fraction of the $\mathrm{SNR}_{\mathrm{tot}}$ accumulates in these frequencies. For any $n \in \{1, 2, \ldots, 10\}$ and $0.9 \leq e_0 \leq 0.99$, γ_n varies by less than 60% of its value shown for $e_0 = 0.9$.

where Ψ_n is the Fourier phase at the origin of the coordinate system set to the Earth's center, given by Equation (35), we set $\dot{\gamma} \equiv 0$, $\gamma \equiv \gamma_c$, $\Theta_{\mathrm{H}}(\cdot)$ denotes the Heaviside function, which is zero and unity for negative and positive arguments[17], respec-

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[17] More precisely, we assume a smoothed truncation of the signal as

$$\Theta_{\mathrm{H}}(e_0 - e) = \begin{cases} 0 & \text{if } e > e_0 \\ \frac{e_0 - e}{\delta e_0} & \text{if } e_0 - \delta e_0 < e \leq e_0 \\ 1 & \text{if } e \leq e_0 - \delta e_0 \end{cases} \tag{D2}$$

where δe_0 is the absolute change of the eccentricity during the first orbit, which from Equations (14) and (15) is

$$\delta e_0 = 2\pi \left| \frac{\dot{e}}{\nu} \right|_{e_0} = \frac{1216\pi^2}{15} \frac{\eta}{\rho_{p0}^{5/2}} \frac{e_0}{(1+e_0)^{5/2}} \left(1 + \frac{121}{304}e_0^2\right), \tag{D3}$$

where $\eta = (\mathcal{M}_z/M_{\mathrm{tot},z})^{5/3}$.

tively, and

$$L_n(e, f_n) = - \left[\frac{A_n \sin^2 \Theta}{4} + \frac{(1+\cos^2 \Theta)}{4} (B_n^+ e^{2i\gamma_c} - B_n^- e^{-2i\gamma_c}) \right]$$
$$\times h_0 F_+ \Lambda_n e^{i(\Delta\Phi_n - \pi/4)}$$
$$- \frac{ih_0 F_\times \Lambda_n \cos\Theta}{2} (B_n^+ e^{2i\gamma_c} + B_n^- e^{-2i\gamma_c}) e^{i(\Delta\Phi_n + \pi/4)}, \quad (D4)$$

where h_0 and Λ_n are given by Equations (33) and (34), A_n and $B_n^\pm$ are given by Equations (5) and (7), γ_c specifies the argument of pericenter, which is assumed to be fixed here, F_+ and $F_\times$ are the antenna factors given by Equation (B12). The factor $\Delta\Phi_n = -2\pi f_n \mathbf{N} \cdot \mathbf{r}$ gives the phase shift of the measured signal between the position of the detector r and the origin of the coordinate system for the nth harmonic (Appendix B). L_n depends on f_n implicitly through $\Delta\Phi_n$ and h_0.

In Equation (D1), $\Theta_\mathrm{H}(e_0 - e)$ accounts for the start of the waveform when the binary forms with initial eccentricity[18] e_0. Along the same lines, a similar term $\Theta_\mathrm{H}(e - e_{\mathrm{LSO}})$ could be incorporated to account for the end of the eccentric inspiral where the waveform transitions to a plunge and ringdown phase. However, we conservatively do not account for such a term, since the waveform near the end of the inspiral is sensitive to higher order post-Newtonian corrections, which are not known and neglected here (Kocsis & Levin 2012; Loutrel & Yunes 2017). Nevertheless, the inspiral rate is sensitive to e_{LSO} in Equations (18) and (19) and (22), which affects L_n and Ψ_n.

For each detector, the square of the SNR for the *NoPrec* waveform, $\mathrm{SNR}_{\mathrm{NoPrec}}^2$, can be obtained by substituting $\tilde{h}_{\mathrm{NoPrec}}$ into Equation (42). We find that the product of sums in $\tilde{h}_{\mathrm{NoPrec}} \tilde{h}_{\mathrm{NoPrec}}^*$ is dominated by the elements such that [19]

$$(\mathrm{SNR}_{\mathrm{NoPrec}})^2 \approx 4 \sum_{n=1}^{\infty} \int_{f_{\min,n}}^{f_{\max,n}} \frac{|L_n(e(f_n), f_n)|^2}{S_n(f_n)} df_n, \quad (D5)$$

where $f_{\min,n}$ is the frequency at which the nth harmonic first enters the detector's sensitive frequency band or when it forms in the band, and similarly $f_{\max,n}$ is the frequency at which the signal exits the detector's sensitive frequency band or when it reaches the LSO,

$$f_{\max,n} = \min(\nu_n(e_{\mathrm{LSO}}), f_{\mathrm{det,max}}), \quad (D6)$$
$$f_{\min,n} = \max(\nu_n(e_0), f_{\mathrm{det,min}}). \quad (D7)$$

Computationally it is practical to change the integration variable from f_n to e as

$$df_n = n \left| \frac{d\nu}{de} \right| de, \quad (D8)$$

thus Equation (D5) can be rewritten generally as

$$\left(\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}\right)^2 \approx 4 \sum_{n=1}^{n_{\max}(e_0)} \int_{e_{\min,n}}^{e_{\max,n}} \frac{n|L_n(e)|^2}{S_n(n\nu(e))} \left| \frac{d\nu}{de} \right| de, \quad (D9)$$

---

[18] The initial eccentricity e_0 does not enter the waveform anywhere else, L_n and Ψ_n are independent of e_0. Due to this term, e_0 and e_{LSO} may be measured independently, and $\Delta\rho_{\mathrm{p0}}$ follows from Equation (50).

[19] Numerically we confirm that cross-terms proportional to $L_n L_m^* \exp(i\Psi_n - i\Psi_m)$ have a negligible contribution for $n \neq m$.

where $\nu(e)$ is given analytically by Equation (18), and $L(e)$ may be obtained from Equation (D4) by substituting $f_n = n\nu(e)$. The integration bounds $e_{\min,n}$ and $e_{\max,n}$ are given by Equations (C9) and (C10). We truncate the calculation beyond a maximum spectral harmonic $n_{\max}(e)$ defined in Equation (37).

$\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ accurately recovers Equation (D5) generally for any value of the initial eccentricity in the range $0 < e_0 < 1$, however the number of considered harmonics $n_{\max}(e_0)$ increases rapidly for high e_0 (Equation 37), and $n_{\max}(e_0) \to \infty$ in the limit $e_0 \to 1$. Therefore, $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ is computationally efficient for low to moderate initial eccentricities ($e_0 \lesssim 0.8$) and it is inefficient for higher e_0. In order to make $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ computationally efficient for high initial eccentricities, we reverse the order of the sum and the integral in Equation (D9) and truncate the sum over harmonics at $n_{\max}(e)$ (O'Leary et al. 2009).[20] Thus, we get

$$\left(\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{High}}\right)^2 \approx 4 \int_{e_{\min,n}}^{e_{\max,n}} \sum_{n=1}^{n_{\max}(e)} \frac{n|L_n(e)|^2}{S_n(n\nu(e))} \left| \frac{d\nu}{de} \right| de. \quad (D10)$$

We use the above introduced trick to derive computationally efficient formulae in the high initial eccentricity limit for Fisher matrix elements in the precession-free case (Appendix D.2.1) and for the SNR and the Fisher matrix elements in the precessing case (Appendices D.1.2 and D.2.2).

D.1.2. *Eccentric inspirals with precession*

We derive the SNR of the precessing model in this section. The Fourier-transformed waveform given by Equation (B14) can be rewritten as

$$\tilde{h}_{\mathrm{Prec}} = \sum_{n=1}^{\infty} K_n(e, f_n) \Theta_\mathrm{H}(e_0 - e) e^{i\Psi_n(e,f_n)}$$
$$+ \sum_{n=1}^{\infty} K_n^+(e, f_n^+) \Theta_\mathrm{H}(e_0 - e) e^{i\Psi_n^+(e,f_n^+)}$$
$$+ \sum_{n=1}^{\infty} K_n^-(e, f_n^-) \Theta_\mathrm{H}(e_0 - e) e^{i\Psi_n^-(e,f_n^-)}, \quad (D11)$$

where the terms K_n, K_n^+, and K_n^- are defined as

$$K_n(e, f_n) = -\frac{h_0 \sin^2 \Theta}{4} A_n \Lambda_n F_+ e^{i(\Delta\Phi_n - \pi/4)}, \quad (D12)$$

$$K_n^+(e, f_n^+) = -\frac{h_0 \cos \Theta}{2} B_n^- \Lambda_n^+ F_\times e^{i(\Delta\Phi_n^+ + \pi/4)}$$
$$+ \frac{h_0 (1+\cos^2 \Theta)}{4} B_n^- \Lambda_n^+ F_+ e^{i(\Delta\Phi_n^+ - \pi/4)}, \quad (D13)$$

$$K_n^-(e, f_n^-) = -\frac{h_0 \cos \Theta}{2} B_n^+ \Lambda_n^- F_\times e^{i(\Delta\Phi_n^- + \pi/4)}$$
$$- \frac{h_0}{4} B_n^+ \Lambda_n^- F_+ \left(1+\cos^2 \Theta\right) e^{i(\Delta\Phi_n^- - \pi/4)}. \quad (D14)$$

The terms K_n and $K_n^\pm$ depend on ν through h_0, which are expressed with $f_n, f_n^\pm$ using Equations (29) and (30). Furthermore these equations depend on f_n and $f_n^\pm$ through $\Delta\Phi_n$ and $\Delta\Phi_n^\pm$ (e.g. $\Delta\Phi_n = -2\pi f_n \mathbf{N} \cdot \mathbf{r}$ and $\Delta\Phi_n^\pm = -2\pi f_n^\pm \mathbf{N} \cdot \mathbf{r}$).

Next, we substitute this waveform $\tilde{h}_{\mathrm{Prec}}$ into Equation (42). Similarly to that of the *NoPrec* signal (Equation D1), cross

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[20] In this case the number of considered harmonics reduces significantly.

terms in the product of sums in $\tilde{h}_{\text{Prec}}\tilde{h}^*_{\text{Prec}}$ have negligible contributions to $(\text{SNR}_{\text{Prec}})^2$, and so

$$(\text{SNR}_{\text{Prec}})^2 \approx 4\sum_{n=1}^{\infty}\int_{f_{\text{min,n}}}^{f_{\text{max,n}}}\frac{|K_n(e(f_n),f_n)|^2}{S_h(f_n)}\,df_n$$
$$+4\sum_{n=1}^{\infty}\int_{f^+_{\text{min,n}}}^{f^+_{\text{max,n}}}\frac{|K^+_n(e(f^+_n),f^+_n)|^2}{S_h(f^+_n)}\,df^+_n$$
$$+4\sum_{n=1}^{\infty}\int_{f^-_{\text{min,n}}}^{f^-_{\text{max,n}}}\frac{|K^-_n(e(f^-_n),f^-_n)|^2}{S_h(f^-_n)}\,df^-_n\,, \qquad (D15)$$

where $f_{\text{max,n}}$ and $f_{\text{min,n}}$ are defined in Equations (D6) and (D7). The integration bounds for the integrals over $f^\pm_n$ are defined similarly to $f_{\text{max,n}}$ and $f_{\text{min,n}}$ in Equation (D6) and (D7),

$$f^\pm_{\text{max},n} = \min(\nu_{n\pm}(e_{\text{LSO}}), f_{\text{det,max}})\,, \qquad (D16)$$
$$f^\pm_{\text{min},n} = \max(\nu_{n\pm}(e_0), f_{\text{det,min}})\,, \qquad (D17)$$

where $\nu_{n\pm}$ is defined in Equation (C22).

Next, we change integration variables from f_n to e using Equation (D8) and similarly from $f^\pm_n$ to e using Equation (30) as

$$df^\pm_n = \left|n\pm\frac{1}{\pi}\frac{d\dot\gamma}{d\nu}\right|\left|\frac{d\nu}{de}\right|\,de\,. \qquad (D18)$$

Here $d\dot\gamma/d\nu$ is given by Equation (25) as

$$\frac{d\dot\gamma}{d\nu} = \frac{\partial\dot\gamma(\nu,e)}{\partial\nu} + \frac{\partial\dot\gamma(\nu,e)}{\partial e}\frac{1}{d\nu/de}$$
$$= \left(\frac{5}{3\nu} - \frac{2e}{1-e^2}\frac{1}{d\nu/de}\right)\dot\gamma\,. \qquad (D19)$$

After truncating the sum over the harmonics to the relevant range as in Equation (D9), Equation (D15) can be written as

$$\left(\text{SNR}^{\text{Gen}}_{\text{Prec}}\right)^2 \approx 4\sum_{n=1}^{n_{\text{max}}(e_0)}\int_{e_{\text{min,n}}}^{e_{\text{max,n}}}\frac{n|K_n(e)|^2}{S_n(\nu_n(e))}\left|\frac{d\nu}{de}\right|\,de$$
$$+4\sum_{n=1}^{n_{\text{max}}(e_0)}\int_{e^+_{\text{min,n}}}^{e^+_{\text{max,n}}}\frac{|K^+_n(e)|^2}{S_n(\nu_{n+}(e))}\left|n+\frac{1}{\pi}\frac{d\dot\gamma}{d\nu}\right|\left|\frac{d\nu}{de}\right|\,de$$
$$+4\sum_{n=1}^{n_{\text{max}}(e_0)}\int_{e^-_{\text{min,n}}}^{e^-_{\text{max,n}}}\frac{|K^-_n(e)|^2}{S_n(\nu_{n-}(e))}\left|n-\frac{1}{\pi}\frac{d\dot\gamma}{d\nu}\right|\left|\frac{d\nu}{de}\right|\,de\,, \qquad (D20)$$

where $K_n(e) \equiv K_n(e,\nu_n(e))$ and $K^\pm_n(e) \equiv K^\pm_n(e,\nu_{n\pm}(e))$. The integration bounds are given by Equations (C14) and (C17).

Similarly to $\text{SNR}^{\text{Gen}}_{\text{NoPrec}}$, $\text{SNR}^{\text{Gen}}_{\text{Prec}}$ is computationally efficient only for low to moderate initial eccentricities ($e_0 \lesssim 0.8$). For high initial eccentricities, the computationally efficient form of $\text{SNR}^{\text{Gen}}_{\text{Prec}}$, $\text{SNR}^{\text{High}}_{\text{Prec}}$, can be given by reversing the order

of the sum and the integral as

$$\left(\text{SNR}^{\text{High}}_{\text{Prec}}\right)^2 \approx 4\int_{e_{\text{min,n}}}^{e_{\text{max,n}}}\sum_{n=1}^{n_{\text{max}}(e)}\frac{n|K_n(e)|^2}{S_n(\nu_n(e))}\left|\frac{d\nu}{de}\right|\,de$$
$$+4\int_{e^+_{\text{min,n}}}^{e^+_{\text{max,n}}}\sum_{n=1}^{n_{\text{max}}(e)}\frac{|K^+_n(e)|^2}{S_n(\nu_{n+}(e))}\left|n+\frac{1}{\pi}\frac{d\dot\gamma}{d\nu}\right|\left|\frac{d\nu}{de}\right|\,de$$
$$+4\int_{e^-_{\text{min,n}}}^{e^-_{\text{max,n}}}\sum_{n=1}^{n_{\text{max}}(e)}\frac{|K^-_n(e)|^2}{S_n(\nu_{n-}(e))}\left|n-\frac{1}{\pi}\frac{d\dot\gamma}{d\nu}\right|\left|\frac{d\nu}{de}\right|\,de\,, \qquad (D21)$$

D.2. *Fisher Matrix*

Due to the similarity of the equations defining the SNR (see Equation (42)) and the Fisher matrix (see Equation (44)), we may follow the same procedure to derive numerically efficient formulae for the Fisher matrix in the limit of high initial eccentricity. Similarly to Appendix D.1, we start the analysis with the *NoPrec* model and then generalize the calculation to the *Prec* model, which accounts for the precessing case.

D.2.1. *Eccentric inspirals without precession*

Let us substitute Equation (D1) into Equation (44). Similarly to the product of sums of orbital harmonics $\tilde{h}_{\text{NoPrec}}\tilde{h}^*_{\text{NoPrec}}$, we find numerically that the cross-terms in $\partial_j\tilde{h}_{\text{NoPrec}}\,\partial_k\tilde{h}^*_{\text{NoPrec}}$ with $j \neq k$ have a negligible contribution to Γ_{jk}. Thus, we find that the stationary phase approximation is applicable if we drop the cross-terms, and thus in Equation (44) we may use

$$\partial_j\tilde{h}_{\text{NoPrec}}\,\partial_k\tilde{h}^*_{\text{NoPrec}} = \sum_{n=1}^{\infty}\tilde{L}_{n,j}(e)\tilde{L}^*_{n,k}(e)\Theta_H(e_0-e)$$
$$+\delta_{e_0,j}\sum_{n=1}^{\infty}\tilde{L}_n(e)\tilde{L}^*_{n,k}(e)\delta(e-e_0)$$
$$+\delta_{e_0,k}\sum_{n=1}^{\infty}\tilde{L}_{n,j}(e)\tilde{L}^*_n(e)\delta(e-e_0)$$
$$+\delta_{e_0,k}\delta_{e_0,j}\sum_{n=1}^{\infty}|L_n(e)|^2\left[\partial_{e_0}\Theta_H(e_0-e)\right]^2\,. \qquad (D22)$$

Here

$$\tilde{L}_n(e) = L_n(f_n(e),e)e^{i\Psi_n(f_n(e),e)}\,, \qquad (D23)$$

and

$$\tilde{L}_{n,j}(e) = \partial_j\left[L_n(f_n)e^{i\Psi_n(f_n)}\right](e)\,, \qquad (D24)$$

where

$$L_n(f_n) \equiv L_n(e(f_n),f_n)\,, \quad \Psi_n(f_n) \equiv \Psi_n(e(f_n),f_n)\,, \qquad (D25)$$

and $L_n(e,f_n)$ and $\Psi_n(e,f_n)$ are given by Equations (D4) and (35). We first differentiate the expressions in the bracket [ ] in Equation (D24) with respect to λ_j, then change the variable from f_n back to e. Note that for all n and e, $\tilde{L}_n(e)$ is independent of e_0, and so $\tilde{L}_{n,j}(e) = 0$ for $\lambda_j = e_0$. In Equation (D22), $\delta_{a,b}$ in the second, third, and fourth terms denote the Kronecker-δ, defined to be unity if $a = b$ and zero otherwise. In Equation (D22), the second and third terms arise due to the Heaviside function in the waveform in Equation (D1), which

represents the start of the waveform with eccentricity e_0. The e_0-derivative of this function is $\delta(e_0-e)$, which denotes the Dirac-δ function. Note that we use a smoothed version of $\Theta_H(e_0-e)$ over a scale δe_0, which is given in Equation (D2), whose derivative is approximately[21]

$$[\partial_{e_0}\Theta_H(e_0-e)]^2 \approx \frac{1}{(\delta e_0)^2} \text{ if } e_0-\delta e_0 \leq e < e_0$$
$$\approx \frac{\delta(e_0-e)}{\delta e_0}. \tag{D26}$$

To avoid confusion note that the numerator denotes the Dirac-δ function, which has a unit integral over $e \approx e_0$, and δe_0 in the denominator is the quantity given by Equation (D3).

Furthermore, we note that

$$\Re\left(\tilde{L}_n(e_0)\tilde{L}_{n,k}^*(e_0)\right) = \frac{1}{2}\partial_k |L_n(e_0)|^2 \tag{D27}$$

in Equation (D22). In these equations f_n enters when substituting f_n/n for ν. By substituting Equation (D22) into Equation (44), changing the integration variable from f_n to e respectively for each harmonic using[22] $f_n = n\nu(e)$ and Equations (18) and (22), and truncating the sum over the harmonics to the relevant range, the Fisher matrix becomes[23]

$$\Gamma_{jk}^{\text{NoPrec,Gen}} \approx 4 \sum_{n=1}^{n_{\max}(e_0)} \int_{e_{\min,n}}^{e_{\max,n}} \frac{\Re\left(\tilde{L}_{n,j}(e)\tilde{L}_{n,k}^*(e)\right)}{S_n(\nu_n(e))} \left|n\frac{d\nu}{de}\right| de$$
$$+2\delta_{k,e_0} \sum_n \left|n\frac{d\nu}{de}\right|_{e_0} \frac{\partial_j |L_n(e_0)|^2}{S_n[\nu_n(e_0)]}$$
$$+2\delta_{j,e_0} \sum_n \left|n\frac{d\nu}{de}\right|_{e_0} \frac{\partial_k |L_n(e_0)|^2}{S_n[\nu_n(e_0)]}$$
$$+4\frac{\delta_{j,e_0}\delta_{k,e_0}}{\delta e_0} \sum_n \left|n\frac{d\nu}{de}\right|_{e_0} \frac{|L_n(e_0)|^2}{S_n(\nu_n(e_0))}. \tag{D28}$$

The limits of integration in Equation (D28) are defined by Equations (C9) and (C10). Here the four terms correspond respectively to the four terms in Equation (D22). The first term is directly analogous to that appearing in the SNR (see Equation D10). Note that in particular, the elements corresponding to $j = e_{\text{LSO}}$ and $k = e_{\text{LSO}}$ terms are nonzero. The e_{LSO}-dependence enters in $\nu(e)$ as shown in Equations (18) and (22). However, the first term in Equation (D28) is zero for the $j = e_0$ and $k = e_0$ elements. If the binary forms in the detector's sensitive frequency band the second, the third, and the fourth terms in Equation (D22) contribute to this element of the Fisher matrix. The eccentricity integral in the Fisher matrix may be carried analytically over the δ function, which yields the second and third terms in Equation (D28). There δ_{j,e_0} is the Kronecker-δ, which is zero unless j corresponds to the parameter e_0, and similarly for δ_{k,e_0}. Note further that only harmonics with $f_{\max,\text{det}}/\nu(e_0) \geq n \geq f_{\min,\text{det}}/\nu(e_0)$ contribute to these boundary terms, since otherwise $S_n(\nu_n(e_0)) = \infty$.

Similarly to the $\text{SNR}_{\text{NoPrec}}^{\text{Gen}}$, $\Gamma_{jk}^{\text{NoPrec}}$ is generally valid for any initial eccentricity in the range $0 < e_0 < 1$, but it is computationally efficient only for low to moderate initial eccentricities ($e_0 \lesssim 0.8$). In order to make the calculation computationally

efficient for high initial eccentricities, we reverse the order of the sum and the integral in the first term in Equation (D28), and truncate the sum over harmonics at $n_{\max}(e)$ as in Appendix D.1.1. We get

$$\Gamma_{jk}^{\text{NoPrec,High}} \approx 4 \int_{e_{\min,n}}^{e_{\max,n}} \sum_{n=1}^{n_{\max}(e)} \frac{\Re\left(\tilde{L}_{n,j}(e)\tilde{L}_{n,k}^*(e)\right)}{S_n(\nu_n(e))} \left|n\frac{d\nu}{de}\right| de$$
$$+2\delta_{k,e_0} \sum_n \left|n\frac{d\nu}{de}\right|_{e_0} \frac{\partial_j |L_n(e_0)|^2}{S_n[\nu_n(e_0)]}$$
$$+2\delta_{j,e_0} \sum_n \left|n\frac{d\nu}{de}\right|_{e_0} \frac{\partial_k |L_n(e_0)|^2}{S_n[\nu_n(e_0)]}$$
$$+4\frac{\delta_{j,e_0}\delta_{k,e_0}}{\delta e_0} \sum_n \left|n\frac{d\nu}{de}\right|_{e_0} \frac{|L_n(e_0)|^2}{S_n(\nu_n(e_0))}. \tag{D29}$$

D.2.2. *Eccentric inspirals with precession*

Following the steps of Appendix D.2.1 for the precession-free model, we may generalize the calculation of the Fisher matrix to include precession similar to Appendix D.1.2. The Fisher matrix, which is computationally efficient for low to moderate initial eccentricities ($e_0 \lesssim 0.8$), can be given as

$$\Gamma_{jk}^{\text{Prec,Gen}} = \Gamma_{jk}^{\text{Gen},n} + \Gamma_{jk}^{\text{Gen},n+} + \Gamma_{jk}^{\text{Gen},n-}, \tag{D30}$$

where

$$\Gamma_{jk}^{\text{Gen},n} \approx 4 \sum_{n=1}^{n_{\max}(e_0)} \int_{e_{\min,n}}^{e_{\max,n}} \frac{\Re\left(\tilde{K}_{n,j}(e)\tilde{K}_{n,k}^*(e)\right)}{S_n(\nu_n(e))} \left|n\frac{d\nu}{de}\right| de$$
$$+2\delta_{j,e_0} \sum_n \left|n\frac{d\nu}{de}\right|_{e_0} \frac{\partial_k |K_n(e_0)|^2}{S_n(\nu_n(e_0))}$$
$$+2\delta_{k,e_0} \sum_n \left|n\frac{d\nu}{de}\right|_{e_0} \frac{\partial_j |K_n(e_0)|^2}{S_n(\nu_n(e_0))}$$
$$+4\frac{\delta_{j,e_0}\delta_{k,e_0}}{\delta e_0} \sum_n \left|n\frac{d\nu}{de}\right|_{e_0} \frac{|K_n(e_0)|^2}{S_n(\nu_n(e_0))}, \tag{D31}$$

and

$$\Gamma_{jk}^{\text{Gen},n\pm} \approx 4 \sum_{n=1}^{n_{\max}(e_0)} \int_{e_{\min,n}^{\pm}}^{e_{\max,n}^{\pm}} de\, \frac{\Re\left(\tilde{K}_{n,j}^{\pm}(e)\tilde{K}_{n,k}^{\pm,*}(e)\right)}{S_n(\nu_n^{\pm}(e))}$$
$$\times \left|n\pm\frac{1}{\pi}\frac{d\dot{\gamma}}{d\nu}\right| \left|\frac{d\nu}{de}\right|$$
$$+2\delta_{j,e_0} \sum_n \left|n\pm\frac{1}{\pi}\frac{d\dot{\gamma}}{d\nu}\right|_{e_0} \left|\frac{d\nu}{de}\right|_{e_0} \frac{\partial_k |K_n^{\pm}(e_0)|^2}{S_n(\nu_n^{\pm}(e_0))}$$
$$+2\delta_{k,e_0} \sum_n \left|n\pm\frac{1}{\pi}\frac{d\dot{\gamma}}{d\nu}\right|_{e_0} \left|\frac{d\nu}{de}\right|_{e_0} \frac{\partial_j |K_n^{\pm}(e_0)|^2}{S_n(\nu_n^{\pm}(e_0))}$$
$$+4\frac{\delta_{j,e_0}\delta_{k,e_0}}{\delta e_0} \sum_n \left|n\pm\frac{1}{\pi}\frac{d\dot{\gamma}}{d\nu}\right|_{e_0} \left|\frac{d\nu}{de}\right|_{e_0} \frac{|K_n^{\pm}(e_0)|^2}{S_n(\nu_n^{\pm}(e_0))}, \tag{D32}$$

---

[21] We neglect the partial derivatives of δe_0 with respect to the physical parameters.

[22] Note that $\nu(e)$ depends on e_{LSO} as seen in Equations (18) and (22).

[23] We label this general expression with "Gen" to distinguish from an approximation "High" used below for high eccentricities.

where the integration bounds are given by Equations (C14)-(C17), and

$$\tilde{K}_{n,j}(e) = \partial_j \left[K_n(f_n) e^{i\Psi_n(f_n)} \right](e), \tag{D33}$$

$$\tilde{K}_{n,j}^{\pm}(e) = \partial_j \left[K_n^{\pm}(f_n^{\pm}) e^{i\Psi_n^{\pm}(f_n^{\pm})} \right](e). \tag{D34}$$

Similar to the precession-free case, here

$$K_n(f_n) \equiv K_n(e(f_n), f_n), \quad \Psi_n(f_n) \equiv \Psi_n(e(f_n), f_n), \tag{D35}$$

$$K_n^{\pm}(f_n) \equiv K_n^{\pm}(e(f_n^{\pm}), f_n^{\pm}), \quad \Psi_n^{\pm}(f_n) \equiv \Psi_n^{\pm}(e(f_n^{\pm}), f_n^{\pm}), \tag{D36}$$

where $K_n(e, f_n)$ and $K_n^{\pm}(e, f_n^{\pm})$ are given by Equations (D12)-(D14), $\Psi_n(e, f_n)$ and $\Psi_n^{\pm}(e, f_n^{\pm})$ are expressed by Equations (35) and (36), and we first differentiate the expressions in the bracket [] in Equations (D33) and (D34) with respect to λ_j, then change variables from f_n and $f_n^{\pm}$ back to e. Similar to the precession-free case, only harmonics with $f_{\mathrm{max,det}}/\nu(e_0) \geq n \geq f_{\mathrm{min,det}}/\nu(e_0)$ and $f_{\mathrm{max,det}}/\nu^{\pm}(e_0) \geq n \geq f_{\mathrm{min,det}}/\nu^{\pm}(e_0)$ contribute to the boundary terms in Equations (D31) and (D32), since otherwise $S_n(\nu_n(e_0)) = \infty$ and $S_n(\nu_n^{\pm}(e_0)) = \infty$.

For high initial eccentricities, the computationally efficient form of $\Gamma_{jk}^{\mathrm{Prec,Gen}}$, $\Gamma_{jk}^{\mathrm{Prec,High}}$, can be derived by reversing the order of the sum and the integral in the first term in Equations (D31) and (D32), and truncating the sum over harmonics at $n_{\mathrm{max}}(e)$ as in Appendix D.1.1. Thus, $\Gamma_{jk}^{\mathrm{Prec,High}}$ can be given as

$$\Gamma_{jk}^{\mathrm{Prec,High}} = \Gamma_{jk}^{\mathrm{High},n} + \Gamma_{jk}^{\mathrm{High},n+} + \Gamma_{jk}^{\mathrm{High},n-}, \tag{D37}$$

where

$$\begin{aligned}
\Gamma_{jk}^{\mathrm{High},n} &\approx 4 \int_{e_{\mathrm{min},n}}^{e_{\mathrm{max},n}} \sum_{n=1}^{n_{\mathrm{max}}(e)} \frac{\Re\left(\tilde{K}_{n,j}(e)\tilde{K}_{n,k}^*(e)\right)}{S_n(\nu_n(e))} \left| n\frac{d\nu}{de} \right| de \\
&+ 2\delta_{j,e_0} \sum_n \left| n\frac{d\nu}{de} \right|_{e_0} \frac{\partial_k |K_n(e_0)|^2}{S_n(\nu_n(e_0))} \\
&+ 2\delta_{k,e_0} \sum_n \left| n\frac{d\nu}{de} \right|_{e_0} \frac{\partial_j |K_n(e_0)|^2}{S_n(\nu_n(e_0))} \\
&+ 4\frac{\delta_{j,e_0}\delta_{k,e_0}}{\delta e_0} \sum_n \left| n\frac{d\nu}{de} \right|_{e_0} \frac{|K_n(e_0)|^2}{S_n(\nu_n(e_0))},
\end{aligned} \tag{D38}$$

and

$$\begin{aligned}
\Gamma_{jk}^{\mathrm{High},n\pm} &\approx 4 \int_{e_{\mathrm{min},n}^{\pm}}^{e_{\mathrm{max},n}^{\pm}} de \sum_{n=1}^{n_{\mathrm{max}}(e)} \frac{\Re\left(\tilde{K}_{n,j}^{\pm}(e)\tilde{K}_{n,k}^{\pm,*}(e)\right)}{S_n(\nu_n^{\pm}(e))} \\
&\times \left| n \pm \frac{1}{\pi}\frac{d\dot{\gamma}}{d\nu} \right| \left| \frac{d\nu}{de} \right| \\
&+ 2\delta_{j,e_0} \sum_n \left| n \pm \frac{1}{\pi}\frac{d\dot{\gamma}}{d\nu} \right|_{e_0} \left| \frac{d\nu}{de} \right|_{e_0} \frac{\partial_k |K_n^{\pm}(e_0)|^2}{S_n(\nu_n^{\pm}(e_0))} \\
&+ 2\delta_{k,e_0} \sum_n \left| n \pm \frac{1}{\pi}\frac{d\dot{\gamma}}{d\nu} \right|_{e_0} \left| \frac{d\nu}{de} \right|_{e_0} \frac{\partial_j |K_n^{\pm}(e_0)|^2}{S_n(\nu_n^{\pm}(e_0))} \\
&+ 4\frac{\delta_{j,e_0}\delta_{k,e_0}}{\delta e_0} \sum_n \left| n \pm \frac{1}{\pi}\frac{d\dot{\gamma}}{d\nu} \right|_{e_0} \left| \frac{d\nu}{de} \right|_{e_0} \frac{|K_n^{\pm}(e_0)|^2}{S_n(\nu_n^{\pm}(e_0))}.
\end{aligned} \tag{D39}$$

E. VALIDATION OF CODES

E.1. *Analytic circular limit without precession*

First, we study the circular limit of the eccentric waveform, $\tilde{h}_{\times}(\mathbf{f})$ and $\tilde{h}_+(\mathbf{f})$, defined by Equations (31) and (32) for the *NoPrec* model. For $e \to 0$, $B_n^+ = \delta_{n,2}$ and $\Psi_n^{\pm} = \Psi_n \mp 2\gamma_c$ implying that γ_c is degenerate with Φ_c, which we henceforth omit. After integrating the term $\nu'/\dot{\nu}(\nu')$ over ν' in Equation (26) according to previous considerations, substituting ν with $f/2$, and expanding the expression of $h_0\Lambda_2$, the polarization components become

$$h_{\times}(f) = -2i\sqrt{\frac{5}{96}} \frac{\mathcal{M}_z^{5/6} f^{-7/6} e^{i\Psi}}{\pi^{2/3} D_{\mathrm{L}}} \cos\Theta, \tag{E1}$$

$$h_+(f) = -\sqrt{\frac{5}{96}} \frac{\mathcal{M}_z^{5/6} f^{-7/6} e^{i\Psi}}{\pi^{2/3} D_{\mathrm{L}}} (1 + \cos^2\Theta), \tag{E2}$$

where the phase function Ψ can be given as

$$\Psi = 2\pi f t_c - \Phi_c - \pi/4 + \frac{3}{4}(8\pi\mathcal{M}_z f)^{-5/3}. \tag{E3}$$

These are indeed the well-known frequency domain polarization components of circular binaries in leading order (Cutler & Flanagan 1994). The parameter set characterizing this waveform is

$$\boldsymbol{\lambda}_{\mathrm{circ}} = \{\ln(D_{\mathrm{L}}), \ln(\mathcal{M}_z), \theta_N, \phi_N, \theta_L, \phi_L, t_c, \Phi_c\}. \tag{E4}$$

For validation tests, we calculate the SNR of circular binaries, $\mathrm{SNR}_{\mathrm{circ}}$, for a single aLIGO detector. $\mathrm{SNR}_{\mathrm{circ}}$ is calculated by substituting Equations (E1) and (E2) into Equation (B13) and then using Equation (42). The Fisher matrix of circular binaries for the parameter set $\boldsymbol{\lambda}_{\mathrm{circ}}$ for each detector $\Gamma_{jk}^{\mathrm{circ}}$ is calculated by substituting Equations (E1) and (E2) into Equation (B13) and then using Equations (44).

E.2. *Eccentric inspiral without pericenter precession*

Next, we discuss validation tests performed for the codes using the *NoPrec* waveform model. In this case, the parameters are

$$\boldsymbol{\lambda}_{\mathrm{NoPrec}} = \{\ln(D_{\mathrm{L}}), \ln(\mathcal{M}_z), \theta_N, \phi_N, \theta_L, \phi_L, t_c, \Phi_c, c_0, e_0, \gamma_c\}. \tag{E5}$$

Compared to $\boldsymbol{\lambda}_{\mathrm{circ}}$, $\boldsymbol{\lambda}_{\mathrm{NoPrec}}$ includes c_0 (set by $M_{\mathrm{tot},z}$ and e_{LSO}, see Equation (22)) and e_0.

E.2.1. *Signal-to-noise ratio*

First, we generate a set of source parameters for comparison $(m_A, m_B, D_{\mathrm{L}}, \theta_N, \phi_N, \theta_L, \phi_L)$, and compare the output of $\mathrm{SNR}_{\mathrm{circ}}$ with the output of $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ in the circular limit for a single aLIGO detector (Table 1). Here and in further validation test, the set of fiducial source parameters generated for comparison are $(m_A, m_B, D_{\mathrm{L}}, \theta_N, \phi_N, \theta_L, \phi_L)$, where $\cos\theta_N$ and $\cos\theta_L$ are drawn from a uniform distribution between $[-1, 1]$, the set of ϕ_N and ϕ_L are drawn from a uniform distribution between $[0, 2\pi]$, m_A and m_B are drawn from a uniform distribution between $[5\,\mathrm{M}_{\odot}, 100\,\mathrm{M}_{\odot}]$, and D_{L} is drawn from a uniform distribution between $[100\,\mathrm{Mpc}, 1000\,\mathrm{Mpc}]$. The generation of other source parameters are described in details in the corresponding paragraph. We assume the fiducial value $t_c = \Phi_c = \gamma_c = 0$ for each binary, and in practice we set $e_0 = 10^{-4}$ and $\rho_{\mathrm{p0}} = 1000$ when considering the circular limit (Appendix A). We find that the relative discrepancy between $\mathrm{SNR}_{\mathrm{circ}}$ and $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ is less than 10^{-3} in all cases.

Next, we examine if the output of $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ agrees with Figure 11 in O'Leary et al. (2009), which shows the source

sky position- and orientation-averaged RMS SNR for a single aLIGO detector as a function of ρ_{p0} and M_{tot}. O'Leary et al. (2009) used the Fourier domain orbit-averaged leading order waveform (Peters & Mathews 1963), which corresponds to our *NoPrec* model. We set the sensitivity curve to that used in O'Leary et al. (2009), set $e = 0.95$, and find the results for several ρ_{p0} and M_{tot} in the range $[5, 100]$ and of $[10\,M_\odot, 1000\,M_\odot]$, respectively. For each ρ_{p0} and M_{tot}, we generate random Monte Carlo samples of source sky location (θ_N, ϕ_N) and binary orientation (θ_L, ϕ_L) as introduced above in this section. We find that the RMS of the $\left(\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}\right)^2$ distributions are in agreement with the results of Figure 11 in O'Leary et al. (2009).

Finally, we generate a set of source parameters for comparison, and compare the output of $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ and $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{High}}$ for several high e_0 and ρ_{p0} and for a single aLIGO detector. In particular, e_0 and ρ_{p0} are drawn from a uniform distribution between $]0.9, 1[$ and $[5, 1000]$ in these calculations, respectively. The relative discrepancy between $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ and $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{High}}$ is less than 10^{-3} in all cases.

E.2.2. *Fisher matrix*

Since the parameter set of the leading order circular and eccentric binaries, λ_{circ} and $\lambda_{\mathrm{NoPrec}}$ differ, see Equations (E4) and (E5), we cannot simply compare the output of $\Gamma_{jk}^{\mathrm{circ}}$ with the output of $\Gamma_{jk}^{\mathrm{NoPrec,Gen}}$ in the circular limit. Therefore, we first restrict $\lambda_{\mathrm{NoPrec}}$ to the parameters set λ_{circ}. Next, we generate a set of source parameters for comparison, and compare the output of $\Gamma_{jk}^{\mathrm{circ}}$ with the output of $\Gamma_{jk}^{\mathrm{NoPrec,Gen}}$ in the circular limit for the Fisher matrix elements corresponding to the circular parameters λ_{circ} and for a single aLIGO detector. The relative discrepancy between $\Gamma_{jk}^{\mathrm{circ}}$ and $\Gamma_{jk}^{\mathrm{NoPrec,Gen}}$ is less than 10^{-2} in all cases.

Finally, we generate a set of source parameters for comparison, and compare the output of $\Gamma_{jk}^{\mathrm{NoPrec,Gen}}$ and $\Gamma_{jk}^{\mathrm{NoPrec,High}}$ for the Fisher matrix elements corresponding to the parameter set $\lambda_{\mathrm{NoPrec}}$ in the high eccentricity limit for several ρ_{p0} and for a single aLIGO detector. Similarly to Appendix E.2.1, e_0 and ρ_{p0} are drawn from a uniform distribution between $]0.9, 1[$ and $[5, 1000]$ in these calculations, respectively. The relative discrepancy between Fisher matrix elements is less than 10^{-3} in all cases.

E.3. *Eccentric pericenter-precessing binary waveforms*

E.3.1. *Signal-to-noise ratio*

We test the numerical accuracy of the precessing waveform using the following theorem for leading order post-Newtonian binary inspirals. The amount of energy radiated in GWs is equal to the loss of mechanical energy of the binary, which is the same for the *NoPrec* and for the *Prec* models. This is due to the fact that a determines the mechanical energy of the binary in the Newtonian approximation, and the orbital elements a and e are not affected by pericenter precession. This also implies that the amount of SNR must be equal for the *NoPrec* and our precessing models for white noise.

Thus, assuming white noise, we generate a set of source parameters, and compare the output of $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ with the output of $\mathrm{SNR}_{\mathrm{Prec}}^{\mathrm{Gen}}$ for several e_0 and ρ_{p0}, where e_0 and ρ_{p0} are drawn from a uniform distribution between $]0, 1[$ and $[5, 1000]$ in these calculations, respectively. The relative discrepancy between $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{Gen}}$ and $\mathrm{SNR}_{\mathrm{Prec}}^{\mathrm{Gen}}$ is less than 10^{-3} in all cases.

Furthermore, we repeat the same analysis for $\mathrm{SNR}_{\mathrm{NoPrec}}^{\mathrm{High}}$ and $\mathrm{SNR}_{\mathrm{Prec}}^{\mathrm{High}}$, where e_0 is drawn from a uniform distribution between $]0.9, 1[$. We find that the discrepancy between SNR values is less than 10^{-3} in all cases.

Finally, we generate a set of source parameters for comparison, and compare the output of $\mathrm{SNR}_{\mathrm{Prec}}^{\mathrm{Gen}}$ with the output of $\mathrm{SNR}_{\mathrm{Prec}}^{\mathrm{High}}$ for several high e_0 and ρ_{p0} and assuming white noise. Similarly to the *NoPrec* model, e_0 and ρ_{p0} are drawn from a uniform distribution between $]0.9, 1[$ and $[5, 1000]$ in these calculations, respectively. We find that the relative discrepancy between $\mathrm{SNR}_{\mathrm{Prec}}^{\mathrm{Gen}}$ and $\mathrm{SNR}_{\mathrm{Prec}}^{\mathrm{High}}$ is less than 10^{-3} in all cases.

E.3.2. *Fisher matrix*

$\Gamma_{jk}^{\mathrm{Prec,High}}$ can be validated by following the procedure for $\Gamma_{jk}^{\mathrm{Prec,Gen}}$, but considering $\Gamma_{jk}^{\mathrm{NoPrec,High}}$ and high e_0 values ($e_0 \geq 0.9$). We find that the relative discrepancy between Fisher matrix elements are less than 10^{-2} in all cases.

Finally, after generating a set of source parameters for comparison, we compare the output of $\Gamma_{jk}^{\mathrm{Prec,High}}$ and $\Gamma_{jk}^{\mathrm{Prec,Gen}}$ for the Fisher matrix elements corresponding to the parameter set λ_{Prec} in the high eccentricity limit for several ρ_{p0} in the range $[5, 1000]$ and for a single aLIGO detector. The discrepancy between Fisher matrix elements is less than 10^{-3} in all cases.

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